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The SNO Collaboration

An Independent Measurement of the Total Active ^{8}B Solar Neutrino Flux Using an Array of ^{3}He Proportional Counters at the Sudbury Neutrino Observatory

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The Sudbury Neutrino Observatory (SNO) used an array of ^{3}He proportional counters to measure the rate of neutral-current interactions in heavy water and precisely determined the total active (ν_x) ^{8}B solar neutrino flux. This technique is independent of previous methods employed by SNO. The total flux is found to be $5.54^{+0.33}_{-0.31}$ (stat)$^{+0.36}_{-0.34}$ (syst) $\times 10^6$ cm^{-2}s^{-1}, in agreement with previous measurements and standard solar models. A global analysis of solar and reactor neutrino results yields $\Delta m^2 = 7.94^{+0.42}_{-0.26} \times 10^{-5}$ eV^2 and $\theta = 33.8^{+1.4}_{-1.3}$ degrees. The uncertainty on the mixing angle has been reduced from SNO's previous results.

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The Sudbury Neutrino Observatory [1] detects ^{8}B solar neutrinos through three reactions: charged-current interactions (CC) on deuterons, in which only electron neutrinos participate; neutrino-electron elastic scattering (ES), which are dominated by contributions from electron-neutrinos; and neutral-current (NC) disintegration of the deuteron by neutrinos, which has equal sensitivity to all active neutrino flavors.

In its first phase of operation, SNO measured the NC rate by observing neutron captures on deuterons and found that a Standard-Electro-Weak-Model description with an undis-

torted ^{8}B neutrino spectrum and CC, NC, and ES rates due solely to ν_e interactions was rejected at 5.3σ [2, 3, 4, 5]. The second phase of SNO measured the rates and spectra after the addition of ~2 Mg of NaCl to the ~1000 Mg of heavy water (D$_2$O). This enhanced the neutron detection efficiency and the ability to statistically separate the NC and CC signals, and resulted in significant improvement in the accuracy of the measured ν_e and ν_x fluxes without any assumption about the energy dependence of the neutrino flavor transformation [6, 7]. In the present measurement, the NC signal neutrons were predominantly detected by an array of ^{3}He proportional counters (Neutral Current Detection, or NCD, array [8]) consisting of 36 "strings" of counters that were deployed in the D$_2$O. Four additional strings filled with ^{4}He were insensitive to the neutron signals and were used to study backgrounds. Cherenkov light signals from CC, NC, and ES reactions were still recorded by the photomultiplier tube (PMT) array, though the rate of such NC events from ^{2}H(n,γ)^{3}H reactions was significantly suppressed due to neutron absorption in the ^{3}He strings. This Letter presents measurements of the CC, NC, and ES rates from SNO's NCD phase.

The data presented here were recorded between November 27, 2004 and November 28, 2006, totaling 385.17 live days. The number of raw NCD triggers was 1,417,811 and the data set was reduced to 91,636 NCD events after data reduction described in [8]. Six strings filled with ^{3}He were excluded from the analysis due to various defects. The number of raw PMT triggers was 146,431,347 with 2381 PMT events passing data reduction and analysis selection requirements similar to those in [5]. Background events arising from β–γ decays were reduced by selecting events with reconstructed electron effective kinetic energies $\geq$ 6.0 MeV and reconstructed vertices within $R_{\mathrm{fit}} \leq 550$ cm.

Thermal neutron capture on the ^{3}He in the proportional counters results in the creation of a proton-triton pair with a total kinetic energy of 764 keV. Due to particles hitting the counter walls [8], the detected ionization energy was between 191 and 764 keV. The signals from each string were amplified logarithmically to provide sufficient dynamic range before they were digitized [8]. The detectors were constructed from ultra-pure nickel produced by a chemical vapor deposition process to minimize internal radioactivity.

The neutron detection efficiency and response of the PMT and NCD arrays have been determined with a variety of neutron calibration sources. Neutron point sources (^{252}Cf and ^{241}AmBe) were frequently deployed throughout the detector volume to measure the temporal stability and the detector gain of the NCD array. The NC neutron detection efficiency was studied by using an isotropic source of neutrons produced by mixing ^{24}Na (t$_{1/2}$ = 14.959 hours), in the form of activated NaCl, into the heavy water in October 2005 and October 2006. Neutrons were produced by deuterium photodisintegration induced by the 2.754-MeV ^{24}Na gammas. The largest uncertainties on the neutron detection efficiency were associated with the knowledge of the ^{24}Na source strength and the ability to determine the uniformity of its mixing in the
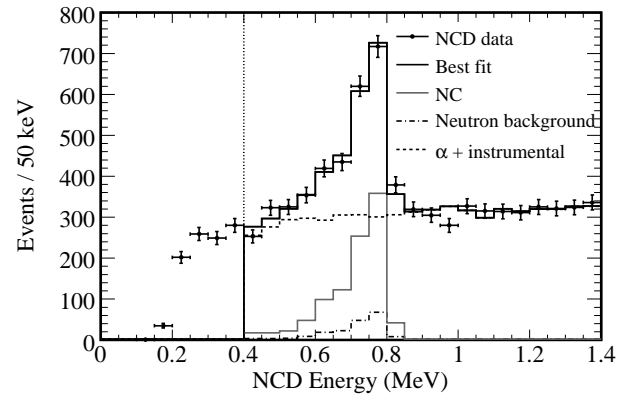

FIG. 1: NCD energy spectrum fit with a neutron calibration spectrum, neutron backgrounds, alpha background derived from Monte Carlo simulation, and low-energy instrumental background distributions. Data are shown after data reduction up to 1.4 MeV and the fit is above 0.4 MeV.

heavy water. The inferred NC neutron capture efficiency for the NCD array was 0.211±0.007 in good agreement with the 0.210 ± 0.003 given by a Monte Carlo simulation verified against point-source data. The fraction of detected neutrons inside the analysis energy range from 0.4–1.4 MeV, including the effects of data reduction, electronic thresholds and efficiency, and digitizer livetime, was 0.862±0.004. The neutron detection efficiency for the PMT array was 0.0485±0.0006.

The energy spectrum of the reduced NCD data set is shown in Fig. 1. The distinctive neutron spectrum peaks at 764 keV. This spectrum was fit with a neutron energy spectrum taken from the ^{24}Na calibration. The alpha background distribution was derived from a Monte Carlo simulation of the proportional counters. The alpha background energy spectrum has several components, U and Th progeny in the bulk of the nickel detector bodies and ^{210}Po on the inner surfaces [8]. These sources resulted in approximately 16 alphas per day detected in the full neutron energy window for the entire NCD array. The Monte Carlo simulation was verified using alpha data from the array above 1.2 MeV and from the ^{4}He strings in the neutron energy region. Several uncertainties were included in the alpha background distribution: depth profile and composition of alpha emitters in nickel, electron drift time, space-charge model parameters, and ion mobility. Low-energy instrumental background events were found on two strings that were excluded from the analysis. Distributions of these events were used to fit for possible additional contamination in the data on the rest of the array.

The optical and energy responses and position and directional reconstruction of the PMT array were updated to include light shadowing and reflection from the NCD array. With improvements to the calibration data analysis and increased high voltage on the PMT array, the introduction of the NCD array did not significantly increase the position or energy reconstruction uncertainties from previous phases. A normalization for the photon detection efficiency based on ^{16}N

calibration data [9] and Monte Carlo simulations was used to set the absolute energy scale. The energy response for electrons can be characterized by a Gaussian function with resolution $\sigma_T = -0.2955 + 0.5031\sqrt{T_e} + 0.0228\,T_e$, where T_e is the electron kinetic energy in MeV. The energy scale uncertainty was 1.1%.

Backgrounds are summarized in Table I. Low levels of ^{214}Bi and ^{208}Tl present in the heavy and light water, NCD counters, and cables can create free neutrons from deuteron photodisintegration and low-energy Cherenkov events from $\beta - \gamma$ decays. Techniques to determine these backgrounds in the water are described in previous works [3, 10, 11, 12]. In addition, alphas from Ra progeny on the NCD tube surfaces can induce ^{17,18}O(α,n) interactions. The background contributions from the NCD array were determined by combining the analyses of the alpha energy spectrum and the time-correlated alpha events in the decay chains. The results from these studies agreed with those from radioassays of the materials prior to the construction of the NCD array. In addition, *in-situ* analysis of the Cherenkov light found three detectable "hotspots" of elevated radioactivity on two strings. Evaluations of their isotopic composition were made upon removal of the NCD array after the end of data taking. Results from the *in-situ* and *ex-situ* methods showed the neutron background uncertainty from the hotspots to be less than 0.7% of the NC signal. Neutron backgrounds from atmospheric neutrino interactions and ^{238}U fission were estimated with NUANCE [13] and from event multiplicities.

Previous results [6, 7] reported the presence of external-source neutrons from the acrylic vessel and light water. Alpha radioactivity measurements of the acrylic vessel neck with Si counters before and after the NCD phase showed values consistent with those from the previous phase within uncertainties. Thus, the external-source neutron contribution from the vessel was taken to be the same as for the previous phase. The contributions from the light water were determined from the measured ^{214}Bi and ^{208}Tl concentrations.

Backgrounds from Cherenkov events inside and outside the fiducial volume were estimated using calibration source data, measured activities, Monte Carlo calculations, and controlled injections of Rn into the detector. These backgrounds were found to be small above the analysis energy threshold and within the fiducial volume, and were included as an additional uncertainty on the flux measurements. Previous phases identified isotropic acrylic vessel background (IAVB) events, which can be limited to 0.3 remaining IAVB events (68% CL) after data reduction for this phase.

A blind analysis procedure was used to minimize the possibility of introducing biases. The data set used during the development of the analysis procedures excluded a hidden fraction of the final data set and included an admixture of neutron events from muon interactions. The blindness constraints were removed after all analysis procedures, parameters, and backgrounds were finalized. A simultaneous fit was made for the number of NC events detected by the NCDs, the numbers of NC, CC and ES events detected by the

TABLE I: Background events for the PMT and NCD arrays, respectively. The internal and external neutrons and the γ-ray backgrounds are constrained in the analysis. "Other backgrounds" include terrestrial $\bar{\nu}$s, reactor $\bar{\nu}$s, spontaneous fission, cosmogenics, CNO νs, and (α,n) reactions. The last two entries are included in the systematic uncertainty estimates for the PMT array.

Source	PMT Events	NCD Events
D_2O radioactivity	7.6 ± 1.2	28.7 ± 4.7
NCD bulk/^{17}O,^{18}O	$4.6^{+2.1}_{-1.6}$	$27.6^{+12.9}_{-10.3}$
Atmospheric ν/^{16}N	24.7 ± 4.6	13.6 ± 2.7
"Other backgrounds"	0.7 ± 0.1	2.3 ± 0.3
NCD "hotspots"	17.7 ± 1.8	64.4 ± 6.4
NCD cables	1.1 ± 1.0	8.0 ± 5.2
Total internal neutron background	$56.4^{+5.6}_{-5.4}$	$144.6^{+13.8}_{-14.8}$
External-source neutrons	20.6 ± 10.4	40.9 ± 20.6
Cherenkov events from β–γ decays	$5.8^{+9.7}_{-2.9}$	–
IAVB	< 0.3 (68% CL)	–

PMTs, as well as the numbers of background events of various types. A Markov-chain Monte Carlo (MCMC) method with the Metropolis-Hastings algorithm [14, 15] was employed to make the fit, which also allowed nuisance parameters (systematics) weighted by external constraints to vary in the fit. The NCD event energy spectrum was fit with an alpha background distribution, a neutron calibration spectrum, expected neutron backgrounds, and two instrumental background event distributions. The PMT events were fit in reconstructed energy, the cosine of the event direction relative to the vector from the sun ($\cos\theta_\odot$), and the reconstructed radial position.

The spectral distributions of the ES and CC events were not constrained to the ^{8}B shape, but were extracted from the data. Fits to the data yielded the following number of events: 983^{+77}_{-76} NC (NCD array), 267^{+24}_{-22} NC (PMT array), 1867^{+91}_{-101} CC, and 171^{+24}_{-22} ES, with 185^{+25}_{-22} and 77^{+12}_{-10} neutron background events in the NCD and PMT arrays respectively. Additionally, the total NCD array background fits including alphas and the two instrumental components yielded 6127 ± 101 events.

Assuming the ^{8}B neutrino spectrum from [17], the equivalent neutrino fluxes derived from the fitted CC, ES, and NC events are (in units of 10^6 cm^{-2}s^{-1}) [18, 19]:

$$\phi_{CC}^{SNO} = 1.67^{+0.05}_{-0.04}(\text{stat})^{+0.07}_{-0.08}(\text{syst})$$
$$\phi_{ES}^{SNO} = 1.77^{+0.24}_{-0.21}(\text{stat})^{+0.09}_{-0.10}(\text{syst})$$
$$\phi_{NC}^{SNO} = 5.54^{+0.33}_{-0.31}(\text{stat})^{+0.36}_{-0.34}(\text{syst}),$$

and the ratio of the ^{8}B neutrino flux measured with the CC and NC reactions is

$$\frac{\phi_{CC}^{SNO}}{\phi_{NC}^{SNO}} = 0.301 \pm 0.033\,(\text{total}).$$

The contributions to the systematic uncertainties on the derived fluxes are shown in Table II.

Two independent analysis methods were used as checks of the MCMC method. Both used maximum likelihood fits but handled the systematics differently. A comparison of results

TABLE II: Sources of systematic uncertainties on NC, CC, and ES fluxes. The total error differs from the individual errors added in quadrature due to correlations.

Nuisance Parameter	NC uncert. (%)	CC uncert. (%)	ES uncert. (%)
PMT energy scale	±0.6	±2.7	±3.6
PMT energy resolution	±0.1	±0.1	±0.3
PMT radial scaling	±0.1	±2.7	±2.7
PMT angular resolution	±0.0	±0.2	±2.2
PMT radial energy dep.	±0.0	±0.9	±0.9
Background neutrons	±2.3	±0.6	±0.7
Neutron capture	±3.3	±0.4	±0.5
Cherenkov/AV backgrounds	±0.0	±0.3	±0.3
NCD instrumentals	±1.6	±0.2	±0.2
NCD energy scale	±0.5	±0.1	±0.1
NCD energy resolution	±2.7	±0.3	±0.3
NCD alpha systematics	±2.7	±0.3	±0.4
PMT data cleaning	±0.0	±0.3	±0.3
Total experimental uncertainty	±6.5	±4.0	±4.9
Cross section [16]	±1.1	±1.2	±0.5

from these three analysis methods after the blindness conditions had been removed revealed two issues. A 10% difference between the NC flux uncertainties was found, and subsequent investigation revealed incorrect input parameters in two methods. After the inputs were corrected, the errors agreed and there was no change in the fitted central values. However, the ES flux from the MCMC fit was 0.5σ lower than from the other two analyses. This difference was found to be from the use of an inappropriate algorithm to fit the peak of the ES posterior distributions. After a better algorithm was implemented, the ES flux agreed with the results from the other two analyses.

The ES flux presented here is 2.2σ lower than that found by Super-Kamiokande-I [20] consistent with a downward statistical fluctuation in the ES signal, as evidenced in the shortfall of signals near $\cos\theta_\odot = 1$ in two isolated energy bins. The ^{8}B spectral shape [17] used here differs from that [21] used in previous SNO results. The CC, ES and NC flux results in this Letter are in agreement ($p = 32.8\%$ [22]) with the NC flux result of the D_2O phase [3] and with the fluxes from the salt phase [7].

The fluxes presented here, combined with day and night energy spectra from the pure D_2O and salt phases [4, 7], place constraints on neutrino flavor mixing parameters. Two-flavor active neutrino oscillation models are used to predict the CC, NC, and ES rates in SNO [23]. A combined χ^2 fit to SNO D_2O, salt, and NCD-phase data [24] yields the allowed regions in Δm^2 and $\tan^2\theta$ shown in Fig. 2(a). In a global analysis of all solar neutrino data (including Borexino [25] and Super-Kamiokande-I [20]) and the 766 ton-year KamLAND reactor anti-neutrino results [26], the allowed regions are shown in Fig. 2(b and c). The best-fit point to the solar global plus KamLAND data yields $\Delta m^2 = 7.94^{+0.42}_{-0.26} \times 10^{-5}$ eV^2 and $\theta = 33.8^{+1.4}_{-1.3}$ degrees, where the errors reflect marginalized 1-σ ranges. In our analyses, the ratio f_B of the total ^{8}B flux to
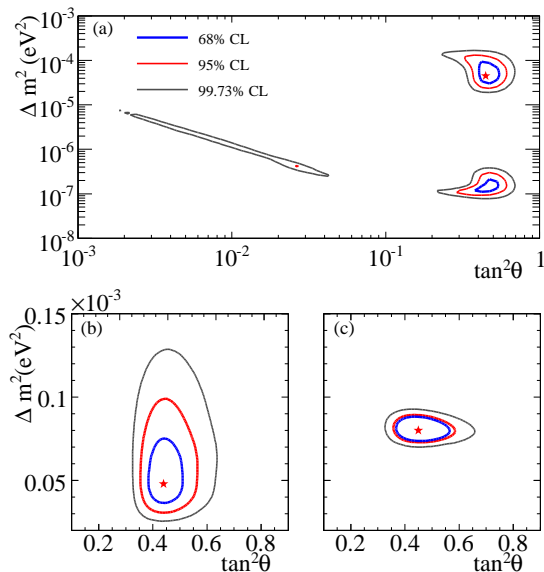

FIG. 2: Neutrino-oscillation contours. (a) SNO only: D_2O & salt day and night spectra, NCD phase fluxes. The best-fit point is $\Delta m^2 = 4.57\times10^{-5}$ eV^2, $\tan^2\theta = 0.447$, $f_B = 0.900$, with χ^2/d.o.f.=73.77/72. (b) Solar Global: SNO, SK, Cl, Ga, Borexino. (c) Solar Global + KamLAND. The best-fit point is $\Delta m^2 = 7.94 \times 10^{-5}$ eV^2, $\tan^2\theta = 0.447$, $f_B = 0.873$.

the SSM [27] value was a free parameter, while the total *hep* flux was fixed at 7.93×10^3 cm^{-2} s^{-1} [28].

In summary, we have precisely measured the total flux of active ^{8}B neutrinos from the sun independently from our previous methods. The flux is in agreement with standard solar model calculations. This Letter presents analysis leading to a reduction in the uncertainty of θ over our previous results. Further improvements are expected with the addition of the latest data from other experiments [29, 30].

[1] SNO Collaboration, Nucl. Instr. and Meth. **A449**, 172 (2000).
[2] SNO Collaboration, Phys. Rev. Lett. **87**, 071301 (2001).
[3] SNO Collaboration, Phys. Rev. Lett. **89**, 011301 (2002).
[4] SNO Collaboration, Phys. Rev. Lett. **89**, 011302 (2002).
[5] SNO Collaboration, Phys. Rev. **C75**, 045502 (2007).
[6] SNO Collaboration, Phys. Rev. Lett. **92**, 0181301 (2004).
[7] SNO Collaboration, Phys. Rev. **C72**, 055502 (2005).
[8] J.F. Amsbaugh *et al.*, Nucl. Instr. and Meth. **A579**, 1054 (2007).
[9] M.R. Dragowsky *et al.*, Nucl. Instr. and Meth. **A481**, 284 (2002).
[10] B. Aharmim *et al.*, arXiv:0803.4162v1 [nucl-ex], submitted to

Nucl. Instr. and Meth.

[11] T.C. Andersen *et al.*, Nucl. Instr. and Meth. **A501**, 399 (2003).

[12] I. Blevis *et al.*, Nucl. Instr. and Meth. **A517**, 139 (2004).

[13] D. Casper, Nucl. Phys. Proc. Suppl. **112**, 161 (2002).

[14] N. Metropolis *et al.*, Journ. Chem. Phys. **21**, 1087 (1953).

[15] W.K. Hastings, Biometrika **57**, 97 (1970).

[16] Cross section uncertainty includes: g_A (0.5%), theoretical cross section (1% for CC and NC, 0.5% for CC/NC; S. Nakumura *et al.* Nucl. Phys. **A707**, 561 (2002)) and radiative corrections (0.3% for CC, 0.1% for NC; A. Kurylov, M.J. Ramsey-Musolf, and P. Vogel, Phys. Rev. **C65** 05501 (2002)).

[17] W.T. Winter *et al.*, Phys. Rev. **C73**, 025503 (2006).

[18] *hep* neutrinos could also be present in the measured fluxes; the SSM contribution would be 0.5%.

[19] Electron-neutrino cross sections are used to calculate all fluxes. The ES cross section is from J.N. Bahcall *et al.*, Phys. Rev. **D51**, 6146 (1995).

[20] J. Hosaka *et al.*, Phys. Rev. **D73**, 112001 (2006).

[21] C.E. Ortiz *et al.*, Phys. Rev. Lett. **85**, 2909 (2000).

[22] A. Valassi, Nucl. Instr. and Meth. **A500**, 391 (2003).

[23] Z. Maki, N. Nakagawa, and S. Sakata, Prog. Theor. Phys. **28** 870 (1962); V. Gribov and B. Pontecorvo, Phys. Lett. **B28** 493 (1969); S. P. Mikheyev and A. Yu. Smirnov, Sov. J. Nucl. Phys. **42** 913 (1985); L. Wolfenstein, Phys. Rev. **D17** 2369 (1978).

[24] See the SNO website http://sno.phy.queensu.ca for details.

[25] C. Arpesella *et al.*, Phys. Letters **B658**, 101 (2008).

[26] T. Araki *et al.*, Phys. Rev. Lett. **94**, 081801 (2005).

[27] J. N. Bahcall, A. M. Serenelli, and S. Basu, Astrophys. J. **621**, L85 (2005).

[28] J. N. Bahcall, A. M. Serenelli, and S. Basu, Ap. J. Supp. **165**, 400 (2006).

[29] S. Abe *et al.*, Phys. Rev. Lett. **100**, 221803 (2008).

[30] J.P. Cravens *et al.*, arXiv:0803.4312v1 [hep-ex], submitted to Phys. Rev. D.