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K2-291b: A Rocky Super-Earth in a 2.2 day Orbit\* \* Based on observations obtained at the W. M. Keck Observatory, which is operated jointly by the University of California and the California Institute of Technology. Keck time has been granted by NASA,...

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K2-291b: A ROCKY SUPER-EARTH IN A 2.2 DAY ORBIT.\*

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Based on observations made with the Italian *Telescopio Nazionale Galileo* (TNG) operated by the *Fundación Galileo Galilei* (FGG) of the *Istituto Nazionale di Astrofisica* (INAF) at the *Observatorio del Roque de los Muchachos* (La Palma, Canary Islands, Spain)

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### ABSTRACT

K2-291 (EPIC 247418783) is a solar-type star with a radius of  $R_* = 0.899 \pm 0.034 R_\odot$  and mass of  $M_* = 0.934 \pm 0.038 M_\odot$ . From the *K2* C13 data, we found one super-Earth planet ( $R_p = 1.589^{+0.095}_{-0.072} R_\oplus$ ) transiting this star on a short period orbit ( $P = 2.225177^{+6.6e-5}_{-6.8e-5}$  days). We followed this system up with adaptive-optic imaging and spectroscopy to derive stellar parameters, search for stellar companions, and determine a planet mass. From our 75 radial velocity measurements using HIRES on Keck I and HARPS-N on Telescopio Nazionale Galileo, we constrained the mass of K2-291 b to  $M_p = 6.49 \pm 1.16 M_\oplus$ . We found it necessary to model correlated stellar activity radial velocity signals with a Gaussian process (GP) in order to more accurately model the effect of stellar noise on our data; the addition of the GP also improved the precision of this mass measurement. With a bulk density of  $\rho = 8.84^{+2.50}_{-2.03} \text{ g cm}^{-3}$ , the planet is consistent with an Earth-like rock/iron composition and no substantial gaseous envelope. Such an envelope, if it existed in the past, was likely eroded away by photoevaporation during the first billion years of the star's lifetime.

*Keywords:* techniques: radial velocities, techniques: photometric, planets and satellites: composition,

## 1. INTRODUCTION

NASA’s *Kepler* and *K2* missions have found hundreds of small, transiting planets with orbital periods less than 10 days. Planets with such short orbital periods are not represented among the solar system planets. In this paper, we describe the discovery and characterization of one such super-Earth sized planet, K2-291 b, orbiting close to its host star ( $P = 2.225177^{+6.6e-5}_{-6.8e-5}$  days).

With a radius of  $R_p = 1.589^{+0.095}_{-0.072} R_{\oplus}$ , K2-291 b lies between two peaks in planet occurrence (Fulton et al. 2017). This bimodality in radius space potentially corresponds to a divide in planet composition (Marcy et al. 2014; Weiss & Marcy 2014; Lopez & Fortney 2014; Rogers 2015). By determining the mass of K2-291 b, we explore this potential boundary between super-Earth and sub-Neptune planets.

Furthermore, one way that sub-Neptunes can transition across this divide to become rocky super-Earths is through photoevaporation, a process where high-energy photons from the star heat and ionize the envelope causing significant portions to escape. Low-mass planets receiving high stellar fluxes will lose a larger portion of their envelopes (Owen & Wu 2013; Lopez & Fortney 2013). This paper explores the potential occurrence of such a process for K2-291 b.

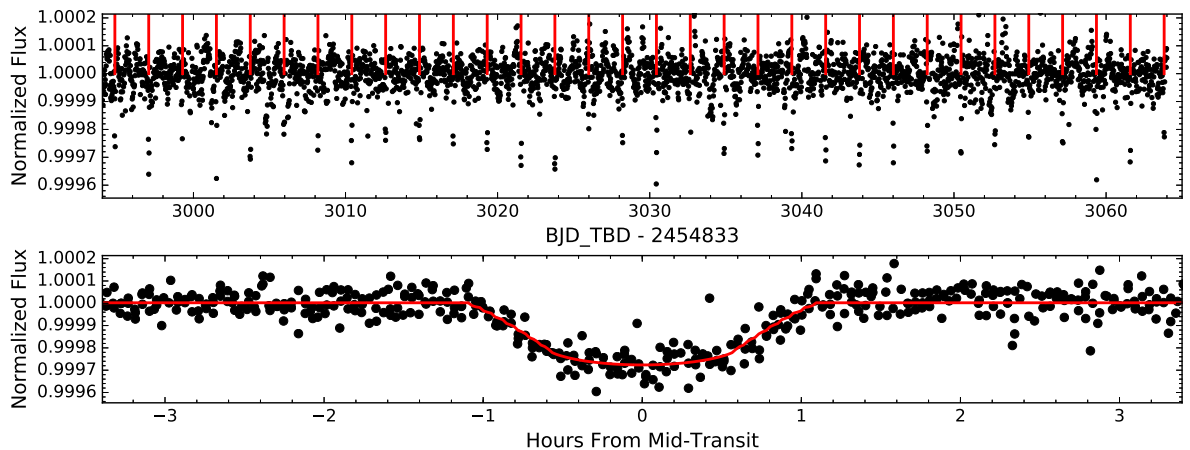
In Section 2 we describe the transit discovery and characterization from the *K2* data. Next, we describe our stellar characterization using both spectra and adaptive optics (AO) imaging in Section 3. Our follow-up radial velocity observations are described and analyzed

in Section 4. We discuss implications of the bulk density of K2-291 b and potential planet evolution through photoevaporation in Section 5. Finally, we conclude in Section 6.

## 2. *K2* LIGHT CURVE ANALYSIS

Photometry of K2-291 was collected during Campaign 13 of NASA’s *K2* mission between 2017 March 8 and 2017 May 27. We processed the *K2* data using a photometric pipeline that has been described in detail in past works by members of our team (Petigura et al. 2018, and references therein). In short, we used the package `k2phot` to analyze the *K2* light curves (Petigura et al. 2015; Aigrain et al. 2016), perform photometry on the *K2* target pixel files, model the time and position dependent photometric variability, and choose the aperture that minimizes noise on three-hour timescales.

We find the signal of one transiting planet at a period of  $P = 2.225177^{+6.6e-5}_{-6.8e-5}$  days (Figure 1, Table 1) in the light curve with the publicly available TERRA algorithm (Petigura et al. 2018). In short, TERRA flags targets with potential transit signals as threshold-crossing events (TCEs); once a TCE is flagged, TERRA masks the previous TCE and is run again on the target star to search for additional signals in the same system. For K2-291, TERRA finds one TCE with a signal-to-noise (S/N) ratio of 21; this signal is consistent with a super-Earth-sized planet transit. After determining the parameters of the host star, described below in Section 3.2, we perform a full Markov chain Monte Carlo (MCMC) analysis on the light curve using a custom Python wrapper of the `batman`<sup>1</sup> transit fitting code (Kreidberg 2015).



**Figure 1.** Top: normalized flux of *K2* light curve with red tick-marks indicating the transit times. Bottom: phase-folded transit data (black points) including our model fit (red line).

<sup>1</sup> Available at <https://github.com/lkreidberg/batman>

**Table 1.** Transit Derived Parameters

Parameter	Name (units)	Value
$T_0$	Time of transit (BJD <sub>TDB</sub> )	$2457830.06163^{+0.00099}_{-0.00104}$
$P$	Period (days)	$2.225177^{+0.000066}_{-0.000068}$
$i$	Inclination (degrees)	$85.26^{+0.23}_{-0.20}$
$R_P/R_*$	Radius of planet in stellar radii (%)	$1.614^{+0.062}_{-0.033}$
$T_{14}$	Total duration (hr)	$1.719^{+0.041}_{-0.032}$
$T_{23}$	Second-to-third contact Transit duration (hr)	$1.625^{+0.043}_{-0.035}$
$R_*/a$	Semimajor axis in stellar radii	$0.1283^{+0.0017}_{-0.0016}$
$b$	Impact parameter	$0.646^{+0.021}_{-0.026}$
$a$	Semimajor axis (AU)	$0.03261^{+0.00044}_{-0.00044}$
$R_P$	Radius ( $R_\oplus$ )	$1.589^{+0.095}_{-0.072}$
$S_{inc}$	Incident stellar flux ( $S_\oplus$ )	$633^{+59}_{-56}$

Our general approach is described further in our previous papers (e.g. [Crossfield et al. 2016](#)). In short, we initialize our `batman` fit with the best-fit parameters from `TERRA` to perform a maximum-likelihood fit and use `emcee`<sup>2</sup> ([Foreman-Mackey et al. 2013](#)) to determine errors. Our model parameters are the time of transit  $T_0$ , orbital period  $P$ , inclination  $i$ , radius of planet in stellar radii ( $R_p/R_*$ ), transit duration  $T_{14}$ , second-to-third contact duration  $T_{23}$ , semimajor axis in stellar radii  $R_*/a$ , impact parameter  $b$ , and quadratic limb-darkening coefficients  $u_1$  and  $u_2$ . [Figure 1](#) shows our best-fit transit model and [Table 1](#) lists the parameters and uncertainties.

### 3. STELLAR CHARACTERIZATION

#### 3.1. Collection of Spectra

We collected 75 radial velocity measurements of K2-291 ([Table 2](#)) with the High Resolution Echelle Spectrometer (HIRES; [Vogt et al. 1994](#)) on the Keck I Telescope on Maunakea and the High Accuracy Radial velocity Planet Searcher in the Northern hemisphere (HARPS-N; [Cosentino et al. 2012](#)) on the Telescopio Nazionale Galileo in La Palma ([Table 2](#)). HARPS-N is an updated version of HARPS at the European Southern Observatory (ESO) 3.6-m ([Mayor et al. 2003](#)).

<sup>2</sup> Available at <https://github.com/dfm/emcee>

We obtained 50 measurements with HIRES between 2017 August and 2018 February. These data were collected with the C2 decker with a typical S/N of 150/pixel (125k on the exposure meter,  $\sim 10$  minute exposures). An iodine cell was used for wavelength calibration ([Butler et al. 1996](#)). We also collected a higher resolution template observation with the B3 decker on 2017 September 6 with 0.8" seeing. The template was a triple exposure with a total S/N of 346/pixel (250k each on the exposure meter) without the iodine cell. See [Howard et al. \(2010\)](#) for more details on this data collection method.

We obtained 25 measurements with HARPS-N between 2017 November and 2018 March as part of the HARPS-N Collaboration's Guaranteed Time Observations (GTO) program. The observations follow a standard observing approach of one or two observations per GTO night, separated by 2–3 hours. The spectra have signal-to-noise ratios in the range  $S/N = 35 - 99$  (average  $S/N = 66$ ), seeing and sky transparency dependent, at 550 nm in 30 minute exposures. This separation was designed to well sample the planet's orbital period and to minimize the stellar granulation signal ([Dumusque et al. 2011](#)).

The HIRES data reduction and analysis followed the California Planet Search method described in [Howard et al. \(2010\)](#). The HARPS-N spectra were reduced with version 3.7 of the HARPS-N Data Reduction Software, which includes corrections for color systematics introduced by variations in seeing ([Cosentino et al. 2014](#)). The HARPS-N radial velocities were computed with a numerical weighted mask following the methodology outlined by [Baranne et al. \(1996\)](#) and [Pepe et al. \(2002\)](#). The resultant radial velocities are presented in [Table 2](#) and in [Figure 6](#).

The HIRES data were collected with three consecutive exposures of 10 minutes each to well sample the stellar p-mode (acoustic) oscillations which occur on a timescale of a few minutes. The HARPS-N data were collected in single observations. Multiple exposures per night were frequently taken, separated by a few hours, to better sample the planet orbital period.

#### 3.2. Stellar Parameters

We derived the stellar parameters by combining constraints from spectroscopy, astrometry, and photometry ([Table 4](#)). The methodology is described in detail in [Fulton & Petigura \(2018\)](#) and summarized in the following paragraphs. We used the HIRES template spectrum to determine the parameters described below. A comparison analysis performed on the HARPS-N data resulted in  $3\sigma$  consistent parameters.

**Table 2.** Radial Velocities

Time ( $BJD_{TDB}$ )	RV <sup>a</sup> ( $\text{m s}^{-1}$ )	RV Unc. ( $\text{m s}^{-1}$ )	S <sub>HK</sub>	Instrument
2457984.09683	-14.53	1.10	0.2227	HIRES
2457985.06918	-7.19	1.33	0.2231	HIRES
2457985.07415	-3.85	1.45	0.2238	HIRES
2457985.07875	-6.89	1.37	0.2247	HIRES
2457994.11807	-7.74	1.28	0.2417	HIRES
2457994.12222	-10.25	1.27	0.2413	HIRES
2457994.12637	-7.25	1.32	0.243	HIRES
2457995.12506	-11.41	1.29	0.2359	HIRES
2457995.12929	-16.88	1.42	0.236	HIRES
2458000.11563	-9.38	1.37	0.2237	HIRES
2458001.12702	-9.92	1.3	0.2134	HIRES
2458001.13405	-13.23	1.33	0.2165	HIRES
2458003.11375	-5.1	1.33	0.2322	HIRES
2458003.11762	-1.76	1.34	0.2337	HIRES
2458003.12159	1.11	1.3	0.2347	HIRES
2458029.07456	-8.94	1.34	0.2629	HIRES
2458030.00982	-3.89	1.38	0.2583	HIRES
2458030.01466	-3.09	1.3	0.2579	HIRES
2458030.01926	-1.76	1.47	0.2595	HIRES
2458096.90078	0.95	1.41	0.2502	HIRES
2458096.90588	0.68	1.43	0.2493	HIRES
2458096.91035	3.57	1.55	0.2526	HIRES
2458097.86564	-3.54	1.34	0.2431	HIRES
2458097.87041	0.29	1.42	0.2414	HIRES
2458097.87537	1.91	1.45	0.2415	HIRES
2458098.89427	8.87	1.53	0.2367	HIRES
2458098.90096	9.69	1.48	0.2363	HIRES
2458098.90727	11.09	1.56	0.2407	HIRES
2458099.86349	1.95	1.37	0.244	HIRES
2458099.86835	5.28	1.38	0.245	HIRES
2458099.87328	4.51	1.44	0.2455	HIRES
2458111.81267	14.64	1.27	0.2229	HIRES
2458111.82241	14.9	1.23	0.2216	HIRES
2458112.83397	6.53	1.4	0.227	HIRES
2458112.83884	4.42	1.43	0.2268	HIRES
2458112.84365	6.73	1.43	0.2281	HIRES
2458113.82544	-0.66	1.09	0.2283	HIRES
2458113.83397	-2.47	1.22	0.2284	HIRES
2458113.84270	1.27	1.18	0.2283	HIRES

**Table 3.** Radial Velocities (continued)

Time ( $BJD_{TDB}$ )	RV <sup>a</sup> ( $\text{m s}^{-1}$ )	RV Unc. ( $\text{m s}^{-1}$ )	S <sub>HK</sub>	Instrument
2458116.77169	7.11	1.56	0.2331	HIRES
2458116.77956	5.75	1.46	0.2322	HIRES
2458116.78838	0.38	1.62	0.2308	HIRES
2458124.91011	-10.25	1.54	0.2131	HIRES
2458149.82066	8.42	1.43	0.2434	HIRES
2458149.82683	1.17	1.47	0.2412	HIRES
2458149.83294	6.7	1.44	0.238	HIRES
2458150.80509	-1.18	1.47	0.2428	HIRES
2458150.81242	-3.06	1.44	0.2394	HIRES
2458150.81994	-0.47	1.54	0.2403	HIRES
2458154.93207	-8.32	1.55	0.2267	HIRES
2458086.52993	25109.80	1.88	0.2471	HARPS-N
2458098.47831	25132.18	0.98	0.2768	HARPS-N
2458102.52715	25125.97	1.61	0.2586	HARPS-N
2458102.66850	25124.74	1.18	0.2663	HARPS-N
2458103.53734	25125.87	1.16	0.2546	HARPS-N
2458111.60905	25135.58	1.8	0.2614	HARPS-N
2458111.68070	25138.34	1.72	0.2659	HARPS-N
2458112.48304	25136.57	1.12	0.2654	HARPS-N
2458119.51674	25125.55	1.87	0.256	HARPS-N
2458120.53530	25128.33	2.08	0.2594	HARPS-N
2458121.58388	25126.18	1.96	0.2507	HARPS-N
2458122.54352	25127.78	2.15	0.2581	HARPS-N
2458143.41096	25128.72	4.73	0.2605	HARPS-N
2458143.50100	25122.77	3.18	0.2479	HARPS-N
2458144.42228	25122.56	1.14	0.2492	HARPS-N
2458144.52492	25120.91	1.94	0.2525	HARPS-N
2458145.42207	25128	1.04	0.2516	HARPS-N
2458145.53009	25123.64	1.41	0.2299	HARPS-N
2458147.53113	25124	1.39	0.2469	HARPS-N
2458172.44662	25122.83	1.18	0.2689	HARPS-N
2458174.35959	25120.61	1.3	0.2545	HARPS-N
2458184.41947	25131.24	1.2	0.2737	HARPS-N
2458187.45055	25150.7	2.91	0.2772	HARPS-N
2458188.44948	25134.35	1.58	0.2859	HARPS-N
2458189.42273	25127.76	1.63	0.2702	HARPS-N

NOTE—<sup>a</sup> HIRES observations report radial velocity changes with respect to the systematic velocity of an observed spectrum whereas HARPS-N observations use a delta-function template with true rest wavelengths.

**Table 4.** Stellar Parameters

Parameter	Name (units)	Value
<b>Name &amp; Magnitude<sup>a</sup></b>		
K2		291
EPIC		247418783
UCAC ID		558-013367
2MASS ID		05054699+2132552
Gaia DR2		3409148746676599168
HD		285181
$K_p$	mag	9.89
$R$	mag	$9.84 \pm 0.14$
$J$	mag	$8.765 \pm 0.032$
$K$	mag	$8.35 \pm 0.02$
$V$	mag	$10.01 \pm 0.03$
<b>Location<sup>b</sup></b>		
RA	Right ascension (deg)	05 05 46.991
DEC	Declination (deg)	+21 32 55.021
$\pi$	Parallax (arcsec)	$0.011076 \pm 6.03\text{e-}05$
$d$	Distance (pc)	$90.23^{+0.51}_{-0.46}$
<b>Stellar Properties</b>		
$A_v$	Extinction (mag)	$0.11740 \pm 0.00061$
$R_*$	Radius ( $R_\odot$ )	$0.899^{+0.035}_{-0.033}$
$M_*$	Mass ( $M_\odot$ )	$0.934 \pm 0.038$
$L_*$	Luminosity ( $L_\odot$ )	$0.682^{+0.014}_{-0.016}$
$T_{\text{eff}}$	Effective temp. (K)	$5520 \pm 60$
$\log(g)$	Surface gravity (cgs)	$4.50 \pm 0.05$
$[Fe/H]$	Metallicity (dex)	$0.08 \pm 0.04$
$v \sin i$	Rotation ( $\text{km s}^{-1}$ )	$< 2.0$
$\log(\text{age})$	Age (yr)	$9.57^{+0.30}_{-0.49}$
$\log(R'_{\text{HK}})$	Chromospheric activity	-4.726

NOTE—<sup>a</sup> MAST, <sup>b</sup> Gaia DR2 (Gaia Collaboration et al. 2018)

Stellar radius is derived from the Stefan Boltzman Law given an absolute bolometric magnitude  $M_{\text{bol}}$  and an effective temperature. We derived stellar effective temperature  $T_{\text{eff}}$ , surface gravity  $\log(g)$ , and metallicity  $[Fe/H]$  by fitting our iodine-free template spectrum

using the Spectroscopy Made Easy<sup>3</sup> (SME) spectral synthesis code (Valenti & Piskunov 2012) following the prescriptions of Brewer et al. (2016). Stellar mass is then calculated using the package isoclassify<sup>4</sup> (Huber et al. 2017). We then derived bolometric magnitudes according to

$$M_{\text{bol}} = m_K - A_k - \mu + BC, \quad (1)$$

where  $m_K$  is the apparent  $K$ -band magnitude,  $A_k$  is the line-of-sight  $K$ -band extinction,  $\mu$  is the distance modulus, and  $BC$  is the  $K$ -band bolometric correction. In our modeling, constraints on  $m_K$  come from 2MASS (Skrutskie et al. 2006) and constraints on  $\mu$  come from the Gaia DR2 parallax measurement (Gaia Collaboration et al. 2018). We derived  $BC$  by interpolating along a grid of  $T_{\text{eff}}$ ,  $\log g$ ,  $[Fe/H]$ , and  $A_V$  in the MIST/C3K grid<sup>5</sup> (C. Conroy et al., in prep.; Dotter 2016; Choi et al. 2016). To find  $A_k$ , we first estimate  $A_v$  from a 3D interstellar dust reddening map by Green et al. (2018), then convert to  $A_k$  using the extinction vector from Schlafly et al. (2018).

The stellar rotation velocity  $v \sin i$ , is computed using the SpecMatch-Syn code (Petigura 2015). Due to the resolution of the instrument the code has been calibrated down to  $2 \text{ km s}^{-1}$ ; values smaller should be considered as an upper limit. Although we measured a value of  $0.2 \text{ km s}^{-1}$ , we adopt  $v \sin i < 2 \text{ km s}^{-1}$ . To determine the chromospheric activity measurement  $\log(R'_{\text{HK}})$ , we measured the flux in the calcium H and K lines relative to the continuum as described in Isaacson & Fischer (2010). Small differences are noted as  $S_{\text{HK}}$  and are tracked to determine if the stellar activity is influencing the radial velocity data.

### 3.3. Search for Stellar Companions

We searched for stellar companions and blended background stars to K2-291 since these stars could contaminate the stellar flux in the  $K2$  aperture, resulting in an inaccurate planet radius and affecting our radial velocity data if bound.

We searched for secondary spectral lines with the ReaMatch algorithm (Kolbl et al. 2015). This algorithm searches for faint orbiting companion stars or background stars that are contaminating the spectrum of the target star. There are no companions detected down to 1% of the brightness of K2-291 with a radial velocity offset of less than  $10 \text{ km s}^{-1}$ .

<sup>3</sup> Available at <http://www.stsci.edu/~valenti/sme.html>

<sup>4</sup> Available at <https://github.com/danxhuber/isoclassify>

<sup>5</sup> Available at [http://waps.cfa.harvard.edu/MIST/model\\_grids.html](http://waps.cfa.harvard.edu/MIST/model_grids.html)

We further looked for stellar companions to K2-291 with adaptive optics (AO). We observed K2-291 on 2017 August 3 UT with NIRC2 on the Keck II AO system (Wizinowich et al. 2000). We obtained images with a 3-point dither pattern in the Br- $\gamma$  and J<sub>cont</sub> filters at an airmass of 1.71. We do not detect any companions down to  $\Delta\text{Br-}\gamma = 6.41$  at 1.03'' as shown in Figure subsection 3.3.

Complementary follow-up observations were taken on 2017 September 7 UT with PHARO-AO on the Hale telescope (Hayward et al. 2001). We obtained images with a five-point dither pattern in the Br- $\gamma$  filter at an airmass of 1.04. The conditions of our observations allowed us to be sensitive down to  $\Delta\text{Br-}\gamma = 8.05$  at 1.05'' as shown in Figure subsection 3.3 and confirm that we detect no companions to EPIC 247418783 above our limits; this also suggests that the transit signal detected is not by a background eclipsing binary.

### 3.4. Stellar Activity Analysis

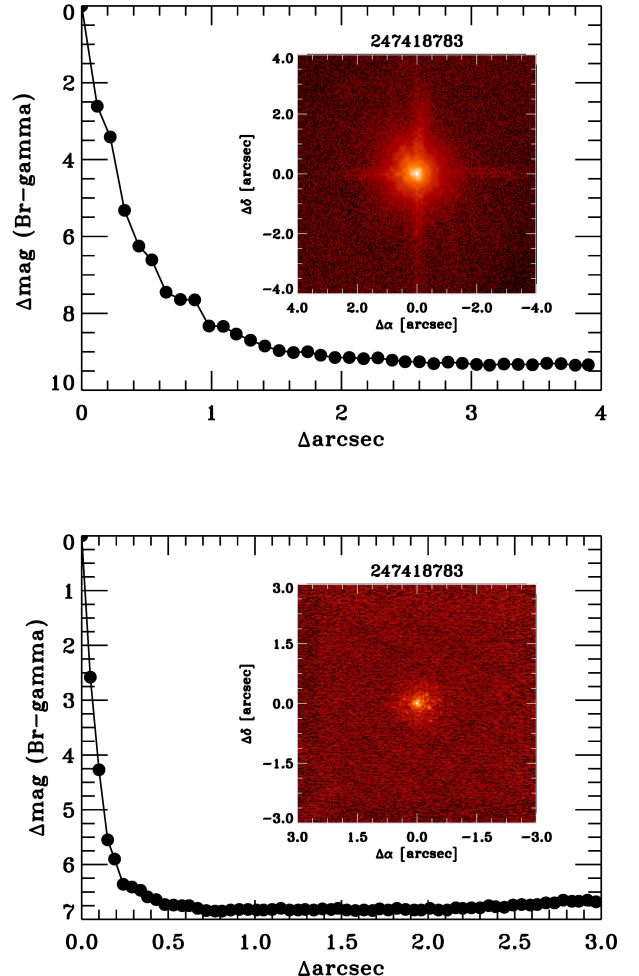
Stars produce intrinsic radial velocity variations due to their internal and surface processes that can be mistaken as planetary signals. The timescales of these radial velocity variations range from a few minutes or hours (p-modes and granulation) to days or years (stellar rotation and large-scale magnetic cycle variations) (Schrijver & Zwaan 2000).

We examine the *K2* light curve periodicity (Figure 3) with a Lomb-Scargle periodogram from `scipy` (Jones et al. 2001) and attribute the clear signal at 18.1 days to rotational modulation of stellar surface features (e.g. spots). There is a secondary peak at half of the strongest peak, and no other significant peaks.

One must consider these timescales when planning radial velocity data collection and analysis to adequately average out or monitor these signals (Dumusque et al. 2011). As described in Section 3.1, we chose the exposure time, spacing, and number of exposures to reduce the effects of p-modes and granulation. We investigated the potential radial velocity signal from the stellar rotation by examining the Calcium II H and K lines ( $S_{\text{HK}}$ , Table 2) in the HIRES and HARPS-N data (Isaacson & Fischer 2010).

We found a clear signal in both the  $S_{\text{HK}}$  and radial velocity data that matches the timescale of the rotation period of K2-291 (Figure 4), as determined from the *K2* light curve; therefore we need to account for this signal in our radial velocity analysis.

We then estimated the correlation coefficient between the measured radial velocity and activity indexes. Due to different zero-points in both radial velocity and  $S_{\text{HK}}$ , we performed the analysis for the two instruments,

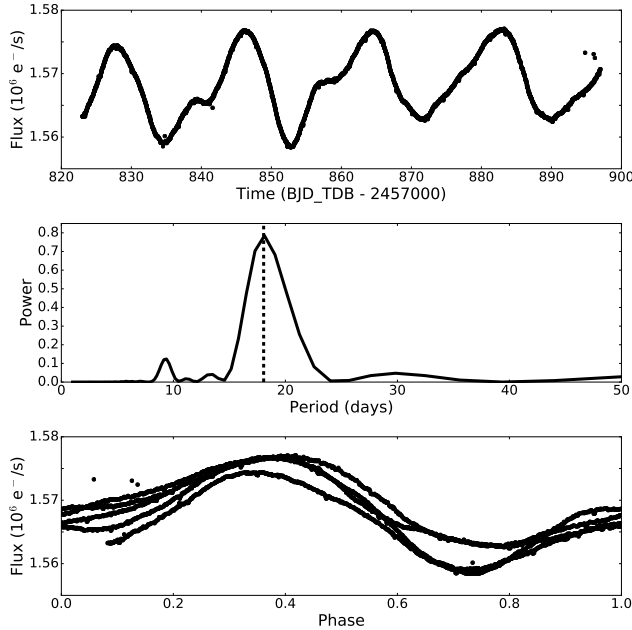


**Figure 2.** We detect no objects near K2-291 with PHARO-AO on the Hale telescope (top) or with Keck/NIRC2 adaptive optics (bottom), as shown in the inset images and the resultant Br- $\gamma$  contrast curves. The curves plotted correspond to a five- $\sigma$  detection limit.

HARPS-N and HIRES, independently. From the calculation of the correlation coefficient value and the knowledge of the sample size,  $p$ -value analysis is often used to reject the null hypothesis of non-correlation at a given significance level. We calculated the  $p$ -value for both datasets using `scipy.stats.pearsonr` (Jones et al. 2001). The HARPS-N radial velocity and  $S_{\text{HK}}$  data have a  $p$ -value of 0.01 allowing us to reject the null hypothesis, therefore suggesting a correlation. The HIRES data, however, have a  $p$ -value of 0.45 which does not support a correlation.

To check any potential flaws in the  $p$ -value test we also used the Bayesian framework described in Figueira et al. (2016) that allows us to estimate the probability distri-

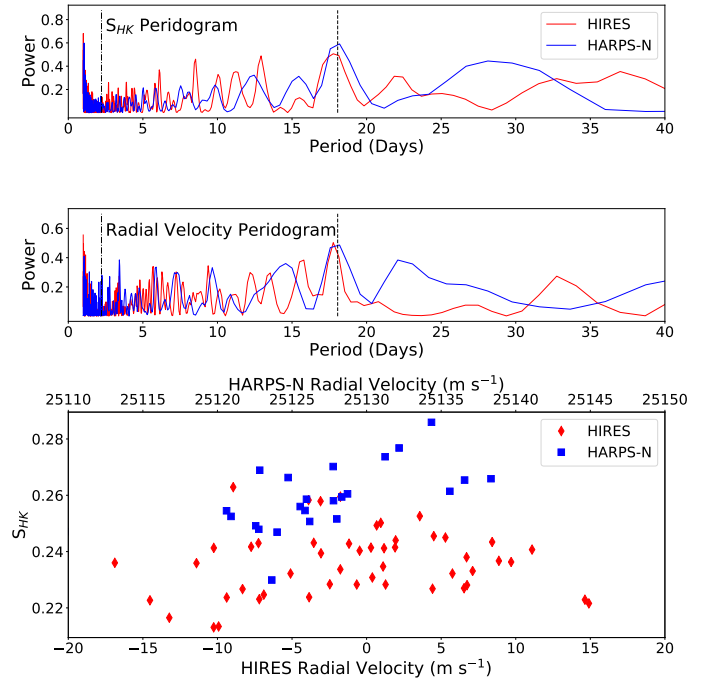




**Figure 3.** Top: light curve of K2-291 from *K2* C13. We attribute the periodicity to stellar rotation and the variation to star spot modulation. Transits are too shallow to be seen by eye, and are shown in Figure 1. Middle: Lomb-Scargle periodogram of *K2* data, illustrating clear periodicity at 18.1 days (dotted line). Bottom: *K2* data phase-folded over 18.1 days.

bution of the coefficient, providing important insight on the correlation presence. This framework calculates the Pearson’s correlation coefficient to test for the presence of a linear correlation, and the Spearman’s rank to test for the presence of a monotonic correlation.

On HARPS-N data we obtain a Pearson’s correlation coefficient of 0.56 with a 95% highest probability density (HPD) between the values [0.29, 0.79], and a Spearman’s rank of 0.63 with 95% HPD of [0.39, 0.83]. This shows that not only the correlation coefficient is large but that its distribution populates essentially positive correlation values. As such, the correlation is strong and significant, both in linear and monotonic terms. On the other hand, for HIRES we obtain an average value of 0.10 with 95% HPD of [-0.17, 0.35] and 0.13 with 95% HPD of [-0.12, 0.39] for Pearson’s correlation coefficient and Spearman’s rank, respectively. The correlation coefficients are low in absolute value and distributed from negative to positive values; its distribution does not support the presence of a correlation. Different instrument properties, such as wavelength ranges and resolution, may explain the differences in the  $S_{HK}$  values and correlation strengths.



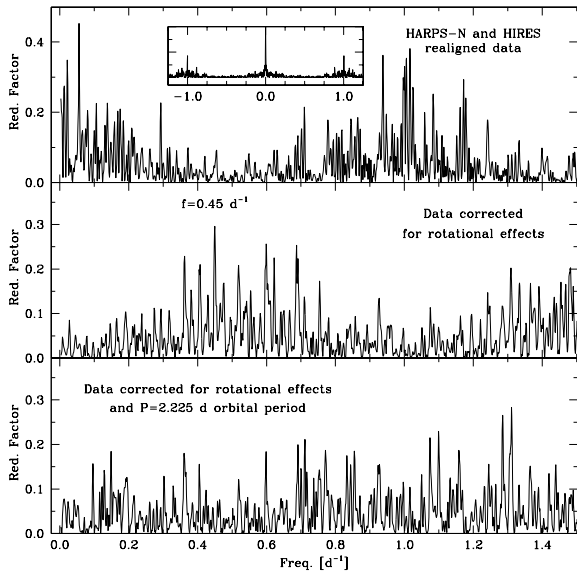
**Figure 4.** Periodograms of  $S_{HK}$  (top), radial velocity (middle), and  $S_{HK}$  vs. radial velocity (bottom). The stellar rotation period (dashed line) and planet orbital period (dashed-dotted line) are shown. There is a strong signal at the stellar rotation period in both datasets.

## 4. RADIAL VELOCITY ANALYSIS

### 4.1. Radial Velocity Planet Search

We first searched for K2-291 b in the combined HIRES and HARPS-N datasets without any priors from our transit analysis to provide an independent planet detection. The radial velocity datasets from HIRES and HARPS-N are merged using the  $\gamma$  values reported in Table 5 to adjust for their different zero-points in this search. The 75 datapoints thus obtained are then analyzed in frequency (Figure 5) using the Iterative Sine-Wave fitting (Vaníček 1971), by computing the fractional reduction in the residual variance after each step (reduction factor). This is an iterative process; peaks should be directly compared within an iteration but not between them. The power spectrum immediately supplies the rotational period at  $f=0.055 \text{ day}^{-1}$  (top panel), corresponding to  $P_{\text{rot}}=18.1$  days. The light curve is very asymmetrical (Figure 3) and therefore signals are visible at the harmonics values,  $f, 2f, 3f$ , and  $4f$ . We were successful in detecting the expected frequency of the planet signal at  $f=0.45 \text{ d}^{-1}$  after including the stellar rotational frequencies in a simultaneous fit (middle panel). We also searched for any other additional signals, but we did not detect any clear peaks (bottom panel). Indeed,

the interaction of the noise with the spectral window (insert in the top panel) prevents any reliable further identification.



**Figure 5.** Top panel: power spectrum of the radial velocity data of EPIC 247418783; the spectral window is shown in the insert. Middle panel: power spectrum obtained considering a long-term trend and  $f=0.0098 \text{ d}^{-1}$ ,  $2f$ ,  $3f$ ,  $4f$  (but not its amplitude and phase) as known constituents. The planet signal is seen at  $f=0.45 \text{ d}^{-1}$ . Bottom panel: power spectrum obtained considering  $f$ ,  $2f$ ,  $3f$ ,  $4f$  and the planet orbital frequency (but not their amplitudes and phases) as known constituents. No clear peaks suggesting additional signals are detected.

#### 4.2. Radial Velocity Fit with RadVel

After this initial, transit-blind radial velocity analysis, we analyzed the radial velocity data using RadVel<sup>6</sup> (Fulton et al. 2018). RadVel is an open source Python package that models Keplerian orbits to fit radial velocity data by first performing a maximum-likelihood fit to the data and then determining errors through an MCMC analysis. We use the default number of walkers, number of steps, and criteria for burn-in and convergence as described in Fulton et al. (2018).

A single planet at an orbital period of  $P=2.225177^{+6.6e-5}_{-6.8e-5}$  days was found in the *K2* photometry (Section 2); we

include a Gaussian prior on the orbital period  $P$  and time of transit  $T_{\text{conj}}$  from the *K2* data (Table 1). We first modeled this system using a one-planet fit including a constant offset for each dataset  $\gamma$ . This fit results in a semi-amplitude for the 2.2 day planetary signal of  $K_p=3.1\pm 1.7 \text{ m s}^{-1}$ .

Next, we tested models including an additional trend ( $\dot{\gamma}$ ), curvature ( $\ddot{\gamma}$ ), and eccentricity ( $e$ ,  $\omega$ ). We used the Bayesian Information Criteria (BIC) to evaluate if the fit improved sufficiently to justify the additional free parameters; a positive  $\Delta\text{BIC}$  indicates an improved fit. The trend is the only additional parameter which has a noticeable  $\Delta\text{BIC}$  ( $\Delta\text{BIC} = 8.29$ ); the trend is  $\dot{\gamma} = 0.07\pm 0.02 \text{ m s}^{-2}$ . There is nearly no change for the curvature ( $\Delta\text{BIC} = 0.84$ ) or eccentric ( $\Delta\text{BIC} = -1.90$ ) cases. All three additional parameters result in semi-amplitudes within  $1\sigma$  of the circular fit.

#### 4.3. Gaussian Process Inclusion and Training

Stellar activity of K2-291 has an appreciable effect on our measured radial velocities. As discussed in Section 3.4, there is a periodic signal in the radial velocity data that matches both the stellar rotation period determined from the *K2* data and the periodicity in the Calcium H and K lines ( $S_{\text{HK}}$ ). We modeled this stellar signal simultaneously with our planet fit using a Gaussian process (GP) with the default GP model available in RadVel (Blunt et al. in prep). GP regression is a non-parametric statistical technique for modeling correlated noise in data. GP regression enables the determination of physical parameter posterior distributions with uncertainties that reflect the confounding effects of stellar activity noise (e.g. Haywood et al. 2014; Grunblatt et al. 2015; López-Morales et al. 2016).

Stellar noise characteristics in GP models are controlled by a kernel function with one or more hyperparameters, but radial velocity data are often too sparse to confidently determine the values of these hyperparameters (see Faria et al. (2016) for a counterexample). To address this problem, authors in the literature use other data sources to constrain the values of the hyperparameters, then incorporate this information into the radial velocity fit as priors on the hyperparameters (e.g. Haywood et al. 2014; Rajpaul et al. 2015). In this paper, we constrain the values of the hyperparameters in our GP model using *K2* photometry.

We modeled the correlated noise introduced from the stellar activity using a quasi-periodic GP with a covariance kernel of the form

$$k(t, t') = \eta_1^2 \exp \left[ -\frac{(t - t')^2}{\eta_2^2} - \frac{\sin^2\left(\frac{\pi(t - t')}{\eta_3}\right)}{\eta_4^2} \right], \quad (2)$$

<sup>6</sup> Available at <https://github.com/California-Planet-Search/radvel>

where the hyper-parameter  $\eta_1$  is the amplitude of the covariance function,  $\eta_2$  is the active-region evolutionary time scale,  $\eta_3$  is the period of the correlated signal,  $\eta_4$  is the length scale of the periodic component (López-Morales et al. 2016; Haywood et al. 2014).

We explore these hyperparameters for this system by performing a maximum-likelihood fit to the *K2* light curve with the quasi-periodic kernel (Equation 2) then determine the errors through an MCMC analysis. We find  $\gamma_{K2} = 1567969.00^{+1766.12}_{-1830.87}$ ,  $\sigma = 54.60 \pm 9.57$ ,  $\eta_1 = 4429.95^{+897.65}_{-673.95}$ ,  $\eta_2 = 25.18^{+3.50}_{-3.59}$ ,  $\eta_3 = 19.41^{+0.68}_{-1.14}$ , and  $\eta_4 = 0.42^{+0.04}_{-0.03}$ . This stellar rotation period ( $\eta_3$ ) is consistent with the results of our periodogram analysis in Section 3.4.

#### 4.4. Gaussian Process Radial Velocity Fit

We then perform a radial velocity fit including a GP to account for the affects of stellar activity on our measurements. We model our GP as a sum of two quasi-periodic kernels, one for each instrument as HIRES and HARPS-N have different properties, such as wavelength ranges, that could alter the way that stellar activity affects the data. Each kernel includes identical  $\eta_2$ ,  $\eta_3$ , and  $\eta_4$  parameters but allows for different  $\eta_1$  values.

We inform the priors on these hyperparameters from the GP light curve fit (Section 4.3).  $\eta_1$  is left as a free parameter as light curve amplitude cannot be directly translated to radial velocity amplitude.  $\eta_2$  has a Gaussian prior describing the exponential decay of the spot features ( $25.18 \pm 3.59$ ).  $\eta_3$  has a Gaussian prior constrained from the stellar rotation period ( $19.14 \pm 1.14$ ).  $\eta_4$  constrains the number of maxima and minima per rotation period with a Gaussian prior ( $0.42 \pm 0.04$ ), as described in López-Morales et al. (2016). We do not include a prior on the phase of the periodic component of the stellar rotation because spot modulation tends to manifest in radial velocity data with a relative phase shift.

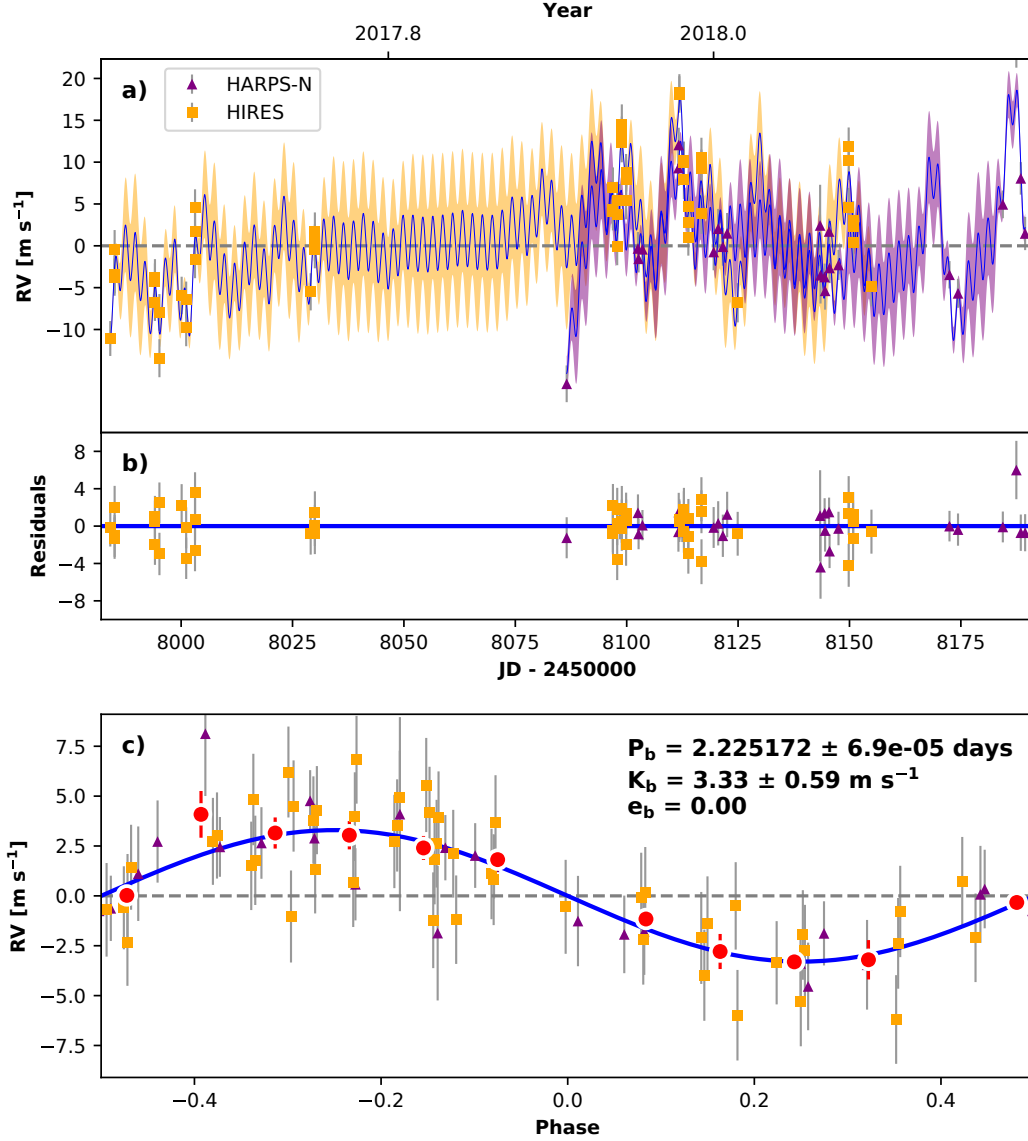
The planet parameters derived from our GP analysis are consistent with our original, non-GP fit within  $1\sigma$ . The uncertainty on the semi-amplitude of the planet signal has decreased by a factor of three to  $K_p = 3.33 \pm 0.59$  m s<sup>-1</sup>. We then investigate the inclusion of additional parameters with our GP fit. All of the tested models increased the BIC value, therefore none of them justified the additional parameters. We adopt the model including the GP with no additional parameters as our best fit, all other models have results within  $1\sigma$ ; our best-fit parameters are listed in Table 5 and the fit is shown in Figure 6.

We choose to include a GP in our analysis to improve the accuracy of our results by including the affects of

**Table 5.** Radial Velocity Fit Parameters

Parameter	Name (Units)	Value
$P_b$	Period (days)	$2.225172^{+6.9e-05}_{-7e-05}$
$T_{\text{conj}_b}$	Time of conjunction (BJD <sub>TDB</sub> )	$2457830.0616^{+0.0011}_{-0.0010}$
$e_b$	Eccentricity	$\equiv 0.0$
$\omega_b$	Argument of periaapse (radians)	$\equiv 0.0$
$K_b$	Semi-amplitude (m s <sup>-1</sup> )	$3.33 \pm 0.59$
$M_b$	Mass ( $M_{\oplus}$ )	$6.49 \pm 1.16$
$\rho_b$	Density (g cm <sup>-3</sup> )	$8.84^{+2.50}_{-2.03}$
$\gamma_{\text{HIRES}}$	Mean center-of-mass velocity (m s <sup>-1</sup> )	$-3.5 \pm 3.2$
$\gamma_{\text{HARPS-N}}$	Mean center-of-mass velocity (m s <sup>-1</sup> )	$25126.2^{+3.4}_{-3.5}$
$\dot{\gamma}$	Linear acceleration (m s <sup>-1</sup> day <sup>-1</sup> )	$\equiv 0.0$
$\ddot{\gamma}$	Quadratic acceleration (m s <sup>-1</sup> day <sup>-2</sup> )	$\equiv 0.0$
$\sigma_{\text{HIRES}}$	Jitter (m s <sup>-1</sup> )	$1.85^{+0.43}_{-0.37}$
$\sigma_{\text{HARPS-N}}$	Jitter (m s <sup>-1</sup> )	$1.43^{+0.85}_{-0.67}$
$\eta_{1,\text{HIRES}}$	Amplitude of covariance (m s <sup>-1</sup> )	$8.45^{+2.21}_{-1.65}$
$\eta_{1,\text{HARPS-N}}$	Amplitude of covariance (m s <sup>-1</sup> )	$8.59^{+2.23}_{-1.77}$
$\eta_2$	Evolution timescale (days)	$26.09^{+3.50}_{-3.62}$
$\eta_3$	Recurrence timescale (days)	$18.66^{+0.95}_{-0.79}$
$\eta_4$	Structure parameter	$0.41 \pm 0.04$

stellar activity. The GP was able to also improve the precision of the mass measurement by a factor of three since the planet orbital period is far from the stellar rotation period, both periods were well sampled with the data, and the stellar activity is dominated by the rotation signal.



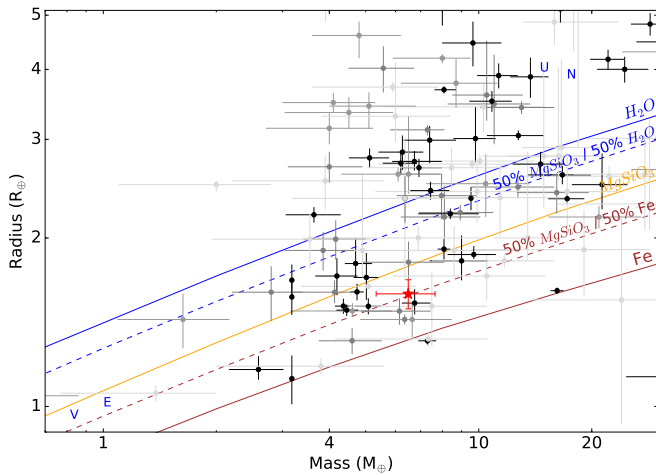
**Figure 6.** Best-fit one-planet Keplerian orbital model for K2-291. The maximum-likelihood model is plotted while the orbital parameters listed in Table 5 are the median values of the posterior distributions. The thin blue line is the best-fit one-planet model with the mean GP model; the colored area surrounding this line includes the  $1\sigma$  maximum-likelihood GP uncertainties. We add the radial velocity jitter term(s) listed in Table 5 in quadrature with the measurement uncertainties for all radial velocities. **b)** Residuals to the best-fit one-planet model and GP model. **c)** Radial velocities phase-folded to the ephemeris of planet b. The small point colors and symbols are the same as in panel a. Red circles are the same velocities binned in 0.08 units of orbital phase. The phase-folded model for planet b is shown as the blue line.

We perform an independent radial velocity analysis using the PyORBIT code <sup>7</sup> (Malavolta et al. 2016, 2018) with results well within  $1\sigma$  with respect to those reported in Table 5.

## 5. DISCUSSION

### 5.1. Mass, Radius, and Bulk Density

Planet compositional models and radial velocity observations of small *Kepler* planets have shown a dividing line between super-Earth and sub-Neptune planets at  $1.5\text{--}2 R_{\oplus}$  (Marcy et al. 2014; Weiss & Marcy 2014; Lopez & Fortney 2014; Rogers 2015; Dressing & Charbonneau 2015). *Kepler* planet radii also display a bimodality in sub-Neptune-sized planets that matches the location of this divide (Fulton et al. 2017). K2-291 b is near the inner edge of the divide ( $1.589^{+0.095}_{-0.072} R_{\oplus}$ ), which makes its composition particularly interesting.



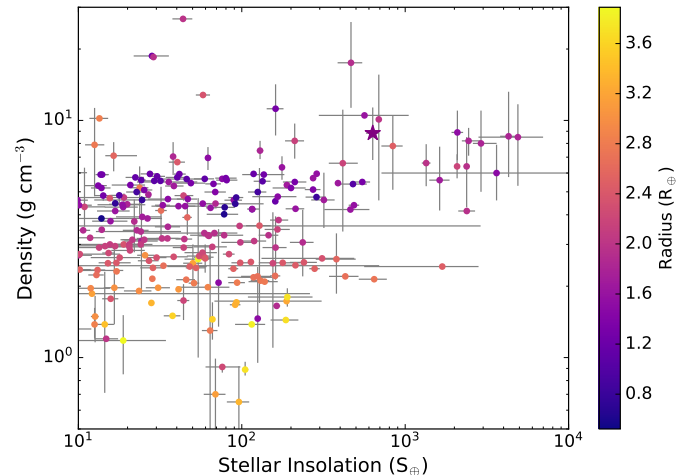
**Figure 7.** Mass-radius diagram for planets between the size of Earth and Neptune with greater than  $2\sigma$  measurements (darker points for lower error). The lines show models of different compositions (Zeng et al. 2016), with solid lines indicating single composition planets and dashed lines for a 50/50 mixture. K2-291 b is shown as a red star along with  $1\sigma$  uncertainties. K2-291 b is consistent with a predominantly rocky composition including an iron core.

As shown in the mass-radius diagram (Figure 7), the composition of K2-291 b is consistent with a silicate planet containing an iron core and lacking substantial volatiles (Zeng et al. 2016). We investigated its composition further using Equation 8 from Fortney et al.

(2007), which assumes a pure silicate and iron composition, to estimate the mass fraction of each. For our mean mass and radius, the mass fraction of silicates is 0.61 and the mass fraction of iron is 0.39, similar to the 0.35 iron core mass fraction of the Earth. For a high gravity case ( $1\sigma$  low radius,  $1\sigma$  high mass), the mass fraction of silicates would be 0.39. For a low gravity case ( $1\sigma$  high radius,  $1\sigma$  low mass), the mass fraction of silicates would be 0.94. In all cases, no volatiles are needed to explain the mass and radius of K2-291 b.

We also estimated the maximum envelope mass fraction of K2-291 b through a model grid from Lopez & Fortney (2014). This grid assumes a solar metallicity envelope with a minimum envelope mass fraction of 0.1%. We generated 100,000 random samples of the envelope fraction from our normal distributions on the mass, radius, age, and flux of K2-291 b. From this, we determined that the  $3\text{-}\sigma$  upper limit on the envelope fraction is 0.3%.

Similarly, *Kepler* planets within 0.15 AU and smaller than  $2 R_{\oplus}$  have an envelope fraction less than 1% (Wolfgang & Lopez 2015). Figure 8 shows the relationship between density and stellar insolation for planets smaller than  $4 R_{\oplus}$ . K2-291 b exhibits a density similar to other small, close-in planets.



**Figure 8.** Density vs. stellar insolation for planets with radii smaller than  $4 R_{\oplus}$ . Point color indicates the radius of the planet. K2-291 b (star) exhibits a density similar to other small, close-in planets.

### 5.2. Photoevaporation

K2-291 b's lack of a substantial volatile envelope could be explained by atmospheric loss. For lower mass planets experiencing a large amount of stellar insolation,

<sup>7</sup> Available at <http://www.github.com/LucaMalavolta/PyORBIT/>

photoevaporation (hydrodynamic escape) is the dominant atmospheric loss process. Photoevaporation occurs when high-energy photons from the host star ionize and heat the atmosphere causing it to expand and escape (Owen & Wu 2013).

K2-291 b is potentially the core of a sub-Neptune planet that underwent photoevaporation. We cannot, however, rule out a scenario where K2-291 b formed with a high density from its onset. In that case, perhaps K2-291 b formed after the gas disk had dissipated, or giant impacts by planetesimals stripped the envelope early in its formation. Although these two scenarios cannot yet be distinguished for an individual planet, population studies can be of use. Swain et al. (2018) finds two separate groups of small planets in radius-insolation-density space. One group is consistent with small solar system bodies and likely has an Earth-like formation, the other forms a bulk density continuum with sub-Neptunes and is likely composed of remnant cores produced by photoevaporation. Another large-scale approach is to look for a radius trend among close orbiting planets of different ages; a trend of smaller young planets compared to larger old planets would suggest photoevaporation. David et al. (2018) finds one such planet and Mann et al. (2017) finds seven close orbiting young planets. There is an emerging trend that these young planets are larger but more planets will need to be found to be statistically significant.

We examine here the possibility that K2-291 b formed by photoevaporation. Due to the hydrodynamic escape of the envelope for close-in planets, the boundary between complete loss and retention of 1% of the envelope is at 0.1 AU for a  $6 M_{\oplus}$  planet orbiting a solar mass star (Owen & Wu 2013). K2-291 b orbits within this boundary at  $a = 0.03261 \pm 0.00044$  AU. For the mass ( $M_p = 6.49 \pm 1.16 M_{\oplus}$ ) and stellar insolation ( $S_{\text{inc}} = 633^{+59}_{-56} S_{\oplus}$ ) of K2-291 b specifically, all of its hydrogen and helium should have been lost between 100 Myr and 1 Gyr, depending on the original hydrogen-helium mass fraction and mass loss efficiency (Lopez & Fortney 2013). We determined an age from the HIRES spectra of  $3.7^{+3.7}_{-2.5}$  Gyr, longer than this photoevaporation timescale.

We ran additional models using the Lopez & Fortney (2014) model grid to calculate the radius K2-291 b would have with an additional hydrogen-helium envelope. Adding 0.1% H/He by mass would result in a planet radius of  $R_p = 1.82 R_{\oplus}$ . Similarly, an additional 1% or 10% would equal a radius of  $R_p = 2.2 R_{\oplus}$  or  $R_p = 3.7 R_{\oplus}$ , respectively. Therefore, a small addition of between 1% and 10% H/He would increase the radius of K2-291 b enough to move the planet across the Fulton gap to the sub-Neptune side.

Together, these analyses imply that K2-291 b may have formed as a sub-Neptune with a substantial volatile envelope and transitioned across the Fulton gap to a super-Earth planet through photoevaporation.

## 6. CONCLUSION

In this paper, we described the discovery and characterization of K2-291 b. From our *K2* analysis (Section 2), we discover K2-291 b, a super-Earth planet with a radius of  $R_p = 1.589^{+0.095}_{-0.072} R_{\oplus}$ . We collected follow-up AO images and spectra to characterize the stellar properties (Section 3). Our radial velocity analysis (Section 4) determined a planet mass of  $M_p = 6.49 \pm 1.16 M_{\oplus}$ .

We accounted for quasi-periodic radial velocity variations induced by the host star’s moderate activity levels using GP regression (S. Blunt et al. in prep, Haywood et al. 2014). This improves the accuracy of our mass determination (e.g. Haywood et al. 2018). In our case, the GP framework also increases the precision of our mass determination over an uncorrelated-noise-only treatment. The increased precision likely results from favorable sampling of the rotational and active-region timescales (López-Morales et al. 2016), combined with the fact that the orbital period is very distinct from these activity timescales.

The density of K2-291 b ( $\rho = 8.84^{+2.50}_{-2.03} \text{ g cm}^{-3}$ ) is consistent with a rock and iron composition. The high density of the planet, along with the high solar flux received by the planet ( $S_{\text{inc}} = 633^{+59}_{-56} S_{\oplus}$ ), indicate that if K2-291 b formed with a substantial envelope, it has been eroded away by photoevaporation.

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