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# STELLAR WEAK INTERACTION RATES<sup>1</sup> FOR INTERMEDIATE MASS NUCLEI. III. RATE TABLES FOR THE FREE NUCLEONS AND NUCLEI WITH A = 21 TO A = 60

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#### ABSTRACT

Stellar electron and positron emission rates and continuum electron and positron capture rates, as well as the associated neutrino energy loss rates, are tabulated for the free nucleons and 226 nuclei with masses between A = 21 and 60. These rates were calculated in accordance with the procedure described in Papers I and II of this series and are presented here in tabular form on an abbreviated temperature and density grid. Results of these calculations on a detailed temperature and density grid are available in computer readable form on magnetic tape upon request to MJN. The stellar weak rate calculation procedure is reviewed, and the results are discussed. Comparison of the stellar weak rates to terrestrial decay rates are made where possible.

Subject headings: equation of state — neutrinos — nuclear reactions — stars: collapsed — stars: interiors — stars: supernovae

#### I. INTRODUCTION

In this work, electron and positron emission rates in stars, continuum electron and positron capture rates in stars, and the associated neutrino energy loss rates are tabulated for free protons and neutrons as well as 226 nuclei with masses between A = 21 and 60. The rate computations were performed following the procedures discussed in Papers I and II of this series (Fuller, Fowler, and Newman 1980, hereafter  $F^2NI$ ; Fuller, Fowler, and Newman 1982, hereafter  $F^2NI$ ). The previous papers are concerned with the details of stellar weak rate calculations and the nuclear physics of the Gamow-Teller strength distribution.

The purpose of this paper is to discuss the results of the stellar rate computations in comparison with other work and with known terrestrial decay rates and provide ready access to printed rates for active investigations in the fields of stellar evolution and nucleosynthesis. The rates presented here in tabular form are reproduced, by necessity, on an abbreviated temperature and density grid which will allow a fair estimate of stellar rates to be made for most astrophysical environments. Where more accurate stellar rates are required, as in stellar evolution and nucleosynthesis calculations, the reader is urged to

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write to MJN and request the stellar rate magnetic tape. This tape presents the computation results on a far more detailed temperature and density grid in a computer readable form suitable for interpolation in temperature and density.

The temperature and density grid employed in these calculations covers conditions ranging from the relatively mild environments characteristic of hydrostatic carbon and oxygen burning through the more extreme conditions characteristic of silicon burning and the onset of core collapse. The temperatures cover the range  $0.01 \le T_9 \le 100$ . (0.862 keV  $\le kT \le 8.617$  MeV), while the densities cover the range  $10 \le \rho/\mu_e$  (g cm<sup>-3</sup>)  $\le 10^{11}$ ; thus, the electron gas ranges from nondegenerate to degenerate conditions, with electron Fermi energies approaching 25 MeV for the highest densities.

The nuclear weak rates computed here are important in determining the neutronization and neutrino energy loss rates during stellar evolution and collapse (Weaver and Woosley 1981; Arnett and Thielemann 1982). The outcome of stellar core collapse and bounce calculations depends on the previous neutronization history of the star and on electron capture during the collapse phase (Van Riper and Lattimer 1981; Baym, Bethe, and Brown 1981). Ultimately, the nuclear weak rates presented here will be important in determinations of the nucleosynthesis yields of the elements and their isotopes, both during hydrostatic burning regimes and the subsequent collapse and explosive burning phases. 280

1982ApJS...48..279F

The nuclei whose stellar weak rates are computed here range in mass from A = 21 up to A = 60. These include many nuclei of astrophysical interest in the stellar conditions discussed above. Subsequent work will extend the rate survey presented here to include neutron-rich nuclei in the A = 60 to A = 75 mass range for use in the supernova problem. Concomitant with this extension in mass range will come an extension in the density range to include  $\rho/\mu_e = 10^{12}$  g cm<sup>-3</sup>. Of interest in the *r*-process and *s*-process are the weak transition rates of heavy neutron-rich nuclei at temperatures and densities which are relatively low compared to those encountered in the late stages of stellar evolution discussed above. In these less extreme conditions unmeasured forbidden transitions and bound state electron capture may be important. The reader is referred to Cosner and Truran (1981) for the most up-to-date treatment of the problem.

#### II. REVIEW OF STELLAR RATE COMPUTATION PROCEDURE

The stellar rate computation problem is shown schematically in Figure 1 for  $T^{<}$  and  $T^{>}$  nuclei. The  $T^{<} \rightleftharpoons T^{>}$  stellar transition rates must be computed. Some 20 discrete states are typically included in each nucleus for many of the cases considered here. The discrete state excitation energies, spins, and parities are taken from the experimental tabulations of Lederer et al. (1978) and Endt and van der Leun (1978) wherever possible. Isospin symmetry is used to obtain excitation energies, spins, and parities in unmeasured nuclei whose mirror nuclei are well studied. For very neutron-rich unmeasured nuclei, discrete state energies, spins, and parities are inferred from isotopes with similar shell structure. The discrete states in the  $T^{<}(T^{>})$  nucleus are denoted by  $E^{<}(E^{>})$  in Figure 1, while the resonant states are denoted by  $R^{<}$  ( $R^{>}$ ). The symbols also designate the energies of the discrete and resonant states.

The weak transition rate from the *i*th state of the parent to the *j*th state of the daughter nucleus is given by

$$\lambda_{ij} = \ln 2 \frac{f_{ij}(T, \rho, U_{\rm F})}{(ft)_{ij}}, \qquad (1)$$

where  $(ft)_{ij}$  is the comparative half-life, which is related to the allowed weak interaction matrix elements by equations (I-2a) and (I-2b) (I and II denote equations in Papers I and II of this series, respectively). The  $f_{ij}(T, \rho, U_F)$  are the phase space factors for either electron or positron emission (eq. [I-3a]), continuum electron or positron capture (eq. [I-3b]), or the associated neutrino energy loss rates for these processes (eqs. [I-6a] and [I-6b]).

In principle, weak transition matrix elements are required between all  $E^{<}$  and  $E^{>}$  discrete states. Experi-



FIG. 1.—A schematic representation of a typical excited state problem is shown.  $E_1^<, E_2^<, \ldots, E_i^<$  are the discrete states of the  $T^<$  nucleus.  $E_1^>, E_2^>, \ldots, E_j^>$  are the discrete states in the  $T^>$ nucleus.  $A_1^<, A_2^<, \ldots, A_j^<$  are the analogs of  $E_1^>, E_2^>, \ldots, E_j^>$ . Transitions from the discrete states of the  $T^<$  nucleus to Gamow-Teller collective resonances  $(R_1^>, R_2^>, \ldots, R_i^>)$  in the  $T^>$  nucleus, and the reverse transitions, are shown with rates  $\lambda_{ii} = \lambda_{er}$ . Similarly, transitions from the discrete states of the  $T^>$  nucleus to collective Gamow-Teller resonances in the  $T^<$  nucleus  $(R_1^<, R_2^<, \ldots, R_j^<)$ , and the reverse transitions are shown with rates  $\lambda_{jj} = \lambda_{re}$ . For clarity the Fermi transitions  $\lambda_{jj}^F$  between  $E_j^> \rightleftharpoons A_j^<$ are not illustrated.

mentally determined transition matrix elements are employed where known, whether or not they correspond to allowed transitions. Unmeasured Gamow-Teller allowed transitions are assigned  $\log ft = 5.0$  (F<sup>2</sup>NI; Gleit, Tang, and Coryell 1968), unless experiment gives an indication of  $\log ft \gg 5$  despite satisfaction of the allowed selection rules. Such cases of hindered allowed transitions were assigned  $\log ft = 99.9$  where they might otherwise be important in the determination of low temperature, low density rates. More discussion of the low temperature, low density rates and the adjustment of experimental log ft values follows in the next section. Fermi transitions are assigned appropriate matrix elements according to equation (I-11). Unmeasured forbidden transitions are neglected because of the dominant effect of the large number of allowed transitions which contribute in most stellar environments. Some log ft values between the parent ground state and low-lying daughter states have been adjusted to reproduce measured laboratory decay rates using accurately computed *f*-values; this procedure is discussed further in § III.

In order to simulate transitions into and from the continuum, appropriately placed Gamow-Teller resonances are included in each stellar weak rate calculation with the procedure discussed in  $F^2$ NII. The excitation energies of the Gamow-Teller resonances are calculated on the basis of a simple shell model employing the tabulated single particle energies of Seeger and Howard (1975). The procedure for computing the resonance excitation energy differs for  $T^{<} \rightarrow T^{>}$  and  $T^{>} \rightarrow T^{<}$  transitions; the reader is referred to §§ III and IV of F<sup>2</sup>NII for an exhaustive discussion. The sum rules for  $T^{<} \rightleftharpoons T^{>}$ transitions are estimated by constructing a zero-order shell model configuration for the parent nucleus ground state and applying equation (II-16). The observed discrete state strength for transitions from the parent ground state is summed and subtracted from the equation (II-16) sum rule strength. The remaining strength is lumped into the ground state  $\rightleftharpoons$  resonance state transition.

Resonances corresponding to excited states are taken into account here with a special treatment of the population index for ground state  $\rightleftharpoons$  resonance transitions. This procedure is based on the assumptions of constant *Q*-value and sum rule for each discrete state  $\rightleftharpoons$  corresponding state transition and is discussed in detail in § V of  $F^2$ NII. In summary, the contribution of the  $E^{<} \rightarrow R^{>}$  transitions to the total  $T^{<} \rightarrow T^{>}$  transition rate can be calculated by including only the  $E_1^{<} \rightarrow R_1^{>}$ transition and setting the occupation index for the  $E_1^{<}$ state in this transition equal to unity (cf. eqs. [II-43] and [II-44]). The  $R^{<}$  to  $E^{>}$  transitions proceed through the thermal population of the  $R^{<}$  resonances. In the special procedure used in this calculation, the contribution of the  $R^{<} \rightarrow E^{>}$  transitions to the total  $T^{<} \rightarrow T^{>}$  transition rate can be calculated by including only the  $R_1^{<} \rightarrow$  $E_1^>$  transition. In addition, the occupation index for the  $R_1^<$  state in this transition must be set equal to  $P = (G^>/G^<) \exp(-R_1^</kT)$ , where  $G^<$  and  $G^>$  are the nuclear partition functions at the ambient temperature for the  $\overline{T}^{<}$  and  $T^{>}$  nuclei, respectively (cf. eqs. [II-52] and [II-53]). The generalization for the contribution of the  $E^> \rightarrow R^<$  and  $R^> \rightarrow E^<$  transitions to the total  $T^> \rightarrow T^<$  transition rate is obvious.

Fermi transitions involving excited states in the  $T^>$  nucleus explicitly obey the constant Q-value and sum rule assumptions discussed above. Thus, the contributions of these excited states plus that of the ground state to the  $T^> \rightarrow T^<$  rate can be taken into account by calculating  $E_1^> \rightarrow A_1^<$ , where  $A_1^<$  designates the analog state in the  $T^<$  nucleus corresponding to the ground state of the  $T^>$  nucleus. In addition, the occupation index for the  $E_1^>$  state in this transition must be set equal to unity. In the  $T^< \rightarrow T^>$  transition the contribu-

tion of the analog states can be calculated for  $A_1^< \rightarrow E_1^>$ , setting the population index for the  $A_1^<$  state equal to  $P = (G^>/G^<) \exp(-A_1^</kT)$ .

The total weak transition rate  $\lambda$  is given by

$$\lambda = \sum_{i} \sum_{j} P_{i} \lambda_{ij}, \qquad (2)$$

where the sum on *i* is over parent states, the sum on *j* is over daughter states, and  $P_i$  is the occupation index defined by equation (I-9a) and modified in the manner discussed above for transitions involving resonances and analog states.

The phase space factors in equation (1) were computed numerically and checked for electron and positron emission at low temperature and density against the tables of log ft by Gove and Martin (1971). A number of low temperature, high density and high temperature, low density results were checked against easily performed analytic calculations. The definite integrals for the electron and positron emission phase space factors were done by 64 point Gaussian guadrature. The integrands of the electron and positron capture phase space integrals are modulated strongly by the electron or positron distribution function, so that the integrand has a characteristically slowly varying part and an exponentially decaying part, corresponding to the shape of the Fermi-Dirac distribution function. The portion of the improper integrals containing the slowly varying part of the integrand was done with 64 point Gaussian quadrature, and the exponential tail was treated with 32 point Gauss-Laguerre quadrature. For each nuclear transition a table of appropriate phase space factors as a function of  $q_n$  (eq. [I-3c]) was prepared at each temperature and density grid point, and the  $f_{ij}$  and  $f_{ij}^{\nu}$  (eqs. [I-3a], [I-3b], [I-6a], [I-6b]) were obtained by cubic spline interpolation in  $q_n$ . This procedure was checked for electron and positron capture in nondegenerate conditions against the analytic phase space factors in Fowler and Hoyle (1964).

The free nucleons are unique "nuclei" in that they have no excited states. As a result the interpolation procedure described above for calculating weak phase space factors was not used; free nucleon phase space factors were performed by direct numerical integration at each temperature and density point. The matrix elements for the free nucleon transitions are  $|M_F|^2 = 1$  (log ft = 3.791 from I-2b) and  $|M_{GT}|^2 = 3$  (log ft = 3.596  $-\log 3 = 3.118$  from I-2a) to give an overall log ft = 3.035. (In F<sup>2</sup>NI the value  $[G_A/G_V]^2 = 1.567$  was used, and the half-life of the neutron was taken to be 936 s.) Nucleon recoil effects (forbidden transitions) are small at the temperatures and densities considered here and are therefore neglected. For a discussion of forbidden transitions see Fuller (1982).

1982ApJS...48..279F

#### **III. RESULTS AND DISCUSSION**

In this section the results of the stellar weak interaction rate calculations are discussed. The terrestrial decay rates of the nuclear species considered in this survey are tabulated in Table 1 based on experimental measurements when available. These rates can be compared with the lowest temperature and density rates calculated in this work. The stellar rates are tabulated on the very abbreviated temperature and density grid shown in Table 2. As noted previously, the rates can be obtained by request on a detailed grid of temperature and density on magnetic tape.

The stellar rate magnetic tape has electron and positron emission rates, continuum electron and positron

TABLE 1A Terrestrial Weak Interaction Rates  $(s^{-1})$ 

Parent (1)	Daughter (2)	Decay (3)	$Q_n$ (MeV) (4)	Log Rate (5)
n	р	β-	1.293	-2.972
<sup>21</sup> O	<sup>21</sup> F	$\beta^{-}$	8.678	(+0  EST)
<sup>21</sup> F	<sup>21</sup> Ne	΄β <sup>-</sup>	6.197	-0.795
<sup>21</sup> Na	<sup>21</sup> Ne	$\beta^+$	3.036	-1.511
<sup>21</sup> Mg	<sup>21</sup> Na <sup>20</sup> Ne	$\beta^+ p$	12.587	+0.751
<sup>22</sup> F	<sup>22</sup> Ne	΄β <sup></sup>	11.364	-0.786
<sup>22</sup> Na	<sup>22</sup> Ne	$\hat{B}^+$	2.331	-8.117
		ε	2.331	- 9.096
		SUM		-8.074
<sup>22</sup> Mg	<sup>22</sup> Na	$\tilde{B}^+$	4.279	-0.746
<sup>23</sup> F	<sup>23</sup> Ne	$B^{-}$	9.021	-0.502
<sup>23</sup> Ne	<sup>23</sup> Na	$\dot{B}^{-}$	4.886	-1.734
<sup>23</sup> Mg	<sup>23</sup> Na	β <sup>+</sup>	3.548	-1.212
<sup>23</sup> Al	<sup>23</sup> Mg <sup>22</sup> Na	$\beta^+ p$	11.727	+0.169
<sup>24</sup> Ne	<sup>24</sup> Na	B <sup>-</sup>	2.979	-2.466
<sup>24</sup> Na	<sup>24</sup> Mg	'B <sup>-</sup>	6.024	-4.892
$^{24}$ Na <sup>m</sup>	<sup>24</sup> Na	ΪT	0.472	+1.535
	<sup>24</sup> Mg	<u>B</u> -	6.496	(-2  EST)
<sup>24</sup> A1	<sup>24</sup> Mg <sup>20</sup> Ne	$\ddot{B}^+ \alpha$	13.367	-0.475
<sup>24</sup> A1 <sup>m</sup>	<sup>24</sup> Al	ÍT	0.439	+0.695
	<sup>24</sup> Mg <sup>20</sup> Ne	$\hat{B}^+ \alpha$	13 806	-0.428
		SUM	121000	+0.727
<sup>24</sup> Si	<sup>24</sup> A1	$\tilde{B}^+$	10 281	(+0.560)
<sup>25</sup> Ne	<sup>25</sup> Na	$\beta^{-}$	7 711	+0.0627
<sup>25</sup> Na	<sup>25</sup> Mg	$\beta^{-}$	4 344	-1937
<sup>25</sup> A1	<sup>25</sup> Mg	$\beta^{+}$	3 767	-1.015
<sup>25</sup> c;	25 1 24 14	$\rho^{+}$	12 226	+ 0.408
<sup>26</sup> No	$26 M_{\odot}$	$\rho_{\rho}^{\mu}$	0.826	-0.180
26A1	26 Mg	$\rho_{\rho^+}$	9.030	-13.602
AI	Ivig	μ	3.494	-14260
		E SUM	3.474	-14.200 -12.516
26 . 1 m	26 Ma	$\rho^+$	2 777	-0.063
26 <b>s</b> ;	26A1	$^{\mu}_{\beta^+}$	J.122 A 553	-0.903
27 No	27 Ma 26 Ma	ρ ρ- "	4.333	-0.304
27 Ma	27A1	$\rho_{\mu}$	9.400	-2.013
27 <b>c;</b>	27A1	$\rho_{\rho^+}$	5.120 1 708	- 2.915
אנג 27 מ <sup>27</sup>	27 <b>c;</b>	$p_{\rho^+}$	4.270	-0.775
28 NIG	31 28 Ma 27 Ma	р 9-"	11.204	$(\pm 0.434)$ $\pm 1.240$
28 Ma	28 A 1	$p_{\rho-}^n$	14.37/	+ 1.349
28 A 1	- AI 28c:	р 0-	2.343	- 5.038
28 n	28 51	$p_{\rho^+}$	5.154	-2.288
28 c	<sup>-1</sup> S1 28 D	р 0+	13.821	$\pm 0.409$
S	<b>P</b>	b.	10.839	(+0.491)

capture rates (in s<sup>-1</sup>), and the associated  $\nu$  and  $\bar{\nu}$  energy loss rates (in MeV s<sup>-1</sup>) for the free nucleons and 226 nuclei with masses between A = 21 and A = 60 on a temperature-density grid which includes  $0.01 \le T_9 \le 100$ and  $10 \le \rho/\mu_e$  (g cm<sup>-3</sup>)  $\le 10^{11}$ . In particular, the rates are calculated at  $T_9 = 0.01, 0.1, 0.2, 0.4, 0.7, 1.0, 1.5, 2.0,$ 3.0, 5.0, 10.0, 30, and 100 for each density point at log ( $\rho/\mu_e$ ) = 1 to 11 by unit increment; thus, there are 143 temperature-density points at which are computed the rates in s<sup>-1</sup> of the four weak processes and the sum of continuum electron capture plus positron emission and continuum positron capture plus electron emission. In addition, the total  $\nu$ -energy loss rate and the total  $\bar{\nu}$ -energy loss rate in MeV s<sup>-1</sup> are given.

At higher temperatures or at higher densities nuclei are completely ionized in astrophysical environments, and continuum electron capture predominates over

 TABLE 1B

 Terrestrial Weak Interaction Rates (s<sup>-1</sup>)

Parent (1)	Daughter (2)	Decay (3)	$Q_n (MeV)$ (4)	Log Rate (5)
<sup>29</sup> Na	<sup>29</sup> Mg <sup>28</sup> Mg	$\beta^{-}n$	13.921	+ 1.207
<sup>29</sup> Mg	<sup>29</sup> Al	$\beta^{-}$	7.973	-0.305
<sup>29</sup> A1	<sup>29</sup> Si	β <sup>-</sup>	4.192	-2.757
<sup>29</sup> P	<sup>29</sup> Si	β <sup>+</sup>	4.433	-0.772
<sup>29</sup> S	<sup>29</sup> P <sup>29</sup> Si	$\beta^+ p$	13.278	+0.562
<sup>30</sup> Na	<sup>30</sup> Mg <sup>29</sup> Mg	$\beta^{-n}$	18.681	+1.108
<sup>30</sup> Mg	<sup>30</sup> Al	΄ <i>β</i> -	6.611	-0.238
<sup>30</sup> A1	<sup>30</sup> Si	'β <sup></sup>	9.050	-0.726
<sup>30</sup> P	<sup>30</sup> Si	$B^+$	3.716	-2.335
		έ	3.716	-5.241
		SUM		-2.335
<sup>30</sup> S	<sup>30</sup> P	$B^+$	5.631	-0.238
<sup>31</sup> Na	<sup>31</sup> Mg <sup>30</sup> Mg	$\beta^{-}n$	15.021	+1.610
<sup>31</sup> Mg	<sup>31</sup> A1	β <sup>-</sup>	11.711	(+1  EST)
<sup>31</sup> Al	<sup>31</sup> Si	'β <sup>-</sup>	8.360	+0.0346
<sup>31</sup> Si	<sup>31</sup> P	'β <sup></sup>	2.002	-4.134
<sup>31</sup> S	<sup>31</sup> P	΄ β <sup>+</sup>	4.884	-0.574
<sup>31</sup> Cl	<sup>31</sup> S	$\ddot{B}^+$	11.463	(+0.405)
<sup>32</sup> Na	$^{32}Mg$	$\tilde{B}^{-}$	19.811	+1.679
<sup>32</sup> Mg	<sup>32</sup> A1	<i>B</i> <sup>-</sup>	8.911	$(\pm 0 \text{ EST})$
<sup>32</sup> A1	<sup>32</sup> Si	$\tilde{B}^{-}$	13.311	(+1  EST)
<sup>32</sup> Si	<sup>32</sup> P	6-	0.724	-9.694
<sup>32</sup> P	$^{32}S$	8-	2.221	-6.250
<sup>32</sup> Cl	<sup>32</sup> S <sup>31</sup> P <sup>28</sup> Si	$\beta^{+}$ n $\alpha$	12.176	+0.367
<sup>32</sup> Ar	$^{32}$ Cl	$B^{+}$	10.608	(+0.308)
<sup>33</sup> Na	<sup>33</sup> Mg	8-	18.711	+1.540
<sup>33</sup> Mg	<sup>33</sup> A1	$\tilde{B}^{-}$	14.011	(+1  EST)
<sup>33</sup> A1	<sup>33</sup> Si	8-	11 711	(+1  EST)
<sup>33</sup> Si	<sup>33</sup> P	<b>B</b> <sup>-</sup>	6.278	-0.952
<sup>33</sup> P	<sup>33</sup> S	$\tilde{B}^{-}$	0.760	-6499
<sup>33</sup> Cl	<sup>33</sup> S	$\tilde{B}^+$	5 072	-0.559
<sup>33</sup> Ar	<sup>33</sup> Cl <sup>32</sup> S	$\tilde{B}^+ n$	11 107	+0.586
<sup>34</sup> Si	<sup>34</sup> P	$\tilde{B}^{-P}$	4 811	-0.606
<sup>34</sup> P	<sup>34</sup> S	$\tilde{B}^{-}$	5.892	-1253
<sup>34</sup> Cl	<sup>34</sup> S	$\tilde{B}^+$	4.982	-0.343
$^{34}\mathrm{Cl}^m$	<sup>34</sup> S	íт	0.146	-3770
<u> </u>	5	$\hat{R}^+$	5 1 28	-3718

No. 3, 1982

...48..279F

1982ApJS

Parent

(1)

<sup>42</sup> Ti.....

<sup>43</sup>Cl.....

<sup>43</sup>Ar.....

 $^{43}K\ldots$ 

<sup>43</sup>Sc.....

<sup>43</sup>Ti.....

<sup>44</sup>Ar.....

<sup>44</sup>K.....

<sup>44</sup>Sc.....

<sup>44</sup>Sc<sup>*m*</sup>...

<sup>44</sup>Ti.....

Daughter

(2)

<sup>42</sup>Sc

<sup>43</sup>Ar

<sup>43</sup>K

<sup>43</sup>Ca

<sup>43</sup>Ca

<sup>43</sup>Sc

<sup>44</sup>K

<sup>44</sup>Ca

<sup>44</sup>Ca

<sup>44</sup>Sc

<sup>44</sup>Ca

<sup>44</sup>Sc

Log Rate

(5)

+0.540

(-2.341)

-2.670

-5.064

-4.305

+0.151

-3.013

-3.282

-4.332

-5.611

-4.310

-5.489

-7.340

-5.483

-9.330

+0.887

TERRESTRIAL WEAK INTERACTION RATES $(s^{-1})$					
Parent Day (1)	ughter Decay (2) (3)	$Q_n$ (MeV) (4)	Log Rate (5)		
<sup>34</sup> Ar <sup>34</sup> (	$\beta^+$	+ 5.548	-0.0855		
<sup>35</sup> P <sup>35</sup> S	ς 'β <sup></sup>	+4.417	-1.831		
<sup>35</sup> S <sup>35</sup> G	C1 'B-	+0.678	-7.037		
<sup>35</sup> Ar <sup>35</sup> O	$C1 \qquad \beta^+$	+5.454	-0.410		
<sup>35</sup> K <sup>35</sup> A	$\beta^+$	+11.369	(+0.592)		
<sup>36</sup> Cl <sup>36</sup> A	νr <sup>'</sup> β <sup></sup>	+1.221	- 13.144		
<sup>36</sup> S		+0.633	-14.857		
-	$\hat{B}^+$	+0.633	-17.905		
	SUM		-13.135		
<sup>36</sup> K <sup>36</sup> A	$B^+$	+12.294	+0.309		
<sup>36</sup> Ca <sup>36</sup> H	$\overline{K}$ $\overline{B}^+$	+10.265	$(\pm 0.538)$		
<sup>37</sup> S <sup>37</sup> C	$\overline{a}$ $\overline{B}^{-}$	+5.365	-2.636		
<sup>37</sup> Ar <sup>37</sup> C	Γ] ε <sup>-</sup>	+0.303	-6.640		
<sup>37</sup> K <sup>37</sup> A	$B^+$	+5.638	-0.249		
<sup>37</sup> Ca <sup>37</sup> H	$\overline{X}^{36}$ Ar $\beta^+ n$	+11.125	+0.603		
<sup>38</sup> S <sup>38</sup> C	$B^{-r}$	+3.447	-4.168		
<sup>38</sup> Cl <sup>38</sup> A	$B^{-}$	+5428	- 3 509		
<sup>38</sup> Cl <sup>m</sup> <sup>38</sup> C	วี่ โา	+0.671	-0.0135		
<sup>38</sup> K <sup>38</sup> A	$\hat{B}^+$	+5402	-2.819		
<sup>38</sup> K <sup>m</sup> <sup>38</sup> A	$\tilde{B}^+$	+5.102 +5.532	-0.128		
<sup>38</sup> Ca <sup>38</sup> I	$\bar{x}$ $\bar{B}^+$	+6.231	+0.197		
<sup>39</sup> Cl <sup>39</sup> A	$B^-$	+3.949	- 3 686		
<sup>39</sup> Ar <sup>39</sup> H	$\bar{x}$ $\bar{B}^-$	+1.076	-10.088		
<sup>39</sup> Ca <sup>39</sup> A	$\tilde{B}^+$	+6.013	-0.0937		
<sup>40</sup> Cl <sup>40</sup> A	$r B^-$	+8.011	-2.068		
<sup>40</sup> K <sup>40</sup> (	$\tilde{a}$ $\tilde{B}^-$	+1.823	- 16 815		
40	ντ ε <sup>-</sup>	+0.994	-17736		
-	- ° B <sup>+</sup>	+0.994	-21765		
	SUM		-16.765		
<sup>40</sup> Sc <sup>40</sup> C	$Ca^{39}K \qquad B^+ n$	+13.809	+0.581		
<sup>40</sup> Ti <sup>40</sup> S	Sc $\beta^+$	+10.976	(+0.645)		
<sup>41</sup> Cl <sup>41</sup> A	$r B^-$	+6.179	-1.691		
<sup>41</sup> Ar <sup>41</sup> H	$\vec{k}$ $\vec{\beta}^-$	+3.003	-3.978		
<sup>41</sup> Ca <sup>41</sup> H	Κ ε	$-0.089^{\circ}$	7 - 12.658		

TABLE 1C

TABLE 1D TERRESTRIAL WEAK INTERACTION RATES  $(s^{-1})$ 

 $Q_n$  (MeV)

(4)

+6.488

+9.351

+5.119

+2.328

+1.709

+6.350

+4.047

+6.170

+3.144

+3.144

+0.271

+3.415

-0.246

+13.185

Decay

(3)

 $\beta^+$ 

'β-

β⁻

 $\beta^{-}$ 

 $\beta^+$ 

β-

 $\beta^{-}$ 

 $\boldsymbol{\beta}^+$ 

ε\_

IT

ε SUM

ε

 $\beta^+ \alpha$ 

SUM

 $\beta^+\epsilon$ 

 $^{44}\tilde{V}\ldots\ldots$ <sup>44</sup>Ti<sup>40</sup>Ca <sup>45</sup>Ca <sup>45</sup>K.....  $\beta^{-}$ +4.709-3.238<sup>45</sup>Sc β-<sup>45</sup>Ca .... -7.313+0.768<sup>45</sup>Sc<sup>*m*</sup> ... <sup>45</sup>Ti..... <sup>45</sup>Sc ĪT +0.0124+0.349<sup>45</sup>Sc  $\beta^+\epsilon$ +1.552-4.205<sup>45</sup>Ti <sup>45</sup>V .....  $\beta^+$ +6.614(-0.034)<sup>45</sup> V<sup>44</sup> Ti  $\beta^+_{\beta^-}p$  $\beta^-_{\perp}$ <sup>45</sup>Cr .... +11.908+1.142<sup>46</sup>Ca <sup>46</sup>K ..... +8.229-2.220<sup>46</sup>Ca <sup>46</sup>Sc.....  $\beta^+\epsilon^-$ +0.872(HI FRB) <sup>46</sup> Ti <sup>46</sup>Sc.....  $\beta^{-}$ +2.878-7.019<sup>46</sup>Sc -1.431 <sup>46</sup>Sc<sup>*m*</sup> ... ĪΤ +0.143<sup>46</sup>Ti <sup>46</sup>V .....  $\beta^+$ +6.5410.214 <sup>46</sup>Cr ....  $\beta^+$ <sup>46</sup>V +7.099 0.426 <sup>47</sup>Ca <sup>47</sup>K .....  $\beta^{-}$ +7.156-1.402<sup>47</sup>Sc  $\beta^{-}$ <sup>47</sup>Ca .... -5.752+2.499<sup>47</sup>Ti  $\beta^{-}$ <sup>47</sup>Sc..... +1.112-5.630<sup>47</sup>Ti <sup>47</sup>V .....  $\beta^+\epsilon^-$ +2.419-3.451<sup>47</sup>Cr .... <sup>47</sup>V  $\beta^+$ +6.872(+0.052)<sup>48</sup>Ca <sup>48</sup>K ..... β +12.507-0.922<sup>48</sup>Sc . β-<sup>48</sup>Ca .... +0.792(HI FRB) <sup>48</sup>Ti β-<sup>48</sup>Sc..... +4.501-5.356 <sup>48</sup>V ..... <sup>48</sup> Ti +3.504-6.597 $\beta^+$ +3.504-6.604SUM -6.299<sup>48</sup>Cr ....  $^{48}V$ +1.144(-5.049)ε <sup>48</sup>Cr <sup>48</sup>Mn ....  $\beta^+\epsilon^-$ (+1 EST)+13.139<sup>49</sup>C1 .... <sup>49</sup>Ar  $\beta^{-}$ ~18.480 (+1 EST)<sup>49</sup>K  $\beta^{-}$ <sup>49</sup>Ar..... ~15.000 (+1 EST)obtained from the terrestrial electron capture decay rate. Terrestrial electron capture decays generally proceed

bound state electron capture. At low density intermediate mass nuclei are thermally ionized in the range  $T_0 = 0.01$  to 0.5, at low temperature nuclei are completely pressure ionized due to the combined effects of nuclear charge screening and continuum lowering. At zero temperature, for example, <sup>56</sup>Fe will become completely pressure ionized for  $\rho/\mu_e \gtrsim 10^4$  g cm<sup>-3</sup>. In general, whenever the density of continuum electrons is larger than the K-shell electron density at the nucleus, then continuum electron capture dominates over bound state capture. In the low temperature-low density corner of the temperature-density grid used here bound state capture is important; we do not calculate bound state electron capture, only continuum capture.

<sup>41</sup>Ca <sup>41</sup>Sc<sup>40</sup>Ca

<sup>42</sup>K

<sup>42</sup>Ca

<sup>42</sup>Ca

<sup>42</sup>Ca

<sup>41</sup>Sc.....

<sup>41</sup>Ti.....

<sup>42</sup>Ar.....

 $^{42}K\ldots$ 

<sup>42</sup>Sc.....

<sup>42</sup>Sc<sup>m</sup> ...

 $\dot{\beta}^+$ 

β

β

 $\beta^+$ 

β+

 $\beta^+ p$ 

+5.984

+1.114

+4.032

+5.912

+6.529

+12.351

+0.0656

+0.938

-9.177

-4.807

+0.0070

-1.952

An estimate of the bound state electron capture rate for nuclear transitions where it is important can be

through the capture of electrons in low-lying atomic levels, usually K-shell electrons. Table 1 lists all of the nuclei considered in this survey, together with their terrestrial decay modes and rates, where these quantities are known. The logarithms of the electron capture rates for a full K-shell can be read from this table. We designate these rates by  $\lambda_{\kappa}$ .

As the temperature rises at a given density in a stellar environment, the occupation of the K-shell is reduced from two electrons to one at a critical value of the temperature. Similarly, at a given temperature the oc284

TABLE 1E Terrestrial Weak Interaction Rates  $(s^{-1})$ 

TABLE 1F		
FERRESTRIAL WEAK INTERACTION RATES	$(s^{-1})$	)

Parent	Daughter	Decay	$Q_n$ (MeV)	Log Rate
(1)	(2)	(3)	(4)	(5)
<sup>49</sup> K	<sup>49</sup> Ca	β <sup>-</sup>	14 038	$\leq -0.460$
<sup>49</sup> Ca	49Sc	$\beta^{-}$	5 779	-2.878
49Sc	<sup>49</sup> Ti	$\beta^{-}$	2.515	-3.693
<sup>49</sup> V	<sup>49</sup> Ti	۳ ۶	0.0908	-7.614
<sup>49</sup> Cr	<sup>49</sup> V	$\tilde{B}^+ \epsilon^-$	2,117	-3560
<sup>49</sup> Mn	<sup>49</sup> Cr	$\beta^+$	7 205	(+0.168)
<sup>49</sup> Fe	$^{49}Mn^{48}Cr$	$\beta^{+}$	12.632	+0.996
<sup>50</sup> K	<sup>50</sup> Ca	$\beta^{P}$	16 511	$\approx \pm 0.364$
<sup>50</sup> Ca	50 Sc	$\beta^{-}$	5 478	-1.305
<sup>50</sup> Sc	<sup>50</sup> Ti	$\beta^{-}$	7 404	-2170
<sup>50</sup> Sc <sup>m</sup>	50 Sc	и ТТ	0.257	+0.297
<sup>50</sup> V	<sup>50</sup> Ti	$\hat{\beta}^+ e^-$	1 702	$(\mathbf{HI} \mathbf{FR} \mathbf{R})$
<sup>50</sup> V	<sup>50</sup> Cr	$\beta^{-}$	1.550	(HI FRB)
<sup>50</sup> Mn	<sup>50</sup> Cr	$\beta^{+}$	7 121	+0.389
$50 \text{ Mn}^m$	<sup>50</sup> Cr	$\beta^{+}$	7 351	-2.178
<sup>51</sup> Sc	<sup>51</sup> Ti	$\beta^{-}$	7.024	-1253
<sup>51</sup> Ti	51 V	$\beta^{-}$	2.977	-2.701
<sup>51</sup> Cr	51V	ب ج	0.240	-6538
<sup>51</sup> Mn	<sup>51</sup> Cr	$\tilde{B}^+ \epsilon^-$	2 697	-3.602
<sup>52</sup> Ti	<sup>52</sup> V	$\beta^{-}$	2.481	-2.168
<sup>52</sup> V	<sup>52</sup> Cr	$\beta^{-}$	4 488	-2.513
<sup>52</sup> Mn	52 Cr	<i>۳</i>	4 200	- 5 986
1 <b>4111</b>	CI	$\tilde{B}^+$	4 200	-6.396
		SUM	1.200	-5.843
$52 \text{ Mn}^m$	<sup>52</sup> Cr	$\beta^+\epsilon^-$	4 578	-3.269
1 <b>411</b>	<sup>52</sup> Mn	μ ΤΤ	0.378	-5.019
	10III	SUM	0.570	-3.262
<sup>52</sup> Fe	<sup>52</sup> Mn	$B^+$	1 861	-4877
10	IVIII	<i>ب</i>	1.861	- 5 000
		SUM	1.001	-4633
<sup>53</sup> Ti	<sup>53</sup> V	$R^{-}$	5 484	-1.678
<sup>53</sup> V	53 Cr	$\beta^{-}$	3 932	-2141
<sup>53</sup> Mn	53 Cr	<u>م</u>	0.0853	-14226
<sup>53</sup> Fe	<sup>53</sup> Mn	$\beta^+ e^-$	3 232	
$5^{3}$ Fe <sup>m</sup>	<sup>53</sup> Fe		3 041	-2 340
<sup>53</sup> Co	53 Fe	$R^+$	7 793	+0.426
$5^{3}Co^{m}$	53 Fe	$\beta^{\mu}_{R^+}$	10 983	+0.436
	52 Fe	ק	1 597	-1 381
	10	SUM	1.071	+0.443
<sup>53</sup> Ni	<sup>53</sup> Co	$R^+n$	12 719	+1.142
141	0	РР	12.717	1 1.142

Parent	Daughter	Decay	$Q_n$ (MeV)	Log Rate
(1)	(2)	(3)	(4)	(5)
<sup>54</sup> V	<sup>54</sup> Cr	$\beta^{-}$	+7.512	-1.793
<sup>54</sup> Mn	<sup>54</sup> Cr	ε	+0.866	-7.590
<sup>54</sup> Mn	<sup>54</sup> Fe	$\beta^{-}$	+1.208	(HI FRB)
<sup>54</sup> Co	<sup>54</sup> Fe	$\beta^+$	+7.731	+0.555
$54^{54}^{60}^{m}$	<sup>54</sup> Fe	$\beta^+$	+7.930	-2.102
<sup>55</sup> Ti	<sup>55</sup> V	'β <sup>-</sup>	+9.151	(+0.130)
<sup>55</sup> V	<sup>55</sup> Cr	'β <sup>-</sup>	+6.607	(-0.486)
<sup>55</sup> Cr	<sup>55</sup> Mn	'β <sup>-</sup>	+3.115	-2.488
<sup>55</sup> Fe	<sup>55</sup> Mn	ε_	-0.280	-8.090
<sup>55</sup> Co	<sup>55</sup> Fe	$\beta^+$	+2.944	+5.072
		ε	+2.944	-5.597
		SUM		-4.959
<sup>56</sup> Sc	<sup>56</sup> Ti	$\beta^{-}$	+16.251	(+0.957)
<sup>56</sup> Ti	<sup>56</sup> V	B <sup></sup>	+6.801	(+0  EST)
<sup>56</sup> V	<sup>56</sup> Cr	B <sup>-</sup>	+9.566	(+0.505)
<sup>56</sup> Cr	<sup>56</sup> Mn	β <sup>-</sup>	+2.154	-2.708
<sup>56</sup> Mn	<sup>56</sup> Fe	8-	+4.206	-4.127
<sup>56</sup> Co	<sup>56</sup> Fe	ε	+4.057	-7.084
		$\tilde{B}^+$	+4.057	-7.713
		SUM		-6.992
<sup>56</sup> Ni	<sup>56</sup> Co	£	+1.625	-5.881
<sup>57</sup> Ti	57 V	<i>B</i> <sup>-</sup>	+11.232	(+0.704)
<sup>57</sup> V	<sup>57</sup> Cr	<i>β</i> <sup>-</sup>	+8.440	(-0.528)
<sup>57</sup> Cr	<sup>57</sup> Mn	Γ́β∸	+5.208	(-0.439)
<sup>57</sup> Mn	<sup>57</sup> Fe	<i>β</i> <sup>-</sup>	+3.203	-2.141
<sup>57</sup> Co	<sup>57</sup> Fe	ε <sup>-</sup>	+0.316	-7.529
<sup>57</sup> Ni	57Co	ε <sup>-</sup>	+2.755	- 5.494
		$\tilde{B}^+$	+2.755	-5.670
		SUM		-5.272
<sup>57</sup> Cu	<sup>57</sup> Ni	$\tilde{\beta}^+ \epsilon^-$	+7.968	+0.586
<sup>57</sup> Zn	<sup>57</sup> Cu <sup>56</sup> Ni	$\beta^+ p$	+14.479	+1.239
<sup>58</sup> Ti	<sup>58</sup> V	'B <sup>-</sup>	+8.611	(+0.239)
<sup>58</sup> V	<sup>58</sup> Cr	'β <sup>-</sup>	+12.911	(+0.176)
<sup>58</sup> Cr	<sup>58</sup> Mn	$B^{-}$	+4.671	(-0.904)
<sup>58</sup> Mn	<sup>58</sup> Fe	B <sup>-</sup>	+6.831	-0.636
<sup>58</sup> Mn <sup>m</sup>	<sup>58</sup> Fe	$\dot{B}^{-}$	+6.861	-1.972
<sup>58</sup> Co	<sup>58</sup> Fe	ε <sup></sup>	+1.797	-7.016
		$\ddot{B}^+$	+1.797	-7.770
		SUM		-6.946
<sup>58</sup> Co	<sup>58</sup> Ni	$B^{-}$	+0.891	(HI FRB)
<sup>58</sup> Co <sup>m</sup>	<sup>58</sup> Co	ΪT	+0.0249	-4.679
58 Ču	<sup>58</sup> Ni	$\hat{B}^+$	+8.052	(-0.664)
		٣		( 3.301)

cupation of the K-shell increases from zero to one as the density rises to a critical value. Following Iben, Kalata, and Schwartz (1967), this critical density *neglecting screening* is given by

$$\left(\frac{\rho}{\mu_{e}}\right)_{c} = 2.536 \times 10^{5} T_{9}^{3/2} \times \exp\left[-1.58 \times 10^{-4} (Z - 0.31)^{2} / T_{9}\right] \text{g cm}^{-3},$$
(3)

where Z = 0.31 is the effective charge of the capturing nucleus. For example, for  ${}^{49}_{23}\text{V} + e^- \rightarrow {}^{49}_{22}\text{Ti} + \nu_e$  one finds  $(\rho/\mu_e)_c = 0.07 \text{ g cm}^{-3}$  for  $T_9 = 0.01$ ,  $3.55 \times 10^3 \text{ g cm}^{-3}$ for  $T_9 = 0.1$ , and  $2.34 \times 10^5 \text{ g cm}^{-3}$  for  $T_9 = 1$ . Below these critical densities the total capture rate is given by the continuum capture rate,  $\lambda_C$ , alone, while above these critical densities the total rate is  $\lambda_C + \lambda_K$ . The transition is not an abrupt one since  $\lambda_C \sim \lambda_K$  near the critical transition density and  $\lambda_C > \lambda_K$  above this density. For somewhat more accurate calculations a monotonically increasing function of  $y = \log (\rho/\mu_e)/(\rho/\mu_e)_c$  can be used passing through  $\lambda_C + \lambda_K/3$  at  $y \approx -0.3$ ,  $\lambda_C + \lambda_K/2$ at y = 0, and  $\lambda_C + 2\lambda_K/3$  at  $y \approx +0.3$ .

The first column in Table 1 gives the parent nucleus, while the second column gives the daughter. In the stellar rate computations both forward and reverse rates are calculated, and the labels "parent" and "daughter" can alternately apply to both nuclei linked by weak transitions. For example, in the laboratory <sup>56</sup>Mn decays

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TERRESTRIAL WEAK INTERACTION RATES $(S^{-1})$					
Parent (1)	Daughter (2)	Decay (3)	$Q_n$ (MeV) (4)	Log Rate (5)	
<sup>59</sup> V	<sup>59</sup> Cr	β-	10.161	(+0.005)	
<sup>59</sup> Cr	<sup>59</sup> Mn	β-	9.011	(-0.002)	
<sup>59</sup> Mn	<sup>59</sup> Fe	β-	5.692	(-1.012)	
<sup>59</sup> Fe	<sup>59</sup> Co	β-	2.076	-6.745	
<sup>59</sup> Ni	<sup>59</sup> Co	$\epsilon^-$	0.562	-12.533	
		$\beta^+$	0.562	-19.357	
		SUM		-12.533	
<sup>59</sup> Cu	<sup>59</sup> Ni	$\beta^+$	4.290	-2.073	
<sup>60</sup> Ti	$^{60}V$	$\beta^{-}$	10.721	(+0.601)	
<sup>60</sup> V	<sup>60</sup> Cr	β-	14.911	(+1  EST)	
<sup>60</sup> Cr	<sup>60</sup> Mn	β-	5.771	(-0.802)	
<sup>60</sup> Mn	<sup>60</sup> Fe	β-	10.228	(+0  EST)	
<sup>60</sup> Fe	<sup>60</sup> Co	$\beta^{-}$	0.720	-13.135	
<sup>60</sup> Co	<sup>60</sup> Ni	β-	3.335	-8.380	
${}^{60}$ Co <sup><i>m</i></sup>	<sup>60</sup> Co	ĪT	0.0586	-2.960	
	<sup>60</sup> Ni	$\beta^{-}$	3.394	-5.561	
		SUM		-2.959	
<sup>60</sup> Cu	<sup>60</sup> Ni	$\beta^+$	5.616	-3.338	
		ε	5.616	-4.461	
		SUM		-3.307	
<sup>60</sup> Zn	<sup>60</sup> Cu	$\beta^+$	3.648	-2.331	
		ε	3.648	-3.840	
		SUM		-2.318	

TABLE 1G

by electron emission to <sup>56</sup>Fe, and this terrestrial decay rate is listed in Table 1, but at high density and temperature <sup>56</sup>Fe  $\rightarrow$  <sup>56</sup>Mn proceeds via continuum electron capture at high electron Fermi energy and/or via positron emission through thermally populated <sup>56</sup>Fe excited states. The stellar rates for both <sup>56</sup>Fe  $\rightarrow$  <sup>56</sup>Mn and <sup>56</sup>Mn  $\rightarrow$  <sup>56</sup>Fe are computed. The third column in Table 1 gives the terrestrial decay mode, either  $\beta^-$  for electron emission,  $\beta^+$  for positron emission,  $\varepsilon^-$  for electron capture, or IT for  $\gamma$ -ray transition from a nuclear isomeric state (listed separately with an "m", e.g., <sup>24</sup>Na<sup>m</sup>) to the ground state. In addition, the sum of the rates for

TABLE 2 Temperature and Density Points at Which Table 3 Rates Are Evaluated (with total

		• • • •	
<b>D</b>	$T_9$	$\log(\rho/\mu_e)$	$W_F$
Point	$(10^{\circ} \text{ K})$	(g cm <sup>3</sup> )	(MeV)
1	1	3	0.046
2	3	3	0.000
3	10	3	0.000
4	100	3	0.000
5	1	7	1.200
6	3	7	1.021
7	10	7	0.196
8	100	7	0.002
9	1	11	23.930
10	3	11	23.925
11	10	11	23.833
12	100	11	14.523

the various laboratory decay processes is denoted as "SUM." Where it is known that proton, neutron, or  $\alpha$ -particle emission from the daughter follows a weak decay process, then that process is designated as, for example,  $\beta^+ p$ ,  $\beta^- n$ ,  $\beta^+ \alpha$ , respectively. In these cases the final nucleus resulting after particle emission from the daughter is listed after the daughter nucleus in column (2). For example, <sup>21</sup>Mg decays by positron emission to <sup>21</sup>Na which subsequently decays by proton emission to <sup>20</sup>Ne.

As indicated in equations (I-3a) and (I-3b), we have found it far simpler to use total particle energies, restmass plus kinetic rather than kinetic, and to use nuclear Q-values rather than atomic. Thus, column (4) in Table 1 lists the nuclear Q-value,  $Q_n$ , for each nuclear transition, where

$$Q_n = M_p c^2 - M_d c^2, \qquad (4)$$

with  $M_p$  the nuclear mass of the parent ground state and  $M_d$  that of the daughter ground state. The dominant allowed transitions may involve excited states in the daughter nucleus. Differential atomic binding energies are negligible and have not been included. In the emission of both positrons and electrons the maximum kinetic energy is given by  $Q_n - m_e c^2 = Q_n - 0.511$  MeV, and in the capture of positrons and electrons the minimum energy of the antineutrinos or neutrinos is  $Q_n + m_e c^2 =$  $Q_n$  +0.511 MeV. On the other hand, the tabulated Fermi energies,  $U_F$ , are kinetic energies but can easily be converted to total energies,  $W_{\rm F} = U_{\rm F} + m_{\rm e}c^2 = U_{\rm F} +$ 0.511 MeV. Finally, column (5) lists the logarithms of the indicated rates. Those entries in parentheses are cases where no terrestrial decay rate is known and the lowest temperature and density stellar rate from our calculations or an estimate (EST) is provided instead. Exoergic transitions, which are highly forbidden so that the "parent" nucleus is essentially stable, e.g., <sup>50</sup>V, are indicated by (HI FRB). At the lowest temperature,  $T_9 = 0.01$ , excited parent states are almost never appreciably populated, while at the lowest density,  $\rho/\mu_e = 10$ g cm $^{-3}$ , the continuum electron density is quite low, and, except for electron capture, the stellar rates should be quite close to the terrestrial values. The exception noted arises because only continuum electron capture and not bound state capture is calculated.

It transpires that the rates calculated for the lowest temperature-density points and the terrestrial rates often disagree. This is because what experimentalists actually measure are branching intensities, not matrix elements. To convert branching intensities for various transitions into  $(\log ft)$ -values, equivalent to matrix elements, requires two input quantities: the overall lifetime of the decaying nucleus, and the *f*-factors (phase space factors) for each of the individual transitions. There is no difficulty with the lifetimes, which are commonly very accurately known, but the appropriate *f*-factors are usually complicated integrals involving electron and neutrino

1982ApJS...48..279F

TABLE 3	STELLAR WEAK RATES EVALUATED AT TABLE 2 TEMPERATURE AND DENSITY POINTS
---------	--

s. Ne	- 990,000,000,000,000,000 888,890,000,000,000,000,000,000,000,000,	2 	27 M M M M M M M M M M M M M M M M M M M
↑ 	$\begin{array}{c} & - & - & - & 0 \\ & - & - & - & - \\ & - & - & - & - \\ & - & -$	0 0 0 0 0 0 0 0 0 0 0 0 0 0	<sup>2</sup> N <sup>0</sup> − − − − − − − − − − − − − − − − − − −
5	0 0 0 0 0 0 0 0 0 0 0 0 0 0	2692429200000000000000000000000000000000	
21	$\begin{array}{c} -\frac{1}{2} \\ -1$	0 -1-50 -1-1-50 -1-1-1-1-1-20 -1-1-1-1-1-1-20 -1-1-1-1-1-1-20 -1-1-1-1-1-1-20 -1-1-1-1-1-1-20 -1-1-1-1-1-1-20 -1-1-1-1-1-20 -1-1-1-1-20 -1-1-1-20 -1-1-1-20 -1-1-1-20 -1-1-1-20 -1-1-20 -1-1-20 -1-1-20 -1-1-20 -1-1-20 -1-1-20 -1-1-20 -1	22 N - 0-0-0-0-0-0-0-0-0-0-0-0-0-0-0-0-0-0-
¢ ₽ Z		– 41– – 41– – 86– –	°°°°°°°°°°°°°°°°°°°°°°°°°°°°°°°°°°°°°
31	- 4-0 - 4-0 - 4-0 - 1-5. 938 - 4-0 - 1-5. 938 - 4-0 - 1-5. 938 - 1-5. 938 - 1-5. 938 - 1-5. 938 - 1-5. 938 - 1-5. 938 - 1-6. 0 -		2.000000000000000000000000000000000000
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'Na → <sup>25</sup> Mg		Ne → <sup>25</sup> Na		Al → <sup>26</sup> Si	
<sup>25</sup> Na → <sup>25</sup> Mg	$ \begin{array}{c} \begin{array}{c} \begin{array}{c} \ \ \ \ \ \ \ \ \ \ \ \ \ \ \ \ \ \ \$	<sup>25</sup> Ne → <sup>25</sup> No	1.57 $1.53$	<sup>26</sup> AI → <sup>26</sup> Si	$-69^{-6}$ $-69^{-6}$ $-69^{-6}$ $-69^{-6}$ $-69^{-6}$ $-69^{-6}$ $-15^{\circ}$ 28.56 $-27^{\circ}$ 835 $-28^{\circ}$ 5.56 $-15^{\circ}$ 28.46 $-35^{\circ}$ 28.46 $-35^{\circ}$ 28.56 $-36^{\circ}$ 28.46 $-36^{\circ}$ 28.56 $-36^{\circ}$ 28.5
<sup>25</sup> Na → <sup>25</sup> Ma	$ \begin{array}{c} 100 \ \nu \\ -26.940 \ -1.785 \ -6.738 \ -1.421 \ -26.940 \ -1.785 \ -6.738 \ -1.421 \ -2.632 \ -2.607 \ -1.723 \ -2.672 \ -1.421 \ -1.428 \ -0.431 \ -2.600 \ -1.777 \ -5.376 \ -1.500 \ -3.558 \ -1.650 \ -2.600 \ -1.777 \ -5.376 \ -1.500 \ -2.600 \ -1.777 \ -5.376 \ -1.500 \ -2.600 \ -2.558 \ -1.650 \ -2.558 \ -1.650 \ -2.558 \ -1.650 \ -2.558 \ -1.650 \ -2.558 \ -1.500 \ -2.558 \ -1.500 \ -2.558 \ -1.500 \ -2.558 \ -1.500 \ -2.558 \ -1.500 \ -2.558 \ -1.500 \ -2.558 \ -1.500 \ -2.558 \ -1.550 \ -2.558 \ -1.550 \ -2.558 \ -1.550 \ -2.558 \ -1.550 \ -2.558 \ -1.550 \ -2.558 \ -1.550 \ -2.558 \ -1.550 \ -2.558 \ -$	25Ne 25Ne 25Na 1	4.3.4         0.074         5.801         0.094           -16.023         0.074         -5.801         0.63           -4.105         1.056         -0.444         1.82           -37.328         0.075         -5.801         0.63           -37.828         0.075         -5.801         0.63           -37.828         0.075         -5.801         0.63           -37.828         0.061         -11.621         0.61           -4.007         1.054         -0.537         1.81           -4.007         1.054         -0.537         1.81           -4.983         2.348         5.009         6.67           -5.032         -7.063         -12.426         0.61           -5.032         -7.063         -12.426         0.61           -6.033         2.348         5.009         66.67           -6.033         2.348         5.003         66.67           -5.032         -7.063         -12.426         66.73           -5.032         -12.426         -6.73         6.95           -6.057         4.288         5.95         9.95	<sup>26</sup> AI → <sup>26</sup> Si	$ \begin{array}{c} 109 \ \nu \\ -0.137 \ -35.380 \ -27.835 \ -28.46 \ -0.137 \ -35.380 \ -27.835 \ -28.46 \ -0.111 \ -15.228 \ -10.826 \ -10.826 \ -10.826 \ -10.826 \ -10.826 \ -10.826 \ -10.826 \ -10.826 \ -10.826 \ -10.826 \ -3.556 \ -3.556 \ -3.566$
Mg → <sup>25</sup> Ng → <sup>25</sup> Mg	$\begin{array}{c ccccccccccccccccccccccccccccccccccc$	$Na \rightarrow 2^5 Ne \qquad 2^5 Ne \rightarrow 2^5 Na$	-4.307       -4.304       0.074       -0.94       0.094         -15.3072       -4.304       0.075       -5.801       0.63         -15.305       -16.023       0.075       -2.833       0.63         -4.603       -4.105       1.056       -0.444       1.82         -37.252       -14.307       0.061       -11.621       0.61         -4.564       -4.007       1.054       -0.537       1.81         -4.554       -4.007       1.054       -0.537       1.81         -4.554       -4.007       1.054       -0.537       1.81         -4.564       -4.007       1.054       -0.537       1.81         -4.564       -4.007       1.054       -0.537       1.81         -4.564       -4.007       1.054       -0.537       1.81         -4.566       -893       -83.33       -83.48       5.009       6.67         -5.161       6.093       2.348       5.009       6.67       -1.81         -4.564       -6.033       2.348       5.009       6.67       -1.81         -5.161       6.093       2.348       -0.537       1.81       -1.81         -4.564       -6.03	Si → <sup>26</sup> AI → <sup>26</sup> Si	$\begin{array}{c ccccccccccccccccccccccccccccccccccc$
<sup>25</sup> Mg → <sup>25</sup> Na → <sup>25</sup> Mg	$ \begin{array}{c ccccccccccccccccccccccccccccccccccc$	$25 \text{Ng} \rightarrow 25 \text{Ne} \qquad 25 \text{Ne} \rightarrow 25 \text{Ne} \qquad 25 \text{Ne} \rightarrow 100 \text{ k}^{-1} + $	<b>-5.291 -5.201 -5</b>	<sup>26</sup> Si → <sup>26</sup> AI <sup>26</sup> Si → <sup>26</sup> Si	$ \begin{array}{c ccccccccccccccccccccccccccccccccccc$

289

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<sup>7</sup> Si		<sup>27</sup> AI	0,0,0,0,0,0,0,0,0,0,0,0,0,0,0,0,0,0,0,	<sup>27</sup> Mg	-7
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<sup>27</sup> AI	-99009000000	<sup>7</sup> Ma		27 <sub>N</sub>	
'Si → <sup>27</sup> AI	- - - - - - - - - - - - - - - - - - -	AI -> <sup>27</sup> Ma	-202	Mq → <sup>27</sup> N	
<sup>27</sup> Si → <sup>27</sup> AI	- 000000000000000000000000000000000000	<sup>27</sup> AI $\rightarrow$ <sup>27</sup> Ma	$ \begin{array}{c} \begin{array}{c} -26 \\ -26 \\ -26 \\ -26 \\ -26 \\ -26 \\ -26 \\ -26 \\ -26 \\ -26 \\ -26 \\ -26 \\ -26 \\ -28 \\ -22 \\ -27 \\ -28 \\ -28 \\ -73 \\ -28 \\ -73 \\ -12 \\ -28 \\ -28 \\ -73 \\ -28 \\ -73 \\ -12 \\ -28 \\ -28 \\ -73 \\ -28 \\ -73 \\ -28 \\ -73 \\ -28 \\ -73 \\ -28 \\ -73 \\ -28 \\ -73 \\ -28 \\ -73 \\ -28 \\ -73 \\ -12 \\ -89 \\ -28 \\ -78 \\ -28 \\ -28 \\ -78 \\ -28 \\$	<sup>27</sup> Mg → <sup>27</sup> N	$\left[ \begin{array}{cccccccccccccccccccccccccccccccccccc$

290

29p		S N N N N N N N N N N N N N	28 Al 0.00 0.1388 0.1388 0.1388 0.1388 0.1388 0.1388 0.1388 0.1388 0.138
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<sup>31</sup> Si → <sup>31</sup> P	$ \begin{bmatrix} \log \beta & -\log \beta & \log \epsilon \\ -4.064 & -8.091 & -4.097 \\ -3.902 & -4.846 & -3.777 \\ -2.359 & -1.823 & -1.036 \\ -4.414 & -13.910 & -4.645 \\ -4.418 & -2.363 & -1.920 & -1.108 \\ -2.363 & -1.920 & -1.108 \\ -2.349 & -6.553 & -4.114 \\ -2.349 & -6.553 & -4.114 \\ -2.349 & -6.553 & -4.116 \\ -2.349 & -6.53 & -1.028 \\ -1.0284 & -13.823 & -9.944 \\ -10.284 & -13.823 & -9.944 \\ 1.966 & 4.078 & 5.733 \\ \end{bmatrix} $	$\begin{array}{cccccccccccccccccccccccccccccccccccc$	$\begin{array}{cccccccccccccccccccccccccccccccccccc$
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38 Ar		38 <sub>C1</sub>	-32.607 -32.607 -12.112 -2.389 6.184	-26.902 -10.424 -2.294 6.184	3.664 3.667 4.073 6.879	38 <sub>S</sub>	-23.607 -23.607 -3.321 -3.321 -18.110		0.837 2.817 6.244
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298

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49 <sub>Ti</sub>	- - - - - - - - - - - - - - - - - - -	* * * * * * * * * * * * * *	* * Construction * * * * * * * * * * * * * * * * * * *
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Cr → <sup>49</sup> Mn	$100 \ \epsilon^{+}$ $100 \ \epsilon^{-}$ $100 \ \epsilon^{-}$ $100 \ \epsilon^{-}$ $1314 \ -0.891$ $-3.995 \ -3.411 \ -3.995 \ -3.411 \ -3.995 \ -3.411 \ -4.033 \ -3.509 \ -5.416 \ -15.105 \ -16.115 \ -4.033 \ -5.59 \ -59.549 \ -59.549 \ -59.549 \ -59.549 \ -59.549 \ -59.5416 \ -15.416 \ -15$	$ \begin{array}{cccccccccccccccccccccccccccccccccccc$	Ti $\rightarrow$ <sup>49</sup> -9.326 - 9.547 -9.326 - 9.547 -5.693 - 5.453 -5.693 - 5.453 -2.270 - 1.205 -7.404 - 6.418 -2.366 - 1.247 -2.366 - 1.247 -2.366 - 1.247 -2.366 - 1.247 -1.247 -1.247 -2.366 - 1.247 -1.
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<sup>49</sup> Cr → <sup>49</sup> Mn	$\begin{bmatrix} 109 & \nu & 109 & \beta & 109 & \epsilon^{+} & 109 & \nu \\ 0.730 & -16.323 & -14.386 & -14.461 \\ 1.235 & -6.147 & -3.995 & -3.411 \\ 7.235 & -6.147 & -3.995 & -3.411 \\ 7.05 & -2.336 & -45.134 & -46.550 \\ 0.881 & -16.707 & -16.102 & -16.115 \\ 1.296 & -6.176 & -4.093 & -3.509 \\ 1.296 & -6.176 & -4.093 & -3.509 \\ 1.296 & -5.399 & -9999 & -9999 \\ 5.509 & -54.835 & -54.579 & -54.543 \\ 6.521 & -17.552 & -16.001 & -15.416 \\ 7.707 & -2.793 & 3.777 & 5.386 \\ \end{bmatrix}$	49 $49^{\vee}$ $49^{\vee}$ $49^{\circ}$ 109 $\nu$ $109^{\circ}$ $109^{\circ}$ -3.526       -50.925 $109^{\circ}$ $109^{\circ}$ -2.950 $15.765^{\circ}$ $-7.749^{\circ}$ $-7.73^{\circ}$ -2.955 $-115.765^{\circ}$ $-7.749^{\circ}$ $-7.73^{\circ}$ -2.955 $-15.765^{\circ}$ $-7.749^{\circ}$ $-7.773^{\circ}$ -1.776 $-50.934^{\circ}$ $-22.574^{\circ}$ $-23.127^{\circ}$ -1.776 $-50.934^{\circ}$ $-22.574^{\circ}$ $-2451^{\circ}$ -1.776 $-50.934^{\circ}$ $-22.574^{\circ}$ $-2451^{\circ}$ -1.776 $-50.934^{\circ}$ $-22.574^{\circ}$ $-2451^{\circ}$ -1.776 $-9.465^{\circ}$ $-9.489^{\circ}$ $-9.489^{\circ}$ -1.776 $-9.465^{\circ}$ $-9.489^{\circ}$ $-9.489^{\circ}$ -1.776 $-9.465^{\circ}$ $-9.489^{\circ}$ $-9.489^{\circ}$ -1.776 $-9.465^{\circ}$ $-9.469^{\circ}$ $-9.489^{\circ}$ -1.776 $-9.465^{\circ}$ $-9.469^{\circ}$ $-9.459^{\circ}$ 5.748^{\circ} $-9.465^{\circ}$ $-9.469^{\circ}$ $-9.745^{\circ}$ 5.774 $-47.776^{\circ}$ <t< th=""><th><sup>49</sup> Ti       <sup>49</sup> Ti       <sup>49</sup> V         <math>\log \nu</math> <math>\log \rho</math> <math>\log \rho</math> <math>\log \rho</math> <math>= 0.249</math> <math>-11.519</math> <math>-9.326</math> <math>-9.547</math> <math>= 0.2133</math> <math>-1.233</math> <math>-5.633</math> <math>-5.453</math> <math>= 7.081</math> <math>-12.270</math> <math>-1.205</math> <math>-9.5453</math> <math>= 7.081</math> <math>-12.270</math> <math>-1.205</math> <math>-9.5433</math> <math>= 7.081</math> <math>-12.826</math> <math>-15.145</math> <math>-13.331</math> <math>= 7.081</math> <math>-12.826</math> <math>-12.270</math> <math>-1.247</math> <math>= 7.091</math> <math>-99.999</math> <math>-99.999</math> <math>-99.999</math> <math>= 7.091</math> <math>-99.999</math> <math>-99.999</math> <math>-99.999</math> <math>= 7.410</math> <math>-11.238</math> <math>-14.267</math> <math>-10.937</math> <math>= 7.411</math> <math>-10.737</math> <math>-4.015</math> <math>-5.630</math></th></t<>	<sup>49</sup> Ti <sup>49</sup> Ti <sup>49</sup> V $\log \nu$ $\log \rho$ $\log \rho$ $\log \rho$ $= 0.249$ $-11.519$ $-9.326$ $-9.547$ $= 0.2133$ $-1.233$ $-5.633$ $-5.453$ $= 7.081$ $-12.270$ $-1.205$ $-9.5453$ $= 7.081$ $-12.270$ $-1.205$ $-9.5433$ $= 7.081$ $-12.826$ $-15.145$ $-13.331$ $= 7.081$ $-12.826$ $-12.270$ $-1.247$ $= 7.091$ $-99.999$ $-99.999$ $-99.999$ $= 7.091$ $-99.999$ $-99.999$ $-99.999$ $= 7.410$ $-11.238$ $-14.267$ $-10.937$ $= 7.411$ $-10.737$ $-4.015$ $-5.630$
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<sup>49</sup> Mn → <sup>49</sup> Cr → <sup>49</sup> Mn	$ \begin{bmatrix} 109 & \beta^{+} & 109 & \epsilon^{-} & 109 & \nu & 109 & \beta^{-} & 109 & \epsilon^{+} & 109 & \nu \\ 0.169 & -4.259 & 0.730 & -16.323 & -14.368 & -14.461 \\ 0.171 & -1.964 & 0.730 & -16.323 & -14.368 & -14.461 \\ 0.2219 & 0.122 & 1.235 & -6.147 & -3.995 & -3.411 \\ 0.222 & 0.559 & 0.870 & -45.36 & -46.530 \\ 0.172 & -0.546 & 0.881 & -16.707 & -16.102 & -16.115 \\ 0.222 & 0.216 & 1.296 & -6.176 & -4.093 & -3.509 \\ 0.172 & -0.546 & 0.881 & -16.707 & -16.102 & -16.115 \\ 0.222 & 0.216 & 1.296 & -6.176 & -4.093 & -3.509 \\ 0.172 & 5.216 & 6.509 & -9999 & -9999 & 999 & 999 & 999 & 999 \\ 0.172 & 5.216 & 6.509 & -54.835 & -54.579 & -54.545 \\ 0.233 & 5.221 & 6.521 & -17.562 & -16.001 & -15.416 \\ 2.476 & 6.014 & 7.707 & -2.793 & 3.777 & 5.386 \\ \end{bmatrix} $	$\begin{array}{cccccccccccccccccccccccccccccccccccc$	$\begin{array}{cccccccccccccccccccccccccccccccccccc$

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${}^{50}S_{6} \rightarrow {}^{50}T_{1}$ ${}^{109}f^{-1}O_{9}f^{$	1.826 4.542 6.163 	$ \begin{array}{c} \begin{array}{c} 1 \\ 1 \\ -1 \\ -1 \\ -1 \\ -1 \\ -1 \\ -1 \\ $	${}^{51}Cr \rightarrow {}^{51}Mn$ ${}^{109}f^{51}r \rightarrow {}^{51}Mn$ ${}^{-13}.108 - {}^{80}.559 - {}^{20}.190 \frac{\nu}{\nu}$ ${}^{-13}.108 - {}^{80}.559 - {}^{20}.190 \frac{\nu}{\nu}$ ${}^{-13}.103 - {}^{30}.052 - {}^{20}.190$ ${}^{20}.599 - {}^{4}.531 - {}^{20}.235$ ${}^{-13}.128 - {}^{10}.233 - {}^{20}.235$ ${}^{-13}.128 - {}^{20}.233 - {}^{20}.235$ ${}^{-13}.128 - {}^{20}.235 - {}^{20}.999$ ${}^{45}.176 - {}^{20}.287 - {}^{20}.247$ ${}^{20}.999 - {}^{99}.999 - {}^{99}.999$ ${}^{45}.176 - {}^{15}.068 - {}^{20}.247$ ${}^{21}.176 - {}^{15}.068 - {}^{20}.247$ ${}^{21}.176 - {}^{15}.068 - {}^{20}.999$ ${}^{21}.176 - {}^{21}.068 - {}^{20}.923$ ${}^{21}.176 - {}^{21}.068 - {}^{20}.923$ ${}^{21}.176 - {}^{21}.068 - {}^{21}.923$ ${}^{21}.176 - {}^{21}.068 - {}^{21}.923$ ${}^{21}.161 - {}^{21}.068 - {}^{21}.923$ ${}^{21}.176 - {}^{21}.068 - {}^{21}.923$ ${}^{21}.161 - {}^{21}.061 - {}^{21}.061 - {}^{21}.061$ ${}^{21}.176 - {}^{21}.068 - {}^{21}.923$ ${}^{21}.176 - {}^{21}.068 - {}^{21}.923$ ${}^{21}.161 - {}^{21}.061 - {}^{21}.061 - {}^{21}.061$ ${}^{21}.061 - {}^{21}.$
$\begin{array}{c ccccccccccccccccccccccccccccccccccc$	8         2.078         4.743         6.425           9         -66.792         3.944         4.945           10         -20.975         3.981         5.015           11         -5.230         3.981         5.015           12         2.191         5.429         7.126	pt $109$ $\beta^{+}$ $109$ $\epsilon^{-}$ $109$ $\epsilon^{-}$ $109$ $\epsilon^{-}$ $109$ $\epsilon^{-}$ $103$ $\epsilon^{-}$ $103$ $\epsilon^{-}$ $103$ $\epsilon^{-}$ $103$ $\epsilon^{-}$ $103$ $\epsilon^{-}$ $133$ $279$ $2 -19.452 -13.077 -13.139$ $2 -19.451 -13.689 -27.459$ $6 -19.451 -11.362 -11.424$ $6 -19.451 -11.362 -11.424$ $6 -19.451 -11.362 -11.424$ $6 -19.451 -11.362 -11.424$ $1.853 -3.738$ $8 -57.331 -26.890 -27.459$ $6 -135 -57.331 -26.890 -27.459$ $1 -5.993 -4.259 -3.738$ $1 -5.974 -2.738$ $1 -5.974 -2.738$ $1 -5.974 -2.738$ $1 -5.974 -2.738$ $1 -5.974 -2.736$ $2 -335$ $1 -5.974 -2.738 -2.738 -$	$\begin{array}{c ccccccccccccccccccccccccccccccccccc$
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<sup>52</sup> Mn 52Mn 52Fe	$ \begin{array}{c} \log \nu & \log \beta & \log \epsilon & \log \frac{1}{2} \\ -5.086 & -42.794 & -16.886 & -17.420 \\ -3.386 & -13.867 & -8.569 & -8.664 \\ -3.107 & 0.731 & 4.111 & 6.019 \\ -3.266 & -42.810 & -22.706 & -23.240 \\ -1.986 & -13.881 & -10.304 & -10.319 \\ -1.008 & -13.881 & -10.304 & -10.319 \\ -0.008 & -13.881 & -10.318 & -10.319 \\ -0.008 & -13.881 & -10.318 & -10.319 \\ -0.008 & -13.881 & -10.318 & -10.318 \\ -0.008 & -13.881 & -10.318 \\ -0.008 & -13.881 & -10.318 \\ -0.008 & -13.881 & -10.318 \\ -0.008 & -10.318 & -10.318 \\ -0.008 & -10.318 & -10.318 \\ -0.008 & -10.318 & -10.318 \\ -0.008 & -10.318 & -10.318 \\ -0.008 & -10.318 & -10.318 \\ -0.008 & -10.318 & -10.318 \\ -0.008 & -10.318 & -10.318 \\ -0.008 & -10.318 & -10.318 \\ -0.008 & -10.318 & -$	$ \begin{array}{cccccccccccccccccccccccccccccccccccc$	$\begin{array}{c ccccccccccccccccccccccccccccccccccc$
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<sup>51</sup> V → <sup>51</sup> Cr	$ \begin{array}{c ccccccccccccccccccccccccccccccccccc$	$\begin{array}{c ccccccccccccccccccccccccccccccccccc$	${}^{51}S_{C} \rightarrow {}^{51}T_{1}$ ${}^{109}$ ${}^{61}$ ${}^{-1.261}$ ${}^{-1.261}$ ${}^{-1.261}$ ${}^{-1.173}$ ${}^{-1.263}$ ${}^{-1.449}$ ${}^{-0.931}$ ${}^{-0.254}$ ${}^{-1.449}$ ${}^{-0.510}$ ${}^{-1.213}$ ${}^{-5.339}$ ${}^{-1.213}$ ${}^{-5.339}$ ${}^{-1.233}$ ${}^{-1.233}$ ${}^{-1.233}$ ${}^{-1.233}$ ${}^{-1.233}$ ${}^{-1.233}$ ${}^{-2.352}$ ${}^{-1.233}$ ${}^{-2.352}$ ${}^{-2.352}$ ${}^{-2.339}$ ${}^{-2.352}$ ${}^{-2.352}$ ${}^{-2.352}$ ${}^{-2.339}$ ${}^{-2.352}$ ${}^{-2.352}$ ${}^{-2.352}$ ${}^{-2.369}$ ${}^{-29.9696}$ ${}^{-29.666}$ ${}^{-29.666$
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55Cr → 55Mn	$\begin{array}{c} \begin{array}{c} \begin{array}{c} \ \ \ \ \ \ \ \ \ \ \ \ \ \ \ \ \ \ \$	$55 \lor 55 $	$\begin{array}{c} 5_{11} \longrightarrow 55_{1} \ 100 \ 130 \$
<sup>55</sup> Cr → <sup>55</sup> Mn	$\begin{array}{c} \begin{array}{c} -09 \\ -21.474 \\ -2.486 \\ -21.847 \\ -2.440 \\ -7.460 \\ -7.460 \\ -7.188 \\ -2.188 \\ -4.111 \\ -1.565 \\ -1.072 \\ -1.208 \\ -1.208 \\ -1.175 \\ -1.175 \\ -1.175 \\ -1.175 \\ -1.175 \\ -1.175 \\ -1.175 \\ -1.175 \\ -1.175 \\ -1.175 \\ -1.175 \\ -1.130 \\ -1.175 \\ -1.130 \\ -$	$5^{\circ}$ $5^{\circ}$ $5^{\circ}$ $109^{\circ}$ $109^{\circ}$ $109^{\circ}$ $109^{\circ}$ $109^{\circ}$ $109^{\circ}$ 133000 133000 133000 133000 133000 133000 1330000 13300000 13300000000000000000000000000000000000	$\begin{array}{c ccccccccccccccccccccccccccccccccccc$
Wn → <sup>55</sup> Cr → <sup>55</sup> Mn	$\begin{array}{c ccccccccccccccccccccccccccccccccccc$	Cr $\rightarrow$ 55 V 55 V 55 V $\rightarrow$ 55 Cr -100 $\epsilon^{-1}$ 100 $\epsilon^{-1}$ 100 $\epsilon^{-1}$ 100 $\epsilon^{-1}$ 100 -14.310 -14.390 -0.236 -6.498 0 -14.310 -14.390 -0.236 -6.498 0 -13.368 -34.262 -0.386 -12.319 0 -12.554 -12.674 -0.314 -5.101 0 -13.761 -5.100 -86 926 -99 999 -67 4.157 5.120 -86 926 -99 999 -67 4.155 5.120 -86 926 -99 999 -67 4.157 5.120 -86 926 -99 999 -79 -13.173 -70 -2.386 -2.286 -29 -29 -2.386 -2.286 -20 -2.489 -2.286 -29 -2.296 -20 -2.489 -2.286 -29 -2.296 -29 -2.296 -20 -2.489 -2.286 -29 -2.296 -20 -2.489 -2.286 -29 -2.296 -20 -2.296 -20 -2.296 -20 -2.296 -20 -2.489 -2.286 -20 -2.480 -2.480 -2.480 -2.480 -2.480 -2.586 -20 -2.296 -2.20 -2.296 -2.20 -2.296 -2.20 -2.296 -2.20 -2.296 -2.20 -2.206 -2.200 -2	V → $5_{11}$ $5_{11$
<sup>55</sup> Mn → <sup>55</sup> Cr → <sup>55</sup> Mn	$ \begin{array}{c ccccccccccccccccccccccccccccccccccc$	$\begin{array}{c ccccccccccccccccccccccccccccccccccc$	<b>55</b> $J_{51} \rightarrow 55_{T1}$ <b>55</b> $J_{51} \rightarrow 55_{T1} \rightarrow 55_{T1} \rightarrow 55_{T1} \rightarrow 55_{T1} \rightarrow 55_{T1} \rightarrow 55_{T1} \rightarrow 55_{T2}$ <b>-79.743 -50.613 -51.192 0.130 -6.320 0.97</b> <b>-7.755 -4.565 -4.099 s</b> <sup>-1</sup> <b>109 s</b> <sup>+</sup> <b>109</b> <b>-7.755 -4.563 -6.320 0.97</b> <b>-7.755 -4.563 -6.329 0.91</b> <b>-7.755 -4.563 -6.329 0.91</b> <b>-7.754 -18.615 -18.702 0.317 -3.209 0.91</b> <b>-7.754 -18.615 -18.702 0.317 -3.209 0.97</b> <b>-7.754 -16.699 -16.987 0.310 -4.914 0.9</b> <b>-7.754 -4.466 -4.000 -2.558 4.956 6.65</b> <b>-79.743 4.146 5.156 -75.931 -99.999 -76.57</b> <b>-7.749 5.161 -24.791 -43.391 -24.93</b> <b>-7.749 5.161 -24.791 -43.391 -24.93</b> <b>-7.749 5.161 -24.791 -43.361 -54.65</b> <b>-7.749 5.161 -24.791 -43.361 -54.95</b> <b>-7.749 5.161 -24.791 -43.361 -54.65</b> <b>-7.749 5.161 -24.791 -43.361 -54.95</b> <b>-7.749 5.161 -24.791 -44.236 -56</b> <b>5.161 -24.791 -24.791 -24.9</b>

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↑ F		<sup>5</sup> C <sup>-</sup> -58.878 -58.878 -58.878 -5.253 -5.253 -5.253 -5.253 -5.253 -5.253 -5.253 -5.253 -5.253 -5.253 -5.254 -5.254 -5.254 -5.254 -5.254 -5.254 -5.254 -5.254 -5.254 -5.254 -5.255 -5.254 -5.255 -5.155 -5.255 -	× 100 ° ° ° ° ° ° ° ° ° ° ° ° ° ° ° ° ° °
8		-28.1412 -28.142 -28.142 -7.964 -7.964 -7.963 -7.960 -7.960 -7.960 -7.960 -28.1412 -28	
1	g-00400000000000	2-00400080010	1-00400080010

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	1-0040000000000	g-0040000005-0	g-004000000555
7Ni	$\begin{array}{c} 109 \ \overline{\nu} \\ -211 \ 729 \\ -2.013 \\ -2.013 \\ -2.105 \\ -2.105 \\ -2.105 \\ -2.105 \\ -2.105 \\ -3.430 \\ -3.0641 \\ -3.52$	<sup>7</sup> Co Co − − − − − − − − − − − − − − − − − − −	<sup>37</sup> Fe 
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	58 <sub>Mn</sub>		<sup>8</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup> <sup>-1</sup>	6.918 -53.730 -5.685 -5.685 6.918 7.612 7.612
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8 \	$\begin{array}{c} \textbf{-0} & \textbf{-0} \\ \textbf{-0} & \textbf{-1} \\ \textbf{-1} &$	22	-11.900 5.612
<sup>1</sup> Ti → <sup>58</sup> V	log         t         log         t           -6.290         0.875         -3.186         1.032           -3.186         1.032         0.875         -3.13           -3.13         2.846         1.032         -3.46           -12.103         0.860         -1.020         -3.46           -4.890         1.020         0.860         -4.890           -4.3368         -5.954         -5.954         -5.954           -5.954         -5.954         5.966         -5.954	Ni → $5^{9}$ Cu → $5^{9}$ Cu → $5^{9}$ Cu → $1-09 \epsilon^{+}$ + $1-06 \epsilon^{-}$ + $1-26.312 - 27.478$ - $1-3.2.328$ - $2.328$ - $2.328$ - $2.328$ - $2.328$ - $2.329$ - $2.339$ - $2.3$	-14.012 -11.900 3.991 5.612
58 <sub>Ti</sub> → <sup>58</sup> √	$ \begin{array}{c ccccccccccccccccccccccccccccccccccc$	$\begin{array}{c ccccccccccccccccccccccccccccccccccc$	-12.126 -14.012 -11.900 0.225 3.991 5.612
<sup>58</sup> Ti → <sup>58</sup> √	log v       log v       log v       log v       log v         -48.774       0.239       -6.290       0.875         -18.105       0.370       -3.186       1.032         -4.158       2.107       0.313       2.846         -4.158       2.107       0.313       2.846         -4.158       2.107       0.313       2.846         -4.5955       0.231       -12.109       0.860         -16.389       0.364       -4.890       1.020         -4.602       2.106       0.226       2.840         -4.918       -78.656       -99.939       -9.256         -5.931       -11.668       -5.954       -5.954         -5.931       2.106       0.226       2.840         -5.991       2.866       -93.935       2.954         -5.991       2.708       -11.668       -5.954         -5.954       6.638       2.086       -5.954         -6.638       2.086       4.314       5.966	59       59       59       59 $109$ $59$ $109$ $59$ $109$ $7$ $109$ $7$ $109$ $7$ $109$ $7$ $109$ $7$ $109$ $7$ $109$ $7$ $109$ $7$ $109$ $7$ $100$ <th< td=""><td><b>6.311</b> -12.126 -14.012 -11.900 7.<b>669</b> 0.225 3.991 5.612</td></th<>	<b>6.311</b> -12.126 -14.012 -11.900 7. <b>669</b> 0.225 3.991 5.612
<sup>1</sup> V → <sup>58</sup> Ti → <sup>58</sup> V	$ \begin{array}{c ccccccccccccccccccccccccccccccccccc$	$ \begin{array}{c ccccccccccccccccccccccccccccccccccc$	5.102 6.311 -12.126 -14.012 -11.900 5.979 7.669 0.225 3.991 5.612
36 ↓ 58 Ti → 58 Ti → 58 V	$ \begin{array}{c ccccccccccccccccccccccccccccccccccc$	$ \begin{array}{cccccccccccccccccccccccccccccccccccc$	-2.680 5.102 6.311 -12.126 -14.012 -11.900 2.920 5.979 7.669 0.225 3.991 5.612

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316

3Continued	$\begin{array}{c ccccccccccccccccccccccccccccccccccc$	$\begin{array}{cccccccccccccccccccccccccccccccccccc$	$ \begin{array}{cccccccccccccccccccccccccccccccccccc$	$\varepsilon^+$ are in s <sup>-1</sup> . Neutrino energy loss Base 10 logarithms are presented.
TABLE	$\begin{array}{c ccccccccccccccccccccccccccccccccccc$	$\begin{array}{cccccccccccccccccccccccccccccccccccc$	$ \begin{array}{cccccccccccccccccccccccccccccccccccc$	NOTE.—Rates $\beta^+$ , $\epsilon^-$ , $\beta^-$ , rates $\nu$ and $\bar{\nu}$ are in MeV s <sup>-1</sup> .

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318

energies and the Coulomb correction factors (cf. eqs. [I-3a], [I-3b], [I-5b] and Gove and Martin 1971, p. 208) which require approximations in actual calculations. Some authors use the tables of  $\log f$  by Gove and Martin (1971), others use approximate f-factor formulae; some authors use f-factors appropriate for first forbidden transitions (denoted with subscript 1 in Lederer et al. 1978), while others use only allowed phase space factors. In short, the  $(\log ft)$ -values tabulated in Endt and van der Leun (1978) and Lederer et al. (1978) are in some cases calculated by methods unknown to us. The f-factors in our stellar rate program are calculated as accurately as possible, using detailed numerical integrations, but these are not applicable if the published  $(\log ft)$ -values have been calculated using an unknown approximation to f.

In order to resolve this problem insofar as possible,  $(\log ft)$ -values were recalculated for cases where the appropriate lowest temperature-density point stellar rate differed from the terrestrial rate by more than 30%  $(\Delta \log \ge 0.1)$ . The lowest temperature-density point stellar rates presented on the magnetic tape and in this paper differ from the terrestrial decay rates, where known, by no more than 30%, except, as noted, in the case of electron capture.

An example of a reaction where the  $(\log ft)$ -values were adjusted in this manner is the  ${}^{51}Sc \rightarrow {}^{51}Ti$  electron emission transition. Several  $(\log ft)$ -values were measured for this and presented in Lederer et al. (1980): from the <sup>51</sup>Sc ground state  $(J^{\pi} = 7/2^{-})$  to state 3  $(7/2^{-})$  of <sup>51</sup>Ti, log ft = 5.5; to state 5  $(5/2^{-})$ , log ft =5.0. In addition, the ground state to state 4  $(5/2^{-})$ transition is clearly allowed, yet no branch was observed, so that this transition was assigned  $\log ft = 99.9$ . When the lowest temperature-density point stellar rate calculation was performed, the electron emission result was  $\log \beta^- = 1.059$ , ~ 50% faster than the known terrestrial decay rate (from the lifetime), namely,  $\log \beta^- =$ 1.2536. The rate for the ground state to state 3 transition was adjusted to  $\log ft = 5.9$ , while the rate for the ground state to state 5 was adjusted to  $\log ft = 5.3$ . When the stellar rates were recalculated, the lowest temperaturedensity point result became  $\log \beta^- = -1.261$ , in agreement with the terrestrial rate to better than 2%. This type of procedure was repeated whenever the calculated rate differed from the terrestrial value by 30% or more. This procedure guaranteed that the input data used in the calculations described in this paper are consistent with measurements of terrestrial decay rates.

Table 3 presents the results of the stellar weak rate calculations for free nucleons and the nuclei in Table 1 on an abbreviated temperature and density grid (Table 2) for log  $(\rho/\mu_e) = 3,7,11$  and for  $T_9 = 1,3,10,100$ . The notation and format of these rate tables are similar to those in F<sup>2</sup>NI, but the entries presenting the temperature and density are replaced with a single entry (1 to 12) describing a temperature-density combination listed in Table 2.

The remaining columns in Table 3 are the logarithm of the positron emission rate,  $\log \beta^+$ , the logarithm of the continuum electron capture rate,  $\log \varepsilon^{-}$ , the logarithm of the total v-energy loss rate (in MeV s<sup>-1</sup>),  $\log v$ , the logarithm of the electron emission rate,  $\log \beta^-$ , the logarithm of the continuum positron capture rate, log  $\varepsilon^+$ . and finally the logarithm of the total  $\bar{\nu}$  energy loss rate (in MeV s<sup>-1</sup>), log  $\bar{\nu}$ . It is again emphasized that this rate table is reproduced on a wide mesh temperature and density grid designed only to provide active investigators with a readily accessible means for estimating stellar nuclear weak rates. For accurate calculations involving stellar evolution or nucleosynthesis problems readers are urged to write MJN and request the stellar weak rate magnetic tape.

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Note added in proof.—We have neglected changes in  $Q_n$  due to Coulomb plasma effects. In the plasma a nucleus has its binding energy increased due to interactions with the dense electron gas. In the standard strong-screening, Wigner-Seitz approximation this extra binding is given by (see Salpeter and Van Horn 1969)

$$Q_{\text{Coul}} = (1.764 \times 10^{-5}) Z^{5/3} (\rho/\mu_e)^{1/3} \text{MeV},$$

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No. 3, 1982

where Z is the nuclear charge and  $\rho/\mu_e(\text{in g cm}^{-3})$  is defined in the text. Because of the charge dependence of this binding, the effective nuclear Q-value,  $Q_n$ , changes at high density. In general, the electron capture Q-value will increase by

$$\Delta Q_n \approx (2.940 \times 10^{-5}) Z^{2/3} (\rho/\mu_e)^{1/3} \text{ MeV},$$

where Z is the charge of the parent nucleus. This is to be compared with the electron Fermi energy (for  $\rho/\mu_e \gtrsim 10^8$  g cm<sup>-3</sup>)

$$W_{\rm F} \approx 5.155 \times 10^{-3} (\rho/\mu_e)^{1/3} \, {\rm MeV}.$$

The ratio of  $\Delta Q_n$  to the electron Fermi energy  $W_F$  is independent of density in the strong-screening approximation and is given by

$$\frac{\Delta Q_n}{W_{\rm F}} \approx 5.703 \times 10^{-3} Z^{2/3} \,\,\mathrm{MeV}.$$

This ratio is at most 0.06 for the largest nuclear charge considered in this paper, Z = 30. We conclude that Coulomb plasma changes in the nuclear masses represent negligible effects in the nuclear weak interaction rates in comparison to other uncertainties in the problem. These uncertainties include the total strength and excitation energy of the Gamow-Teller resonances which dominate the rates at high  $W_F$  where  $\Delta Q_n$  is largest and the low-lying discrete Gamow-Teller strength distributions which dominate the rates for low  $W_F$  where  $\Delta Q_n$  is smallest.

We have also neglected the change in energy of highly excited nuclear states relative to the ground state due to Coulomb plasma effects. At low densities where the experimentally determined discrete states dominate the reaction rates these effects are small. At high densities the Gamow-Teller resonances dominate the reaction rates. Because of the collective nature of these resonances, the Coulomb plasma effects will be small.

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