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# Probing Gravitational Lensing of the CMB with SDSS-IV Quasars

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## ABSTRACT

We study the cross-correlation between the *Planck* CMB lensing convergence map and the eBOSS quasar overdensity obtained from the Sloan Digital Sky Survey (SDSS) IV, in the redshift range  $0.9 < z < 2.2$ . We detect the CMB lensing convergence-quasar cross power spectrum at  $5.4\sigma$  significance. The cross power spectrum provides a quasar clustering bias measurement that is expected to be particularly robust against systematic effects. The redshift distribution of the quasar sample has a median redshift  $z \approx 1.55$ , and an effective redshift about 1.51. The best fit bias of the quasar sample is  $b_q = 2.43 \pm 0.45$ , corresponding to a host halo mass of  $\log_{10} \left( \frac{M}{h^{-1} M_\odot} \right) = 12.54^{+0.25}_{-0.36}$ . This is broadly consistent with the previous literature on quasars with a similar redshift range and selection. Since our constraint on the bias comes from the cross-correlation between quasars and CMB lensing, we expect it to be robust to a wide range of possible systematic effects that may contaminate the auto correlation of quasars. We checked for a number of systematic effects from both CMB lensing and quasar overdensity, and found that all systematics are consistent with null within  $2\sigma$ . The data is not sensitive to a possible scale dependence of the bias at present, but we expect that as the number of quasars increases (in future surveys such as DESI), it is likely that strong constraints on the scale dependence of the bias can be obtained.

**Key words:** large-scale structure of Universe – quasars: general – cosmic background radiation

## 1 INTRODUCTION

Cosmic microwave background (CMB) temperature fluctuations provide invaluable information about our Universe, and can give extremely tight constraints on cosmological parameters (Kofman et al. 1993; Hinshaw et al. 2012; Planck Collaboration et al. 2015b, 2018). The primary CMB anisotropy encodes information about the primordial universe, measured at  $z \approx 1100$ . However, since the discovery of the CMB, a lot of progress has also been made on the secondary CMB anisotropies, such as gravitational lensing (Smith et al. 2007; Hirata et al. 2008; Lewis & Challinor 2006), the thermal and kinetic Sunyaev-Zel’dovich (tSZ, kSZ) effects (Sunyaev & Zel’dovich 1980), and the integrated Sachs-Wolfe effect (ISW) (Sachs & Wolfe 1967). These effects can act as foreground for the primary CMB, but they also encode information about the growth of structure at

lower redshifts, a powerful probe of Dark Energy, Modified Gravity and neutrino masses (Lewis & Challinor 2006).

Secondary CMB anisotropies are produced by large scale structures (LSS) in the late-time universe (Aghanim et al. 2008) and can be easily detected through cross-correlation with the LSS. In particular, CMB lensing traces the matter density field at intermediate redshifts ( $z \lesssim 1100$ ). As CMB photons travel to the observer, they are gravitationally deflected by the matter, leaving an imprint on the observed CMB temperature and polarization fluctuations. Weak lensing of the CMB introduces off-diagonal correlations between Fourier modes, allowing the CMB lensing deflection field  $\mathbf{d}$  to be estimated (Hu & Okamoto 2002).

We use the CMB lensing convergence field as a tracer of the dark matter field, and thus the large scale structure of the Universe. More specifically, we use the CMB lensing convergence field to study properties of quasars, which trace the large scale structure at intermediate redshifts. Quasars are thought to be luminous accreting supermassive black holes

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at the centers of distant galaxies (Salpeter 1964). Like galaxies, they are tracers of the 3D distribution of dark matter at different redshifts. With the understanding that almost every galaxy hosts a supermassive black hole at its center (Kormendy & Richstone 1995), quasars can be thought as a phase in the galaxy evolution. Properties of quasars, such as the characteristic mass of their host halos (Tinker et al. 2010; DiPompeo et al. 2014), can be inferred by studying the relationship between the dark matter distribution and quasar clustering. The information about quasar properties can reveal much about the growth of structure over the history of the universe (Marziani & Sulentic 2014; Mortlock 2015).

Both CMB lensing and the observed quasar overdensity depend on the projected matter overdensity, and the quasar redshift distribution matches well with the CMB lensing kernel, so the CMB lensing convergence and the quasar overdensity should have a relatively strong correlation (Peiris & Spergel 2000). Quasars are biased tracers of the underlying matter density field (Kaiser 1984), meaning that the observed cross power spectrum is proportional to the quasar bias. This factor parametrizes the properties of the clustering of quasars and encapsulates the information about the processes of galaxy formation and evolution that are currently not very well understood (Amendola et al. 2017). Measuring this bias factor would be crucial to the understanding of galaxy formation and the evolution of supermassive black holes within the standard structural formation framework (Shen et al. 2009).

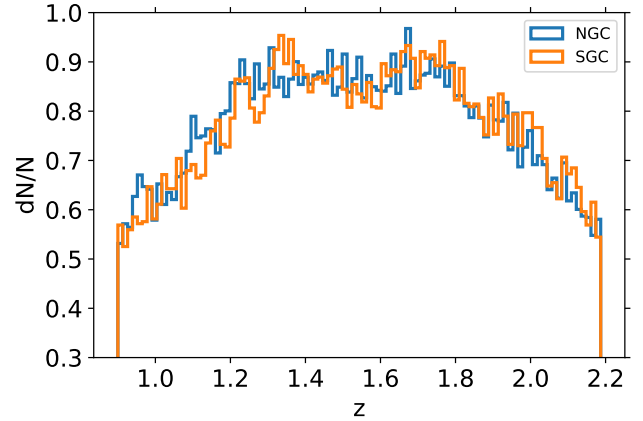
Laurent et al. (2017) have analyzed the auto-correlation of the eBOSS quasars, and put constraints on the quasar bias, as well as the corresponding host halo mass. In this paper, we use an alternative way to constrain the quasar bias, by cross-correlating the CMB lensing map from Planck and quasar overdensities drawn from eBOSS Data Release 14 (Dawson et al. 2016). We measure a quasar bias that is consistent with the auto-correlation result. We then calculate the corresponding characteristic host halo mass of the quasars. All calculations assume a cosmology with the Planck TT+lowP+lensing parameters (Planck Collaboration et al. 2015b).

The remainder of the paper is organized as follows: in Section 2 we present the theoretical background of quasar linear bias and the CMB lensing-quasar angular cross power spectrum. Section 3 includes the data samples and estimators we used to evaluate the observed power spectrum. In Section 4, we estimate the cross power spectrum, the quasar bias, and the characteristic host halo mass. In Section 5, we discuss errors, systematic effects and a null test performed on the data. We draw our conclusions in Section 6.

## 2 THEORETICAL BACKGROUND

### 2.1 Overview

Quasars reside in the nuclei of distant galaxies and are expected to be biased tracers of the matter overdensity on large scales. In other words, the number density of quasars is related to the dark matter overdensity by a bias factor, i.e.  $\delta_q = b_q \delta_m$ , where  $b_q$  can be a function of scale, redshift, formation history or other environment related factors



**Figure 1.** The redshift distributions of the selected quasars in the range  $0.9 < z < 2.2$  in the North and South Galactic Caps (NGC, in blue; and SGC, in orange). The normalized redshift distribution is plotted on the y-axis.

(White & Rees 1978). The amplitude of the deflection by CMB lensing in a given direction depends on the projected matter density in that direction. Thus we expect the quasar number density to be correlated with CMB lensing convergence (Peiris & Spergel 2000).

### 2.2 Angular cross power spectrum

To relate the CMB lensing to the matter overdensity field, we define the lensing convergence,  $\kappa \equiv -\frac{1}{2}\nabla \cdot \mathbf{d}$ , where  $\mathbf{d}$  is the lensing deflection field. The lensing convergence is a weighted projection of the matter overdensity in direction  $\hat{n}$  along the line of sight (Lewis & Challinor 2006):

$$\kappa(\hat{n}) = \int_0^{z_{\text{CMB}}} dz W(z) \delta_m(\chi(z)\hat{n}, z) \quad (1)$$

where  $z_{\text{CMB}} \approx 1100$  is the redshift at the last scattering surface,  $W(z)$  is the CMB lensing kernel,  $\delta_m(\chi(z)\hat{n}, z)$  is the matter overdensity at redshift  $z$  in the direction  $\hat{n}$ , and  $\chi(z)$  is the comoving distance at redshift  $z$ . Assuming a flat universe,  $W(z)$  is given by

$$W(z) = \frac{3H_0^2\Omega_{m,0}}{2cH(z)}(1+z)\chi(z) \left(1 - \frac{\chi(z)}{\chi_{\text{CMB}}}\right) \quad (2)$$

where  $\chi_{\text{CMB}}$  is the comoving distance to the last scattering surface,  $H_0$  is the current Hubble parameter,  $H(z)$  is the Hubble parameter at redshift  $z$ , and  $\Omega_{m,0}$  is the current matter density parameter. Since the lensing potential  $\phi$  is a 2D projection of the gravitational potential, we can assume CMB lensing as an unbiased tracer of the underlying matter overdensity field (Lewis & Challinor 2006).

The quasar overdensity field is related to the matter overdensity field by a window function  $f(z)$ , such that the projected surface density is  $q(\hat{n}) = \int_0^{z_{\text{CMB}}} dz f(z) \delta_m(\chi(z)\hat{n}, z)$  (Peiris & Spergel 2000):

$$f(z) = \frac{b(z)dN/dz}{\int dz' \frac{dN}{dz'}} + \frac{3}{2H(z)}\Omega_0 H_0^2 (1+z)g(z)(5s-2). \quad (3)$$

In the previous equation, the first term is the normalized,

bias-weighted redshift distribution of the quasars. The second term is the magnification bias, which accounts for the change in the density of the sources due to lensing magnification (Moessner et al. 1997; Scranton et al. 2005). This term is negligible compared to the intrinsic clustering of the quasars, and for this reason we ignore it for simplicity. For a full expression of  $g(z)$ , see Peiris & Spergel (2000); Sherwin et al. (2012).

On large scales, we expect the quasar bias to be a constant. On smaller scales, however, the scale-dependence of the bias has been supported by many measurements and it is predicted by theory (Amendola et al. 2017; Giusarma et al. 2018).

Many bias models have been proposed. We will consider an effective power law parametrization of the scale dependence of the bias:

$$b(k) = b_1 + b_2 \left( \frac{k}{k_0} \right)^n \quad (4)$$

where  $k_0$  is an arbitrary reference scale we set to be  $1h \text{ Mpc}^{-1}$ , such that  $b_2$  is a dimensionless parameter (Amendola et al. 2017). The case  $n = 0$  corresponds to a scale-independent bias.

Desjacques et al. (2016) and Modi et al. (2017) reported an  $n = 2$  behavior at scales  $0.1 \lesssim k \lesssim 0.5h \text{ Mpc}^{-1}$  for the linear halo bias, based on results from N-body simulations. We will test this form for the scale-dependent quasar bias.

If the selection functions of the dark matter tracers are slowly varying compared to the scale we are probing, the Limber approximation (Limber 1954; Lewis & Challinor 2006) is expected to be valid at  $\ell \gtrsim 30$ . Assuming a flat universe, the quasar-CMB lensing convergence angular cross-power spectrum is given by:

$$C_l^{\kappa q} = \int \frac{dz}{c} \frac{H(z)}{\chi^2(z)} W(z) f(z) P_{mm} \left( k = \frac{l}{\chi(z)}, z \right) \quad (5)$$

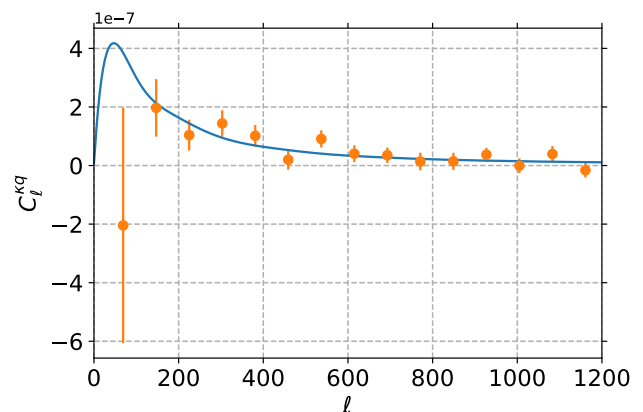
where  $f(z)$  is the bias-weighted redshift distribution, and  $P_{mm}(k, z)$  is the 3D matter power spectrum.

An advantage of using the cross-correlation between CMB lensing and quasars over doing quasar auto-correlation, is that the quasar clustering-matter cross power spectrum has a linear dependence on the quasar bias, from the bias-weighted redshift distribution. Moreover, measuring this cross-correlation in addition to the auto-correlation of quasars helps break the degeneracy between quasar bias  $b_q$  and amplitude of fluctuations<sup>1</sup>  $\sigma_8$ , thus improving our constraints on  $\sigma_8$ . The cross-correlation is also less likely to be affected by systematics in the quasar sample (Sherwin et al. 2012; Geach et al. 2013; Bleem et al. 2012).

### 3 DATA AND METHODS

#### 3.1 CMB lensing map

We use the CMB lensing convergence map published by the Planck Collaboration (Planck Collaboration et al. 2015c). The *Planck* satellite, which was launched in 2009, observed



**Figure 2.** The CMB lensing-quasar overdensity angular cross-power spectrum. The data points are in orange, and the blue solid curve is the calculated theory curve. The significance of the cross-power spectrum signal is  $5.4\sigma$ .

the temperature and polarization fields of the cosmic background radiation over the whole sky at various frequencies. Maps of the temperature and polarization fields of the CMB covering 70% of the sky are produced (Planck Collaboration et al. 2015a). The Planck minimum-variance CMB lensing potential field is reconstructed using the CMB maps produced by the SMICA code, and combines the five quadratic estimators of the correlations of the CMB temperature ( $T$ ) and polarizations ( $E, B$ ). The map underwent several systematic and null tests, that showed that any contamination is small compared to the statistical errors.

#### 3.2 Quasar map

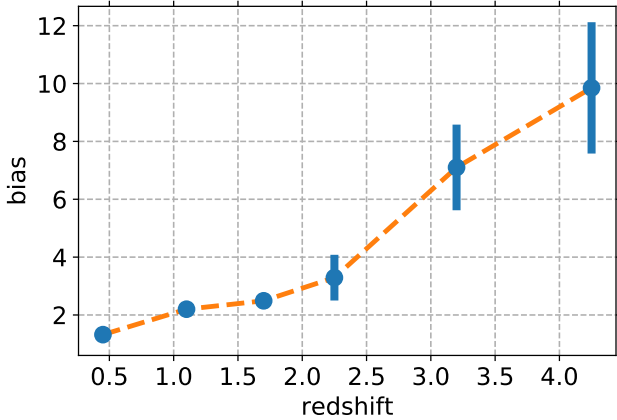
We use quasars from the extended-Baryon Oscillation Spectroscopic Survey (eBOSS, Dawson et al. (2016); Zhao et al. (2016)), which started in July 2014, as an extension to the Baryon Oscillation Spectroscopic Survey (BOSS) (Dawson et al. 2013). BOSS probed the BAO at a scale of roughly  $100 h^{-1} \text{ Mpc}$ , using mostly galaxies at  $z < 0.7$  and neutral hydrogen clouds in the Lyman- $\alpha$  forest at  $z > 2.1$ .

eBOSS aims to probe four different dark matter tracers at redshift ranges that are not covered in previous surveys, and map the large scale structures over the redshift range  $0.6 < z < 2.2$ , which is previously unconstrained by BOSS. The full eBOSS quasar catalog (Myers et al. 2015) is expected to contain 500,000 spectroscopically-confirmed quasars over an area of  $7500 \text{ deg}^2$  by the end of the survey and provide the first BAO distance measurement over the range  $0.9 < z < 2.2$ . The eBOSS quasars will also provide tests of General Relativity on the cosmological scales through measurements of the redshift-space distortion, and new constraints on the summed mass of all known neutrino species.

We use the quasars from the eBOSS DR14 LSS catalog<sup>2</sup> (Myers et al. 2015), which contains 142,017 quasars between

<sup>1</sup> This is because the auto-correlation measures  $b_q^2 \sigma_8^2$ , while the cross correlation is proportional to  $b_q \sigma_8^2$

<sup>2</sup> <https://data.sdss.org/sas/ebosswork/eboss/lss/catalogs/catalogs-DR14/>



**Figure 3.** The fiducial bias-redshift model used in the calculation, obtained by interpolating the data points in Shen et al. (2009). The paper also provides estimates of the error in the quasar bias, which are shown as error bars in the plot. The dashed line in orange is the interpolated result.

$0.9 < z < 2.2$  and has an effective redshift of 1.51. The redshift distribution of the selected quasars is shown in Fig. 1. We construct an overdensity map ( $q_i = \frac{n_i - \bar{n}}{\bar{n}}$ , where  $i$  is the pixel number) of these quasars. The map is converted into HEALPix format with  $N_{\text{side}} = 2048$  to match the resolution of the CMB lensing convergence map. We find the quasar footprint by downgrading the resolution of the quasar map to  $N_{\text{side}} = 32$  and identifying the empty pixels in the map.

### 3.3 Estimator for the angular power spectrum

We use a pseudo- $C_l$  estimator (Wandelt et al. 2000) to calculate the angular cross-power spectrum from the data

$$\hat{C}_l^{\kappa q} = \frac{1}{f_{\text{sky}}^{\kappa q} (2l+1)} \sum_{m=-l}^l \kappa_{lm}^* q_{lm} \quad (6)$$

where  $f_{\text{sky}}^{\kappa q}$  is the fraction of the sky shared by the quasar map and the CMB lensing convergence map.  $\kappa_{lm}$  is the spherical harmonic transform of the CMB lensing convergence map, and  $q_{lm}$  is the spherical harmonic transform of the quasar overdensity map.

In the Fisher approximation, the theoretical error in each bin  $A$  of  $\hat{C}_l^{\kappa q}$  can be estimated using (Cabr e et al. 2008)

$$\frac{1}{\sigma^2(A)} = \sum_{l_{\min}(A) < l < l_{\max}(A)} \frac{f_{\text{sky}}^{\kappa q} (2l+1)}{(C_l^{\kappa q})^2 + C_l^{\kappa \kappa} C_l^{q q}} \quad (7)$$

where  $C_l^{\kappa \kappa}$  and  $C_l^{q q}$  are the CMB lensing and quasar auto-power spectra, including both signal and noise. The contribution of error from the  $C_l^{\kappa \kappa} C_l^{q q}$  term should dominate the contribution from the cross term. The auto-spectra can be estimated similarly:

$$\hat{C}_l^{\kappa \kappa} = \frac{1}{f_{\text{sky}}^{\kappa} (2l+1)} \sum_{m=-l}^l |\kappa_{lm}|^2 \quad (8)$$

and

$$\hat{C}_l^{q q} = \frac{1}{f_{\text{sky}}^q (2l+1)} \sum_{m=-l}^l |q_{lm}|^2 \quad (9)$$

where  $f_{\text{sky}}^{\kappa}$  is the sky fraction of the CMB lensing convergence map, and  $f_{\text{sky}}^q$  is the sky fraction of the quasar overdensity map. We bin the cross-power spectrum into 15 bands in the range  $30 \leq \ell < 1200$ . We choose  $\ell_{\min} = 30$  because the Limber approximation breaks down on larger scales. We choose  $\ell_{\max} = 1200$  because of the uncertainty on modeling the bias and power spectrum on smaller scales. The signal-to-noise also drops significantly for  $\ell > \ell_{\max}$  (Lewis & Challinor 2006; Kirk et al. 2015).

## 4 RESULTS

### 4.1 Cross-correlation

The cross-correlation results are shown in Fig. 2. The theoretical curve is calculated using Equation 5. We use the redshift distribution in Fig. 1 and the CMB lensing kernel in Equation 2. The sample variance fluctuations are of order  $\sim 10\%$  per bin in redshift, due to the large number of quasars in the bin. We use the full  $dN/dz$  from Figure 1 in the theory calculation, which yields a smooth result since it is integrated over a broad kernel. The theory curve should not be sensitive to binning and interpolation, since the weighting functions are slowly varying with redshift. We use CAMB (Lewis et al. 2000) to compute the matter power spectrum. The nonlinear matter power spectrum (HALOFIT, Smith et al. (2003); Takahashi et al. (2012)) is used in this calculation. The linear matter power spectrum produces similar results because the signal mainly comes from angular scales ( $\ell < 600$ ) corresponding to linear scales.

We assume a fiducial bias-redshift model from Shen et al. (2009) in the theory calculation, shown in Fig. 3. The fiducial bias model is based on the amplitude of the quasar correlation function from the SDSS DR5 quasar sample. We fit a scaled version of the fiducial bias-redshift relation to the data to find the best-fit scaling parameter ( $b_q/b_{\text{fid}}$ ), and the theory curve is a good fit. With 14 degrees of freedom, the chi-squared value for the best-fit theory curve is  $\chi_{\text{th}}^2 = 12.9$ .

The significance of the cross-correlation is  $\sqrt{\chi_0^2 - \chi_{\text{th}}^2} = 5.4\sigma$ , where  $\chi_0^2$  is the chi-squared value for the null hypothesis. The detection significance is also the best-fit scaling parameter divided by its error.

All points are included in the model fits to the data. Despite the theory curve being a good fit to the data points, the first bin is more than  $1\sigma$  from the theoretical prediction and shows an anti-correlation between the CMB lensing map and the quasar overdensity, albeit having a large uncertainty. Giannantonio et al. (2016), which uses the CMB lensing data from the South Pole Telescope (Story et al. 2015), and Pullen et al. (2016) also reported a deficit of power in the low  $\ell$  region of the CMB lensing-galaxy angular cross power spectrum. We do not have an explanation for the cause of this deficit of power. However, we can rule out a list of systematics, described in detail later in Section 5, as causes of this deficit.

$n$	$b_0$	$\sigma(b_0)$	$b_1$	$\sigma(b_1)$	$\chi^2$
-2	2.80	0.50	-0.0010	0.0006	11.0
-1	3.53	0.81	-0.067	0.041	11.2
0	2.43	0.45	-	-	12.9
1	1.85	0.89	6.37	8.71	13.3
2	2.26	0.59	13.0	33.3	13.7

**Table 1.** Selected results for the scale-dependent bias fit. In the third row,  $n = 0$  corresponds to a scale-independent bias.

## 4.2 Quasar bias

The fiducial bias-redshift model used in the calculation is obtained by interpolating the data in Shen et al. (2009). Although it uses a different quasar catalog than the one in our analysis, we choose this as a convenient model because the theoretical cross-power spectrum does not have a strong dependence on the detailed form of the bias model (Sherwin et al. 2012). The fiducial model is shown in Fig. 3. From this model we find  $b_q/b_{\text{fid}} = 1.01 \pm 0.19$ . At the effective redshift of our quasar sample ( $z \approx 1.51$ ), the fiducial model gives a bias  $b_{\text{fid}} = 2.4$ . Combining these results we get a quasar linear bias of  $b_q = 2.43 \pm 0.45$ , with  $5.4\sigma$  significance.

We also fit for the scale-dependent bias in Equation 4, by fixing  $n$  at various values. Table 1 shows some of the results. In the  $n = 2$  case, we have  $b_1 = 2.26 \pm 0.59$  and  $b_2 = 13.0 \pm 33.3$ . We conclude that the data does not yield a strong constraint on the scale dependent bias. This, however, could be due to the low number density of quasars in the survey. We expect that the better sensitivity and resolution of future surveys will allow better constraints on the scale dependence of both the bias and the matter power spectrum (Abazajian et al. 2016).

## 4.3 Quasar host halo mass

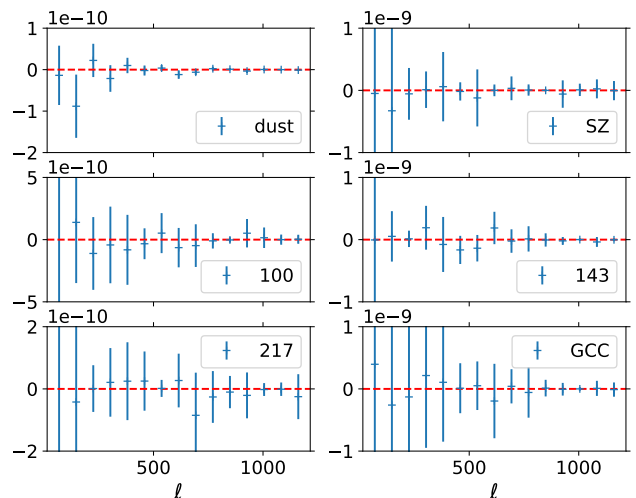
As shown in Fig. 3, the quasar bias generally increases with redshift, and the bias is expected to increase with halo mass. However, at higher redshifts, the halos also have less time to grow. Therefore, we would expect a roughly constant halo mass-redshift relation.

We use the bias model provided in Tinker et al. (2010) to relate the scale-independent quasar bias to the peak height of the linear density field,  $\nu = \frac{\delta_c}{\sigma(M)}$ , where  $\delta_c = 1.686$  is the critical overdensity for collapse, and calculate a corresponding characteristic host halo mass. We assume the ratio between the halo mass density and the average matter density of universe is  $\Delta = 200$ . We find the characteristic host halo mass to be  $\log_{10}\left(\frac{M}{h^{-1}M_\odot}\right) = 12.54^{+0.25}_{-0.36}$ . This is consistent with previous estimates for BOSS/eBOSS quasars at similar redshifts (White et al. 2012; Laurent et al. 2017).

## 5 MEASUREMENT SYSTEMATICS AND UNCERTAINTIES

### 5.1 Systematic effects

Residual foregrounds in the CMB map that are correlated with the large scale structure probed by the eBOSS quasars can lead to biases to the CMB lensing cross-correlation (van Engelen et al. 2014; Osborne et al. 2014; Ferraro & Hill 2018;



**Figure 4.** Check for possible systematic effects on the cross power spectrum due to contaminants. Here we show the right hand side of Equation 10 for different foregrounds. The dust plot is the bias due to dust emission. The SZ plot is the bias due to the Planck SZ catalog. The 100, 143, and 217 plots are the biases from Planck Catalog of Compact Objects, corresponding to the labeled frequency. The GCC plot is the bias from the Planck Galactic cold clumps.

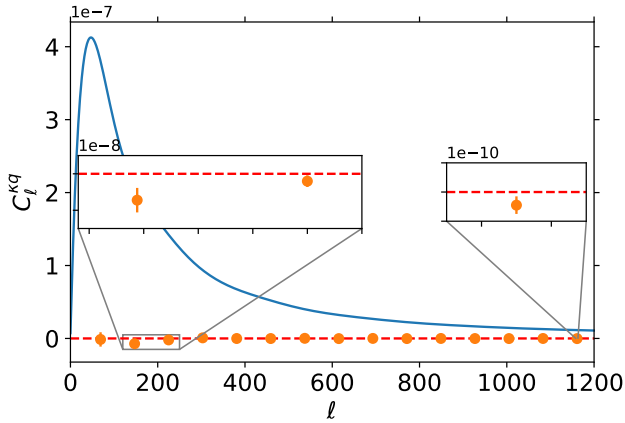
Pullen et al. 2016). Mitigation strategies have been proposed (Madhavacheril & Hill 2018; Schaan & Ferraro 2018), and based on previous work we expect the bias to cross-correlations with Planck lensing to be at most a few percent, considerably smaller than our statistical significance.

Nonetheless, we check for contamination from galactic dust emission, point sources, and SZ effect. We use the Second Planck SZ Catalog (Planck Collaboration et al. 2015e), which includes sources detected through the SZ effect (Sunyaev & Zel'dovich 1980), the Schlegel et al. (1998) dust infrared emission map for estimation of CMB radiation foregrounds, the Planck Catalog of Galactic cold clumps (Planck Collaboration et al. 2015f), and the overdensity maps constructed from the Second Planck Catalog of Compact Sources (Planck Collaboration et al. 2015d) at frequencies 100 GHz, 143 GHz, and 217 GHz.

If systematic effects were added linearly to the observed CMB lensing map and quasar map (Ross et al. 2011; Ho et al. 2012), the bias to the cross correlation would be given by (Giannantonio et al. 2016):

$$\Delta\hat{C}_l^{kq} = \sum_s \frac{\hat{C}_l^{ks} \hat{C}_l^{qs}}{\hat{C}_l^{ss}} \quad (10)$$

where  $s$  is the map for the systematics. In Equation 10, we estimate the amplitude of the systematic  $s$  for each data set by cross-correlating the data and the systematic template, and propagate these to the bias in the observed cross power spectrum. Although the lensing map is obtained through non-linear operations on the CMB map, and therefore the assumption of linearity is not satisfied, estimating the quantity above is still a powerful null-test. If significant contamination was found, Equation 10 should not be used to correct for the bias, but more sophisticated mitigation techniques



**Figure 5.** The cross-power spectrum from the null test. The error bars are obtained the same way as before (Equation 7). The blue curve is the best-fit theoretical cross-power spectra. The zoomed-up subplots show points more than  $1\sigma$  from null. The result is consistent with zero correlation.

should be employed (Schaan & Ferraro 2018; Madhavacheril & Hill 2018; Osborne et al. 2014).

Fig. 4 shows the right hand side of Equation 10. The effects are consistent with null at most scales and we conclude that there is no significant systematic effects due to the contaminants considered above. We calculate the overall systematic error by adding the average absolute biases at each angular scale, weighted by the inverse variance, and find it to be less than 7% of the signal.

## 5.2 Null test

We use a simple null test (Sherwin et al. 2012; Geach et al. 2013) on the CMB lensing-quasar overdensity cross power spectrum to check our result and procedure by cross-correlating the CMB lensing convergence map on one part of the sky with the quasar map on another part of the sky. The result of the null test is shown in Fig. 5. Most bins of the null cross-spectrum fall within  $1\sigma$  of null, and fitting the theoretical spectrum to the null result yields a bias measurement of  $b/b_{\text{fid}} = -0.005 \pm 0.003$ . The best fit chi-square value for the null hypothesis is 11.23, with 14 degrees of freedom. The distribution of the points is consistent with a Gaussian centered at zero.

## 5.3 Covariance matrix

The theoretical error bar for each bin is calculated using Equation 7, which assumes the bins are independent. Limited sky fraction may induce correlation between  $\ell$  bins, and this assumption is only valid when the bins are relatively large ( $\Delta\ell \gtrsim 2/f_{\text{sky}}$ ) (Gaztañaga et al. 2012; Cabré et al. 2008). In our case, the bins are large and should be roughly independent in the limit of large  $\Delta\ell$ .

To compute the full covariance matrix of  $\hat{C}_l^{kq}$ , we use quasar mocks and CMB lensing simulations. The quasar mocks are taken from the QSO EZmocks (effective Zel’dovich approximation mock catalogs) (Chuang et al. 2014), which include 1000 realizations of the quasar map

with the same number of randomly distributed sources. The CMB lensing simulations include 100 realizations of simulated lensing convergence maps (Planck Collaboration et al. 2015c) containing both signal and noise. We cross correlate 100 pairs of the quasar mocks and lensing simulations, and calculate the average of the cross power spectra, to estimate the covariance matrix  $\text{cov}[i, j] = \langle (C(i) - E(C(i)))(C(j) - E(C(j))) \rangle$ . Note that the covariance estimated via this route does not include the  $C_l^{kq}$  part in Equation 7, because the quasar mocks are not correlated with the CMB lensing simulations.

The off-diagonal elements of the covariance matrix are small compared to the on-diagonal elements (Fig. 6), and the diagonal elements mostly agree with the theoretical values, calculated using Equation 7. In both the theoretically predicted error and the covariance matrix, the error in the cross power spectrum decreases with increasing  $\ell$  for  $\ell < 1200$ . The shot noise of the quasars should be a constant contribution of the power spectrum error at all scales. On smaller scales, the error increases again, due to reconstruction noise in the lensing map.

The central value and the uncertainty of the bias estimate change slightly when we use the full covariance matrix, which gives a bias of  $2.42 \pm 0.44$  with a significance of  $5.4\sigma$  and  $\chi_{\text{th}}^2 = 13.9$  for 14 degrees of freedom.

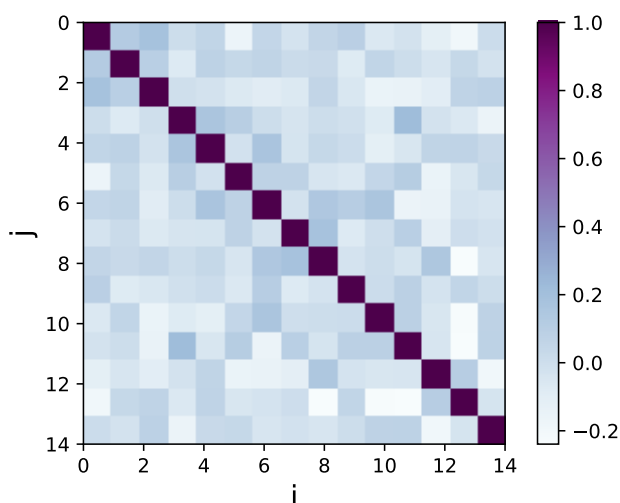
## 6 CONCLUSIONS

We studied the cross-correlation between the Planck CMB lensing convergence map and the eBOSS DR14 quasar map at redshift range  $0.9 < z < 2.2$ , with an effective redshift of  $z_{\text{eff}} \approx 1.51$ , and measure the quasar bias. We found correlation between CMB lensing and the eBOSS quasars, and a quasar bias  $b_q = 2.43 \pm 0.45$  at  $5.4\sigma$  significance, using the theoretically calculated covariance matrix. This is consistent with the result in Laurent et al. (2017). We obtained the covariance matrix from the quasar mocks and lensing simulations, and found it consistent with the theoretical covariance matrix. While the theory curve is a good fit on most of the scales, the first bin shows low cross-correlation between CMB lensing and quasar clustering. The origin of this deficit of power at low- $\ell$  is not known at present.

We performed a simple null test for the cross power spectrum, and the result is mostly consistent with null, with the exception of two low- $\ell$  bins and one near  $\ell_{\text{max}}$ . We checked for several systematics and found no significant contributions from the considered contaminants.

Using the Tinker et al. (2010) model of the relation between halo mass function and clustering, we calculate a characteristic host halo mass for the eBOSS DR14 quasar catalog:  $\log_{10} \left( \frac{M_{200}}{1h^{-1}M_{\odot}} \right) = 12.54^{+0.25}_{-0.36}$ . This is consistent with previous estimates of the quasar host halo mass at similar redshifts (White et al. 2012; Laurent et al. 2017). We also attempted to fit for a scale dependent bias, but did not find evidence for a scale dependent term.

The significance and accuracy of the quasar bias measurement depend on the sample size and number density of the quasar survey (Seljak et al. 2009), so we would expect the signal-to-noise ratio of the detection using this method to improve, as eBOSS continues to expand its sample size (Dawson et al. 2016) and new surveys such as DESI (DESI



**Figure 6.** The normalized covariance matrix of the angular cross power spectrum ( $(\text{cov}[i, j]/\sqrt{\text{cov}[i, i] * \text{cov}[j, j]})$ ), where  $i$  and  $j$  are labels of the bins.

Collaboration et al. 2016) and Euclid (Laureijs et al. 2011) become operational. This will provide more precise measurements of the quasar bias as a function of redshift, scale, etc, and open paths to better understanding of the various properties of quasars, including the host halo mass and duty cycle (Martini & Weinberg 2001). The improved bias measurement could also provide good constraints on the galaxy formation models (Contreras et al. 2013), general relativity and modified gravity (Acquaviva et al. 2008), and the properties of dark matter and dark energy (Das & Spergel 2009).

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## REFERENCES

Abazajian K. N., et al. 2016, preprint, ([arXiv:1610.02743](https://arxiv.org/abs/1610.02743))  
 Acquaviva V., Hajian A., Spergel D. N., Das S. 2008, *Phys. Rev. D*, **78**, 043514

Aghanim N., Majumdar S., & Silk J. 2008, Reports on Progress in Physics, **71**, 066902. ([arXiv:0711.0518](https://arxiv.org/abs/0711.0518))  
 Amendola L., Menegoni E., Di Porto C., Corsi M., Branchini E. 2017, *Phys. Rev. D* **95**, 023505  
 Baron E., Nugent P. E., Branch D., Hauschildt P. H. 2004, *ApJ*, **616**, L91  
 Blanton M. R., et al. 2017, *AJ*, **154**, 28  
 Bleem, L. E., van Engelen, A., Holder, G. P., et al. 2012, *ApJ*, **753**, L9  
 Bridle S. L., Zehavi I., Dekel A., Lahav O., Hobson M. P., Lasenby A. N. 2001, *MNRAS*, **321**, 333  
 Cabré A., Fosalba P., Gaztañaga E., Manera M. 2008, *MNRAS*, **381**, 11  
 Contreras S., Baugh C. M., Norberg P., Padilla N. 2013, *MNRAS*, **432**, 2717  
 Chuang C-H, Kitaura F-S, Prada F., Zhao C., Yepes G. 2014, *MNRAS*, **446**, 2621  
 Das S., Spergel D. N. 2009, *Phys. Rev. D*, **79**, 043509  
 Dawson K. S., Schlegel D. J., Ahn C. P., et al. 2013, *AJ*, **145**, 10  
 Dawson K. S., Kneib J-P, Percival W. J., et al. 2016, *AJ*, **151**, 44  
 DESI Collaboration et al. 2016, preprint, ([arXiv:1611.00036](https://arxiv.org/abs/1611.00036)).  
 Desjacques V., Jeong D., Schmidt F., 2016, *Physics Reports*, **733**, 1. ([arXiv:1611.09787](https://arxiv.org/abs/1611.09787))  
 DiPompeo M. A., et al. 2015, *MNRAS* **446**, 3492. ([arXiv:1411.0527](https://arxiv.org/abs/1411.0527))  
 Ferraro, S., & Hill, J. C. 2018, *Phys. Rev. D*, **97**, 023512  
 Gaztañaga E., et al., 2012, *MNRAS*, **422**, 2904  
 Geach J. E., et al., 2013, *ApJ Letters*, **776**, L41. ([arXiv:1307.1706](https://arxiv.org/abs/1307.1706))  
 Giannantonio T., et al., 2016, *MNRAS*, **456**, 3213  
 Giusarma E., et al., 2018, preprint. ([arXiv:1802.08694](https://arxiv.org/abs/1802.08694))  
 Hinshaw G., et al. 2012, *ApJS*, **208**, 19  
 Hirata C. M., Padmanabhan N., Seljak U., Schlegel D., Brinkmann J. 2004, *Phys. Rev. D*, **70**, 103501  
 Hirata C. M., Ho S., Padmanabhan N., Seljak U., Bahcall N. 2008, *Phys. Rev. D*, **78**, 043520  
 Ho S., Cuesta A., et al. 2012, *ApJ*, **761**, 14  
 Hu W. & Okamoto T. 2002, *ApJ*, **574**, 566  
 Kaiser, N. 1984, *ApJ*, **284**, L9  
 Kirk, D., et al. 2015 *MNRAS*, **459**, 21  
 Kofman, L. A., Gnedin, N. Y., & Bahcall, N. A. 1993, *ApJ*, **413**, 1  
 Kormendy J. & Richstone D. 1995, *ARA&A*, **33**, 581  
 Kuntz A. 2015, *A&A*, **584**, A53  
 Laureijs R., et al. 2011, preprint. ([arXiv:1110.3193](https://arxiv.org/abs/1110.3193))  
 Laurent P., Eftekharzadeh S., Le Goff J.-M. et al. 2017, *JCAP*, **07**, 017  
 Lewis A., Challinor A., & Lasenby, A. 2000, *ApJ*, **538**, 473  
 Lewis A. & Challinor A. 2006, *Phys. Rept.*, **429**, 1. [arXiv:astro-ph/0601594](https://arxiv.org/abs/astro-ph/0601594)  
 Limber D. N. 1954, *ApJ*, **119**, 655  
 Madhavacheril, M. S., & Hill, J. C. 2018, [arXiv:1802.08230](https://arxiv.org/abs/1802.08230)  
 Martini P. & Weinberg D. H. 2001, *ApJ*, **547**, 12  
 Marziani P., Sulentic J. 2014, *Advances in Space Research*, **54**, 1331  
 Modi C., Castorina E., Seljak U. 2017, *MNRAS*, **472**, 3959  
 Moessner R., Jain B., Villumsen J. V. 1997, *MNRAS*, **294**, 291. ([arXiv:astro-ph/9708271](https://arxiv.org/abs/astro-ph/9708271))  
 Mortlock D. J. 2015, preprint. ([arXiv:1511.01107](https://arxiv.org/abs/1511.01107))  
 Myers A. D., Palanque-Delabrouille N., Prakash A., et al. 2015, *ApJS*, **221**, 27. [arXiv:1508.04472](https://arxiv.org/abs/1508.04472)  
 Osborne, S. J., Hanson, D., & Doré, O. 2014, *J. Cosmology Astropart. Phys.*, **3**, 024  
 Peiris H. V. & Spergel D. N. 2000, *ApJ*, **540**, 605  
 Peacock J. A. & Smith R. E. 2000, *MNRAS*, **318**, 1144  
 Planck Collaboration et al. 2015a, preprint, ([arXiv:1502.01582](https://arxiv.org/abs/1502.01582))  
 Planck Collaboration et al. 2015b, *A&A*, **594**, A13  
 Planck Collaboration et al. 2015c, *A&A*, **594**, A15  
 Planck Collaboration et al. 2015d, *A&A*, **594**, A26



- Planck Collaboration et al. 2015e, *A&A*, 594, A27  
Planck Collaboration et al. 2015f, *A&A*, 594, A28  
Planck Collaboration et al. 2018, preprint, ([arXiv:1807.06209](#))  
Pullen A. R., Alam S., He S., Ho S. 2016, *MNRAS*, 460, 4098  
Ross A. J., Ho S., et al., 2011, *MNRAS*, 417, 1350  
Sachs, R. K. & Wolfe, A. M. 1967, *ApJ*, 147, 73  
Salpeter E. E. 1964, *ApJ*, 140, 796  
Schaan, E., & Ferraro, S. 2018, [arXiv:1804.06403](#)  
Schlegel D. J., Finkbeiner D. P., Davis M. 1998, *ApJ*, 500, 525  
Scranton R., et al. 2005, *ApJ*, 633, 589  
Seljak U. 2000, *MNRAS*, 318, 203  
Seljak U., Hamaus N., Desjacques V. 2009, *Phys. Rev. Lett.*, 103, 091303  
Shen Y., Strauss M. A., Ross N. P., et al. 2009, *ApJ*, 697, 1656  
Sherwin B. D., Das S., et al. 2012, *Phys. Rev. D* 86, 083006  
Smith K. M., Zahn O., Dore O. 2007, *Phys. Rev. D* 76, 043510. ([arXiv:0705.3980](#))  
Smith R. E., Peacock J. A., Jenkins A., et al. 2003, *MNRAS*, 341, 1311  
Spergel D. N., Bolte M., Freedman W. 1997, *PNAS*, 94, 6579  
Story K. T., et al. 2015, *ApJ*, 810, 50. ([arXiv:1412.4760](#))  
Sunyaev R. A., & Zel'dovich Ya. B. 1980, *ARA&A*, 18, 537  
Takahashi R., Sato M., Nishimichi T., Taruya A., Oguri M. 2012, *ApJ*, 761, 152  
Tinker J. L., Robertson B. E., Kravtsov A. V., et al. 2010, *ApJ*, 724, 878  
van Engelen, A., Bhattacharya, S., Sehgal, N., et al. 2014, *ApJ*, 786, 13  
Wandelt B. D., Hivon E., Gorski K. M. 2000, preprint. ([arXiv:astro-ph/0008111](#))  
Weinberg D. H., Dav'e R., Katz N., Kollmeier J. A. 2003, *AIP Conference Proceedings* 666, 157  
White S. D. M. & Rees M. J. 1978, *MNRAS*, 183, 341  
White M., Myers A. D., Ross N. P., et al. 2012, *MNRAS*, 424, 933  
Zaldarriaga M. & Seljak U. 1999, *Phys. Rev. D* 59, 123507  
Zhao G.-B., Wang Y., Ross A. J., et al. 2016, *MNRAS*, 457, 2377

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